

Molecular gas and star formation within 12 strong galactic bars observed with IRAM-30 m[★]

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ABSTRACT

Context. While some galactic bars show recent massive star formation (SF) along them, some others do not. Whether bars with low level of SF are a consequence of low star formation efficiency, low gas inflow rate, or dynamical effects remains a matter of debate.

Aims. In order to study the physical conditions that enable or prevent SF, we perform a multi-wavelength analysis of 12 strongly barred galaxies with total stellar masses $\log_{10}(M_*/M_\odot) \in [10.2, 11]$, chosen to host different degrees of SF along the bar major axis without any prior condition on gas content. We observe the CO(1–0) and CO(2–1) emission within bars with the IRAM-30 m telescope (beam sizes of 1.7–3.9 kpc and 0.9–2.0 kpc, respectively; 7–8 pointings per galaxy on average).

Methods. We estimated molecular gas masses (M_{mol}) from the CO(1–0) and CO(2–1) emissions. SF rates (SFRs) were calculated from GALEX near-ultraviolet (UV) and WISE 12 μm images within the beam-pointings, covering the full bar extent (SFRs were also derived from far-UV and 22 μm).

Results. We detect molecular gas along the bars of all probed galaxies. Molecular gas and SFR surface densities span the ranges $\log_{10}(\Sigma_{\text{mol}}/[M_\odot \text{ pc}^{-2}]) \in [0.4, 2.4]$ and $\log_{10}(\Sigma_{\text{SFR}}/[M_\odot \text{ pc}^{-1} \text{ kpc}^{-2}]) \in [-3.25, -0.75]$, respectively. The star formation efficiency (SFE; i.e., $\text{SFR}/M_{\text{mol}}$) in bars varies between galaxies by up to an order of magnitude ($\text{SFE} \in [0.1, 1.8] \text{ Gyr}^{-1}$). On average, SFEs are roughly constant along bars. SFEs are not significantly different from the mean value in spiral galaxies reported in the literature ($\sim 0.43 \text{ Gyr}^{-1}$), regardless of whether we estimate M_{mol} from CO(1–0) or CO(2–1). Interestingly, the higher the total stellar mass of the host galaxy, the lower the SFE within their bars. In particular, the two galaxies in our sample with the lowest SFE and Σ_{SFR} (NGC 4548 and NGC 5850, $\text{SFE} \lesssim 0.25 \text{ Gyr}^{-1}$, $\Sigma_{\text{SFR}} \lesssim 10^{-2.25} M_\odot \text{ yr}^{-1} \text{ kpc}^{-2}$, $M_* \gtrsim 10^{10.7} M_\odot$) are also those hosting massive bulges and signs of past interactions with nearby companions.

Conclusions. We present a statistical analysis of the SFE in bars for a sample of 12 galaxies. The SFE in strong bars is not systematically inhibited (either in the central, middle, or end parts of the bar). Both environmental and internal quenching are likely responsible for the lowest SFEs reported in this work.

Key words. galaxies: star formation – galaxies: ISM – galaxies: structure – galaxies: evolution – galaxies: statistics

1. Introduction

Stellar bars are frequent stellar structures (e.g., [de Vaucouleurs 1963](#); [Knapen et al. 2000](#); [Eskridge et al. 2000](#); [Aguerri et al. 2009](#); [Méndez-Abreu et al. 2012](#); [Díaz-García et al. 2016b](#)).

The bar potential creates lanes of molecular gas and dust in disk galaxies (e.g., [de Vaucouleurs & de Vaucouleurs 1963](#); [Combes et al. 1990](#); [Athanassoula 1992](#); [Friedli & Benz](#)

1993; [Sellwood & Wilkinson 1993](#); [Knapen et al. 1995, 2002](#); [Kormendy & Kennicutt 2004](#); [Pérez et al. 2004](#); [Comerón et al. 2009](#); [Sánchez-Menguiano et al. 2015](#)). These lanes are the loci of shocks in the gas flow ([Prendergast et al. 1962](#)) that can induce star formation (SF).

Different distributions of SF in bars have been found based on the detection of ionised gas, which is typically traced by H α or ultraviolet (UV) emission (e.g., [Fraser-McKelvie et al. 2020](#); [Díaz-García et al. 2020](#), and references therein). H II regions are detected all along the bar in some galaxies, whereas SF is only detected in the nuclear or circumnuclear regions in others (e.g., [Martin & Friedli 1997](#); [Sheth et al. 2002](#); [Verley et al. 2007b](#)).

* Tabulated CO(1–0) and CO(2–1) spectra are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/654/A135>

With the aid of stacking techniques applied to far-UV images, Díaz-García et al. (2020) showed that massive barred galaxies (total stellar masses $M_* \gtrsim 10^{10} M_\odot$) are typically characterised by a dip in the radial distribution of SF that is not seen in non-barred systems (see also James et al. 2009). These ‘SF deserts’ indeed exist in barred galaxies (James & Percival 2016, 2018; Spinoso et al. 2017; Donohoe-Keyes et al. 2019).

Bar-driven central starbursts have also been proposed as the mechanism that depletes the gas, and thus induces SF quenching (e.g., Cheung et al. 2013; Gavazzi et al. 2015; Newnham et al. 2020). Specifically, simulations by Khoperskov et al. (2018) of gas-rich disk isolated galaxies show that stellar bars reduce the star formation rates (SFRs) by a factor of ten in less than 1 Gyr, right after the bar strength reaches its saturation level. This supports the important role of the Galactic bar on the star formation history of the Milky Way (e.g., Haywood et al. 2016).

What drives SF in bars? In this work our aim is to compare the molecular gas mass (M_{mol}) with the SFR along galactic bars, quantifying star formation efficiencies (SFE; i.e., $\text{SFR}/M_{\text{mol}}$). This comparison allows us to distinguish whether the lack of SF in some bars is due to a small amount of gas (most likely caused by a low inflow rate) or rather due to a low star formation efficiency.

Here we estimated SFRs based on both Galaxy Evolution Explorer (GALEX) UV imaging and Wide-field Infrared Survey Explorer (WISE) data. The UV radiation from young massive stars is partially absorbed by interstellar dust, which gets heated and emits infrared (IR) photons (e.g., Lagache et al. 2005). The mid-IR emission in galaxies mainly comes from dust continuum associated with very small grains (VSGs) and polycyclic aromatic hydrocarbons (PAHs) (e.g., Leger & Puget 1984; Puget & Leger 1989), and can be used as a robust SFR indicator (e.g., Zhu et al. 2008). There is a strong correlation between the $8 \mu\text{m}$ and $24 \mu\text{m}$ emission of H II regions and their H α and UV counterparts (e.g., Calzetti et al. 2004). Not only can UV-based SFRs be underestimated, but also IR-based SFRs, since not all the starlight is re-processed by dust (Kennicutt et al. 2009); a hybrid combination of different tracers is advocated (e.g., Leroy et al. 2019).

Giant molecular clouds (GMCs), from which stars are formed, are mainly composed of H₂, dust, and other molecules such as carbon monoxide (CO; Solomon et al. 1987; Heyer et al. 2009; Narayanan et al. 2011). GMCs can be probed from the CO emission (Dickman et al. 1986). Using the IRAM-30 m radiotelescope (Sierra Nevada, Spain), we observed the $J = 1-0$ and $J = 2-1$ rotational transitions of CO along the strong galactic bars of 12 nearby massive spiral galaxies ($M_* \gtrsim 10^{10.2} M_\odot$). We chose bars with different degrees of SF, hosting and lacking H II regions, as traced from ancillary H α images, and with no prior information about the molecular gas content.

Since the pioneering work of Schmidt (1959), it has been a highly important endeavour to find a universal power law that relates the SFR surface density (Σ_{SFR}) and the molecular gas surface density (Σ_{mol} , e.g., Young et al. 1996; Kennicutt 1998b; Kennicutt et al. 2007; Verley et al. 2010). The molecular gas mass within GMCs is indeed correlated with the SF activity in galaxies (e.g., Leroy et al. 2008). Bigiel et al. (2011) analysed the $\Sigma_{\text{mol}} - \Sigma_{\text{SFR}}$ relation by exploiting CO data from the HERA CO Line Extragalactic Survey (HERACLES; Leroy et al. 2009), GALEX far-UV, and Spitzer 24 μm imaging, over thousands of positions in 30 nearby disks at a resolution of 1 kpc. The authors imply a median molecular gas depletion time ($\tau = 1/\text{SFE} = M_{\text{mol}}/\text{SFR}$; including helium) of ~ 2.35 Gyr ($\sigma = 0.24$ dex), that is, a median SFE of 0.43 Gyr^{-1} (see also Bigiel et al. 2008);

these values are used as a benchmark in this work. Our aim is to assess the relationship between Σ_{SFR} and Σ_{mol} , known as the Kennicutt-Schmidt (KS) law, within stellar bars. The analysis of the KS law in the central regions of gas-rich simulated galaxies by Khoperskov et al. (2018) reveals a flattening of its slope in the bar region.

In general, SFEs in bars have been only studied in a few objects (e.g., Knapen & Beckman 1996; Muraoka et al. 2016; Yajima et al. 2019). It has been shown that SF can be suppressed even in the presence of molecular gas (Maeda et al. 2020, and references therein). The SFE in the spiral arms has been shown to be higher than in the bars (e.g., Handa et al. 1991; Momose et al. 2010; Hirota et al. 2014, and references therein). In this work we search for possible gradients in SFEs along bars.

Some studies using CO indicate that the onset of SF along bars depends on the strength of shocks and shear stress (e.g., Reynaud & Downes 1998; Zurita et al. 2004). The formation of GMCs can also be inhibited by diverging streamlines in bars (e.g., Regan & Elmegreen 1997; Sheth et al. 2002). However, observations of H II regions along bars have not always confirmed these hypotheses (e.g., Martin & Friedli 1997; Zurita & Pérez 2008). Thus, several physical processes might come into play to explain the occurrence of SF in bars; for instance, Sheth et al. (2000) suggest that the disruption of GMCs in bars can be neutralised by self-gravity. Disentangling the physical conditions that enable or prevent SF in bars remains an unsolved astrophysical problem, despite all the observational and theoretical work referred to above.

This paper is organised as follows. In Sect. 2 we present the sample and the set of UV, optical, near- and mid-infrared, and millimetre data used. For the millimetre data we explain the observing strategy, data acquisition, and reduction of CO spectra; the spectra are shown in Appendix A. In Sect. 3 we describe the methodology used to infer molecular gas masses from CO(1–0) spectra (see Appendix B); Σ_{mol} profiles derived from CO(2–1) spectra are analysed in Appendix E. Section 4 describes the use of aperture photometry to compute SFRs within the same regions covered by the IRAM-30 m pointings (see also Appendix C); the resulting values are tabulated in Appendix D. In Sect. 5 we analyse the distribution of Σ_{SFR} and Σ_{mol} along the stellar bars. In Sect. 6 we discuss the physical properties that determine the presence or absence of star formation in galactic bars (see also Appendix F). Finally, Sect. 7 summarises the most important results of this work.

2. Sample and data

Our sample of 12 galaxies was selected from the extragalactic database HyperLeda¹ (Paturel et al. 2003) and met the following criteria:

- They host prominent bars, according to the visual classification from ancillary optical images;
- They probe both star-forming and quiescent bars, based on the perusal of archival H α or mid-IR imaging;
- They have morphological types spanning the range S0/a-Sc (i.e., Hubble types $T \in [0, 5]$);
- They are nearby, with recessional velocities $< 3000 \text{ km s}^{-1}$;
- They have low disk inclinations ($i < 50^\circ$), enabling the study of surface densities (of SF and gas masses) and minimising the effect of dust absorption in optical and near-IR wavelengths;

¹ <http://leda.univ-lyon1.fr>

Table 1. Sample and properties.

Galaxy (1)	Morphological classification (2)	T (3)	D (Mpc) (4)	i (°) (5)	$\log_{10} M_\star$ (M_\odot) (6)	$\log_{10} M_{\text{HI}}$ (M_\odot) (7)	Σ_3^A (8)	D_{25} (‘) (9)	R_{bar} (‘) (10)	A_2 (11)	ϵ_{bar} (12)	N_{CO} (13)	θ_{beam} (kpc) (14)
NGC 3504	(R'L)SAB(rs, nl)a	1	20.0	12.8	10.40	8.68	0.70	147.3	33.97	1.09	0.63	7	2.1
NGC 4123	SB _x (rs)ab	2	21.9	46.9	10.29	9.64	0.54	189.7	52.51	0.62	0.68	10	2.3
NGC 4394	(RL)SB(rs, bl, nl)0/a	0	16.8	30.4	10.44	8.73	1.68	208.0	45.71	0.55	0.55	7	1.8
NGC 4535	SAB(s)c	5	17.0	25.6	10.69	9.63	0.78	487.7	43.08	0.52	0.65	6	1.8
NGC 4548	SB(rs, bl)ab	2	16.2	39.0	10.70	8.84	2.03	329.7	57.98	0.68	0.64	9	1.7
NGC 4593	(R')SB(rs, bl)a	1	33.9	33.4	10.93	9.31	1.34	143.9	52.51	0.81	0.62	7	3.6
NGC 5383	SB(rs, nr)ab	1	37.6	34.7	10.80	9.88	0.24	147.3	50.16	0.81	0.57	10	3.9
NGC 5850	(R')SB(r, bl, nr, nb)ab	2	23.1	36.0	10.77	9.27	1.44	198.7	62.76	0.74	0.65	11	2.4
NGC 5921	SB(rs)b	3	19.9	33.5	10.41	9.36	-0.26	181.2	45.55	0.40	0.73	7	2.1
NGC 6217	(R'L)SB(rs)b	3	23.9	23.5	10.45	9.75	-0.17	134.3	36.76	0.56	0.74	7	2.5
NGC 6951	SAB(rs)bc	4	23.5	39.4	10.90	9.51	-	189.7	44.01	-	0.54	7	2.5
NGC 7723	(R')SB(rs)ab	3	27.4	47.8	10.68	8.95	-0.47	189.7	19.13	0.38	0.58	5	2.9

Notes. Column 1: Galaxy identification. Column 2: Morphological classification. Column 3: Hubble type. Column 4: Distance. Column 5: Disk inclination. Column 6: Total stellar mass. Column 7: Total H I mass. Column 8: Projected surface density to the third nearest neighbour galaxy. Column 9: Projected major axis at the isophotal level 25 mag arcsec $^{-2}$ in the B -band. Column 10: Bar radius (calculated via ellipse fitting). Column 11: Bar normalised $m = 2$ Fourier amplitude. Column 12: Bar deprojected ellipticity. Column 13: Number of pointings done with IRAM-30 m. Column 14: CO(1–0) IRAM-30 m FWHM in kpc. For CO(2–1) the FWHM is one half of the listed value.

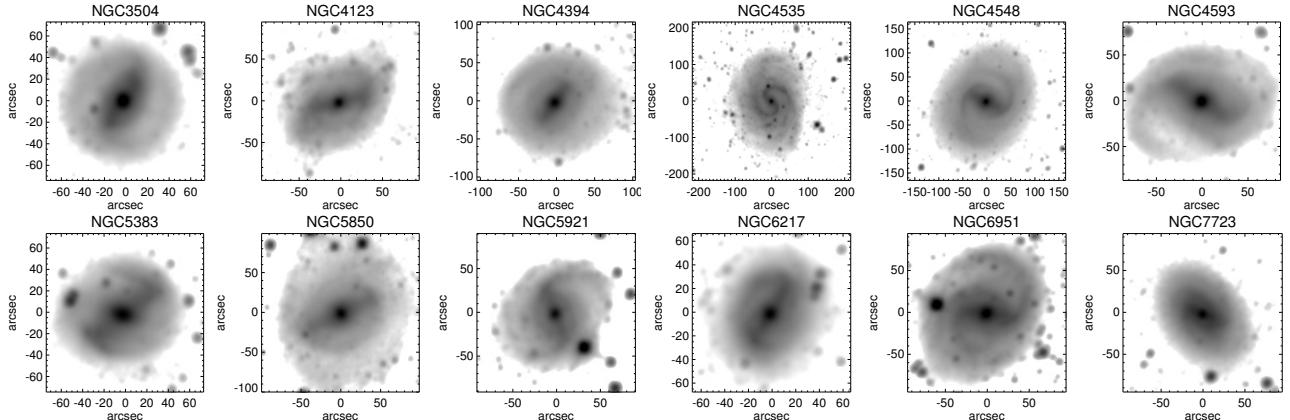


Fig. 1. Sky-subtracted WISE 1 images of the 12 galaxies in our sample, in magnitude scale with surface brightness ($\mu_{3.4\mu\text{m}}$, AB) in the range [17–25] mag arcsec $^{-2}$.

- They have galactic declinations $\delta > -15^\circ$, avoiding high air masses during IRAM-30 m observations;
- They have $D_{25} > 2'$, where D_{25} is the length of the projected major axis of a galaxy at the isophotal level 25 mag arcsec $^{-2}$ in the B -band, ensuring good spatial resolution.

The galaxy sample is listed in Table 1 together with some of the main galaxy properties. Their WISE 3.4 μm images are shown in Fig. 1. We use the mean of the redshift-independent distances (D) available in the NASA/IPAC Extragalactic Database (NED)². Total stellar masses are taken from Muñoz-Mateos et al. (2015), calculated from 3.6 μm imaging using the calibration of the mass-to-light ratio (M/L) by Eskew et al. (2012), which is also the M/L applied in this work to estimate stellar densities within bars (Sect. 4), and from Harmsen et al. (2017) in the case of NGC 6951 (using the K band and $M/L = 0.6$; Bell & de Jong 2001).

According to the catalogue of quasars and active galactic nuclei by Véron-Cetty & Véron (2010), NGC 4593 is classified as Seyfert 1.0, NGC 4548 is a Seyfert 3.0 or hosts a low-ionisation nuclear emission-line region (LINER), and NGC 6217

and NGC 3504 harbour nuclear H II regions. NGC 6951 has been reported to host an active galactic nucleus (AGN) (Ho et al. 1997; Kohno et al. 1999; Pérez et al. 2000). The rest of the galaxies in our sample are not active.

Díaz-García et al. (2020) classified the distribution of massive SF in bars in the *Spitzer* Survey of Stellar Structure in Galaxies (S⁴G; Sheth et al. 2010), using continuum-subtracted H α and far-UV images, based on the detection of H II knots, clumps, or filaments along the major axis. For NGC 6951 (not in S⁴G) we inspected the GALEX far-UV image. One-half of our galaxies (NGC 3504, NGC 4535, NGC 4593, NGC 5383, NGC 5921, NGC 7723) harbour H II regions along the bar major axis, whereas in the other galaxies these regions are scarce. We note that in the case of NGC 6217 the distribution of SF is asymmetric and only presents a clear SF gap in the north-eastern part of the bar.

In Table 1 we also show the morphological classification by Buta et al. (2015) (S⁴G) and/or NED (Col. 2). Eleven of the galaxies in our sample host inner features, the exception being NGC 4535. Of these, the majority have inner pseudo-rings (i.e., made of tightly wrapped spiral arms) and only NGC 5850 harbours a closed inner ring. This is relevant for our work

² <http://ned.ipac.caltech.edu>

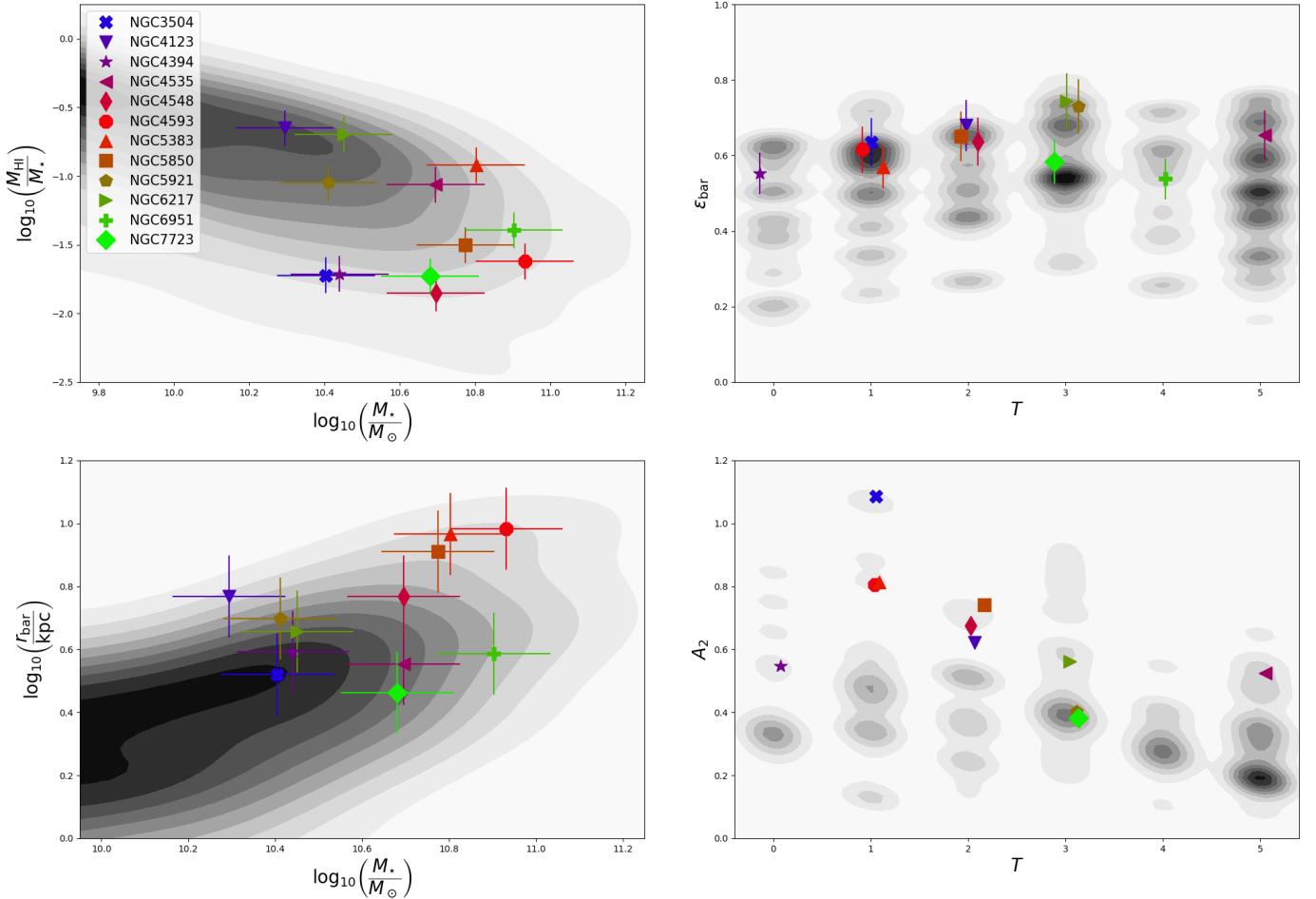


Fig. 2. Properties of our sample of 12 barred galaxies (coloured symbols; see legend) and comparison with 508 barred galaxies in the S⁴G survey (grey contours; disk inclinations lower than 50°), as reported in Díaz-García et al. (2016b) and Díaz-García et al. (2019a). *Upper left panel:* gas fraction as a function of the total stellar mass. Error bars indicate the error on M_{\star} associated with a 30% uncertainty on the mass-to-light ratio (Eskew et al. 2012). *Lower left panel:* bar length (in kiloparsecs) vs total stellar mass. Vertical error bars are associated with the distance uncertainty, which is typically 15% (Muñoz-Mateos et al. 2015). *Upper right panel:* deprojected bar ellipticity as a function of the revised Hubble stage. Error bars indicate the typical 10% uncertainty due to 2D deprojection effects for inclinations lower than 60° (Zou et al. 2014). *Lower right panel:* bar normalised $m = 2$ Fourier amplitude (available for 11 galaxies) vs Hubble type. On the x -axes of the last two panels small random offsets ($\lesssim 0.3$) were added to the T values to avoid point overlapping.

since the presence of inner rings and the abundance of SF fuel along the bar can be closely connected (Neumann et al. 2019; Díaz-García et al. 2020, and references therein).

None of the 12 studied galaxies are currently merging. The degree of isolation of the S⁴G galaxies in our sample was quantified by Laine et al. (2014); Col. 8 of Table 1 lists the projected surface density of galaxies,

$$\Sigma_3^A = \log_{10} \left(\frac{3}{\pi R_3^2} \right), \quad (1)$$

where R_3 is the projected distance to the third nearest neighbour galaxy, given in Mpc. Among the galaxies in our sample with the highest values of Σ_3^A , NGC 4548 and NGC 5850 are known to host signs of recent interactions with a nearby companion (e.g., Higdon et al. 1998; Vollmer et al. 1999), and NGC 4394 likely belongs to the M 85 subgroup of the Virgo Cluster (e.g., Vicari et al. 2002).

The selection of galaxies hosting high-amplitude bars biases our sample towards massive galaxies, spanning $10.2 < \log_{10}(M_{\star}/M_{\odot}) < 11$. This is shown in the upper left panel of Fig. 2, where we display the gas fraction as a function of the

total stellar mass. Total atomic gas (H I) masses are estimated following Giovanelli et al. (1988) and Erwin (2018),

$$M_{\text{HI}} = 2.356 \times 10^5 \times D^2 \times 10^{0.4 \times (17.4 - m_{21c})}, \quad (2)$$

where m_{21c} is the corrected 21 cm line flux in magnitude from HyperLeda. For direct comparison with the physical properties of barred galaxies in the local Universe, in Fig. 2 we display in grey the same values for the S⁴G sample (Díaz-García & Knapen 2020) after imposing an inclination cut-off of 50°, which shows good agreement. Our galaxies have H I gas masses in the range $10^{8.7} - 10^{9.9} M_{\odot}$, resulting in gas fractions (M_{HI}/M_{\star}) spanning 1.5 dex.

2.1. Bar structural parameters

In order to characterise the basic properties of the bars in our sample, we use structural parameters derived from S⁴G 3.6 μm images by Herrera-Endoqui et al. (2015) (bar position angles and sizes) and Díaz-García et al. (2016b) (bar strengths). The disk orientation parameters are from Salo et al. (2015). The bar parameters were deprojected to the disk plane following the

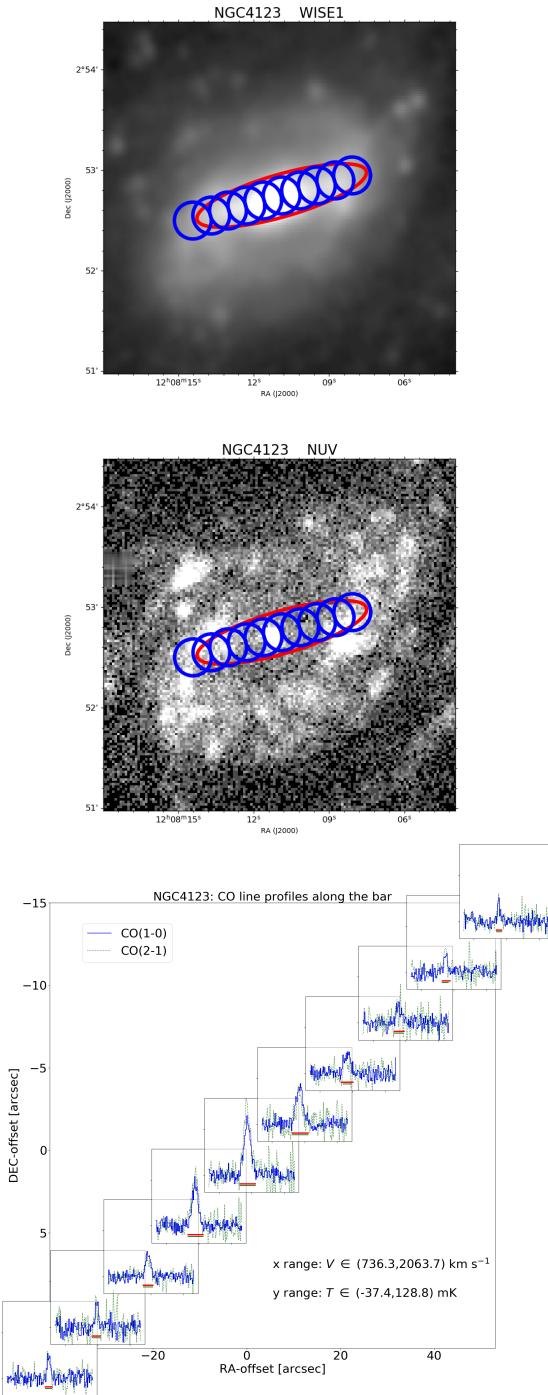


Fig. 3. WISE 1 (top panel) and GALEX far-UV (middle panel) images of NGC 4123. The bar isophote, based on ellipse fitting of $3.6\text{ }\mu\text{m}$ images from Herrera-Endoqui et al. (2015) and Díaz-García et al. (2016b), is highlighted with a red ellipse. We show the circular beams where aperture photometry was performed (IRAM-30 m $22''$ pointings) in blue, spaced by $11''$ and covering the whole bar. Bottom panel: CO emission-line spectra along the bar of NGC 4123. The blue and green lines show the CO(1–0) and CO(2–1) spectra, respectively. The spectra are located on the plot in such a way that their centre corresponds to the offset of the pointing, relative to the central one. The red and green horizontal segments respectively show the zero-level line width of the CO(1–0) and CO(2–1) lines adopted for the determination of the velocity integrated intensities. The velocity (optical convention and in the Local Standard of Rest frame, in units of km s^{-1}) and intensity ranges of the spectra (in the main beam temperature scale, T_{mb} , in units of mK) are also indicated in the lower right corner.

same approach as in Gadotti et al. (2007). Bar sizes and position angles were calculated from the maximum of the ellipticity profile, combined with the constancy of the position angle, at the bar region (Wozniak et al. 1995).

As for NGC 6951, which is not in the S⁴G, we also ran ellipse fitting using the IRAF procedure ellipse (Jedrzejewski 1987) over the *i*-band image from del Río & Cepa (1999). The disk inclination and position angle are taken from Buta et al. (2003): $i = 39.4^\circ$ and $\text{PA} = 143.1^\circ$ (which is consistent with the optimal photometric estimation by Marquez & Moles 1993). The bar parameters are listed in Table 1.

The right panels of Fig. 2 show the distribution in the Hubble sequence of the intrinsic ellipticity of the bars (ϵ_{bar}) and their normalised $m = 2$ Fourier amplitude (A_2). The grey points again correspond to the values in the S⁴G ($i < 50^\circ$) reported by Díaz-García et al. (2016b) (see their Table B.1). The lower left panel of Fig. 2 shows the bar radii in kiloparsecs as a function of the total stellar masses of the host galaxies. The bar parameters in our sample lie well within typical ranges for nearby disk galaxies of similar morphological types found in the literature (see also Martin 1995; Erwin 2005; Aguerri et al. 2009; Pérez et al. 2012). Our sample of galaxies comprises fairly strong bars, as determined from either visual (Buta et al. 2015) or quantitative (A_2 , ϵ_{bar}) measurements.

2.2. IRAM-30 m CO spectra

We observed the CO(1–0) emission line (2.6 mm) along the bars of the 12 galaxies in our sample with the IRAM-30 m radiotelescope at a central frequency of 115 GHz. This allowed us to obtain molecular gas mass surface density profiles (Sect. 3).

Observations were carried out between June and November 2008 using the dual polarisation receivers configuration AB with the $512 \times 1\text{ MHz}$ filterbanks. We used a wobbler switching mode with a wobbler throw of 200 arcsec in azimuthal direction. The pointing was monitored on nearby quasars, Mars, or Jupiter every 60–90 minutes. The weather conditions were good during the observation period, with pointing accuracy better than 4 arcsec (data with poorer pointings were excluded). The average system temperature was $\sim 300\text{ K}$ at 115 GHz on the T_A^* scale. For the data reduction only observations of good quality were chosen (suitable weather conditions and showing a flat baseline). The spectra were summed over the individual positions and the baseline was then subtracted.

We carried out several pointings (depending on the bar length, see Table 1) separated by 11 arcsec and aligned with the bar major axis. The full width at half maximum (FWHM) is $21.5''$ at the frequency of our observations, resolving 1–4 kpc depending on the galaxy (see Col. 13 in Table 1). The spatial distribution of the regions where the CO was measured is displayed in the upper panels of Fig. 3 for NGC 4123 (the same plots can be found in Appendix A for all the galaxies in our sample), based on the visual measurement of the bar size and position angle. We show the WISE 1 and far-UV images, with the CO pointings and the bar ellipse superimposed. The integration times per pointing ranged between 10 and 115 min and was divided into individual observing scans of 4–6 min duration. The resulting CO(1–0) spectra are presented in the lower panel of Fig. 3 and in Appendix A. The x - and y -axes show the offset of the pointings (in RA and DEC) with respect to the centre of the galaxy.

The CO spectra and intensities are expressed on the main beam temperature scale (T_{mb}), which is computed from the forward and beam efficiency: $T_{\text{mb}} = T_A^* \times F_{\text{eff}} / B_{\text{eff}}$ (at 115 GHz,

$F_{\text{eff}} = 0.95$, and $B_{\text{eff}} = 0.75$ for the IRAM-30 m). The velocity resolution varies between 10.4 and 10.5 km s $^{-1}$, which is adequate for our purposes because it is similar to the typical velocity dispersion of molecular clouds (e.g., Stark 1984). We simultaneously observed CO(2–1) (230 GHz); these data are suitable for studying the gas excitation (see Appendix F) and molecular gas masses within a beam of size 0.9–2.0 kpc ($\text{FWHM} \approx 10.75''$), which is two times better resolution than CO(1–0). The obtained CO(1–0) and CO(2–1) spectra are available online at the CDS associated with this publication.

In addition to the statistical error of the velocity-integrated line intensities, a typical flux calibration error has to be taken into account. We adopt the errors determined in Lisenfeld et al. (2019), 15% for CO(1–0) and 30% for CO(2–1), derived from observations of several sources on different days. Despite the higher resolution achieved using CO(2–1), here we opt to use CO(1–0) to estimate molecular gas masses, given the higher calibration error and poorer quality of the CO(2–1) spectra. Even so, in Appendix E we show that the results in this paper are consistently similar when CO(2–1)-based Σ_{mol} estimates are used, confirming that the resolution is not a big issue (on kiloparsec scales) in the statistical trends presented in this paper.

The current availability of homogeneous and state-of-the-art UV and IR datasets (Sects. 2.3 and 2.4) makes the exploitation of these CO(1–0) data especially timely, as it allows accurate measurements of star formation efficiencies (Sect. 5). While the use of interferometric data would maximise the spatial resolution (which is a caveat of this work), single-dish observations guarantee the recovery of all the CO flux along the bar major-axis.

2.3. GALEX ultraviolet imaging

We use the sky-subtracted and masked images from the GALEX/S⁴G UV-IR catalogue by Bouquin et al. (2018), that had been gathered from the GALEX GR6/7 Data Release³ and reduced following Gil de Paz et al. (2007). Specifically, we use near-UV (NUV; $\lambda_{\text{eff}} = 2267\text{\AA}$) and far-UV (FUV; $\lambda_{\text{eff}} = 1516\text{\AA}$) photometry. Following Gil de Paz et al. (2007), we adopt a 14.8% calibration error (0.15 mag) for both the FUV and NUV bands, but we note that Morrissey et al. (2007) give smaller errors (4.7% and 2.8% for FUV and NUV, respectively).

For NGC 6951 (not in S⁴G) we obtained and processed the GALEX NUV GR6/7 image ourselves; to our knowledge no usable FUV image exists for this galaxy. Emission at FUV wavelengths trace recent SF, of the order of several tens to 100 Myr, while NUV traces \lesssim 300 Myr populations (Kennicutt 1998a; Boquien et al. 2014).

Unless stated otherwise, we hereafter use NUV as the ultraviolet SF tracer for the sake of maximising the sample coverage. In Appendix C we show that FUV and NUV yield similar results; this is also discussed in the following sections.

2.4. WISE IR data

Dust re-radiates the absorbed UV in mid-IR wavelengths. In order to trace the dust-enshrouded SF (Sect. 4) we use data from the Wide-field Infrared Survey Explorer (Wright et al. 2010) in the WISE 3 ($\sim 12\mu\text{m}$) and WISE 4 ($\sim 22\mu\text{m}$) passbands, that exist for all the spiral galaxies in our sample (Jarrett et al. 2011). Stellar masses (M_{\star}) within the CO pointings are likewise computed using WISE 1 ($\sim 3.4\mu\text{m}$), while WISE 2 ($\sim 4.6\mu\text{m}$) and WISE 3 are used to calculate IR colours and get a hint of

non-stellar emission and AGN activity. Hereafter, we refer to the WISE passbands with the letter ‘W’.

The WISE images were downloaded from the NASA/IPAC Infrared Science archive⁴. We subtracted the global background using the Python package photutils. In this process, we masked emission that was above 2σ for at least five connected pixels as belonging to the source. Both the mask making and the background subtraction use sigma-clipped statistics to calculate the median and mean values. After subtracting the global background, the WISE maps were converted from digital numbers into janskys (Jy) using the photometric zero points given in Table 1 of the Explanatory Supplement to the WISE All-Sky Data Release Products⁵.

Following Jarrett et al. (2013, see their Sect. 3.5), we performed three corrections for extended source photometry. First, we applied an aperture correction that accounts for the point spread function (PSF) profile fitting used in the WISE absolute photometric calibration. The corrections are 0.034, 0.041, 0.030, and 0.029 mag for the W1, W2, W3, and W4 bands, respectively, and the uncertainty is 1%. The second correction is a colour correction that accounts for the spectral signature of the source convolved with the WISE relative system response. The colour correction is given in Wright et al. (2010) and in the WISE Explanatory Supplement 21, Section IV.4.h. For bands W1, W2, and W4 this corrections is always very small, $\lesssim 1\%$, and we neglect it, but for W3 it can be as large as $\sim 10\%$. The third correction, related to a discrepancy in the calibration between the WISE photometric standard blue stars and red galaxies, only applies to the W4 images and accounts for a factor of 0.92 (Jarrett et al. 2011) for non-bulge dominated objects ($\text{W2-W3} > 1.3$ mag, which is the case for our objects). The calibration accuracy of the WISE maps is 2.4%, 2.8%, 4.5%, and 5.7% for the W1, W2, W3, and W4 images, respectively (Jarrett et al. 2011).

3. Integrated CO spectra and molecular gas mass

From the CO(1–0) spectra measured along the bars (Sect. 2.2), we obtain the velocity-integrated intensities (in K km s $^{-1}$)

$$I_{\text{CO}(1-0)} = \int T \text{d}V, \quad (3)$$

by summing the spectra over the visually determined zero-level velocity width, using the Python *SciPy* package for the Simpson’s rule (i.e., approximating each part of the line with parabolas to perform the numerical integration). We estimate the uncertainty as

$$\text{Error}(I_{\text{CO}(1-0)}) = \sigma \sqrt{\Delta V \delta V}, \quad (4)$$

where δV is the channel width (10.45 km s $^{-1}$), ΔV is the total line width, and σ is the root mean square (rms) noise of the spectrum, calculated outside the line emission. We checked that increasing ΔV by 25% and 50% results in a median difference in $I_{\text{CO}(1-0)}$ of 3.2% and 4.2%, respectively, which is lower than the median error ($\approx 8\%$) on $I_{\text{CO}(1-0)}$ (Eq. (4)). We also confirmed that the velocity-integrated intensities determined within the line width agree well with those obtained from Gaussian fits, the difference being 5.5% (consistent with Alatalo et al. 2016). We hence conclude that the measurement procedure is not a source of major uncertainty in the $I_{\text{CO}(1-0)}$ estimates.

⁴ <https://irsa.ipac.caltech.edu/applications/wise/>

⁵ https://wise2.ipac.caltech.edu/docs/release/allsky/expsup/sec2_3f.html

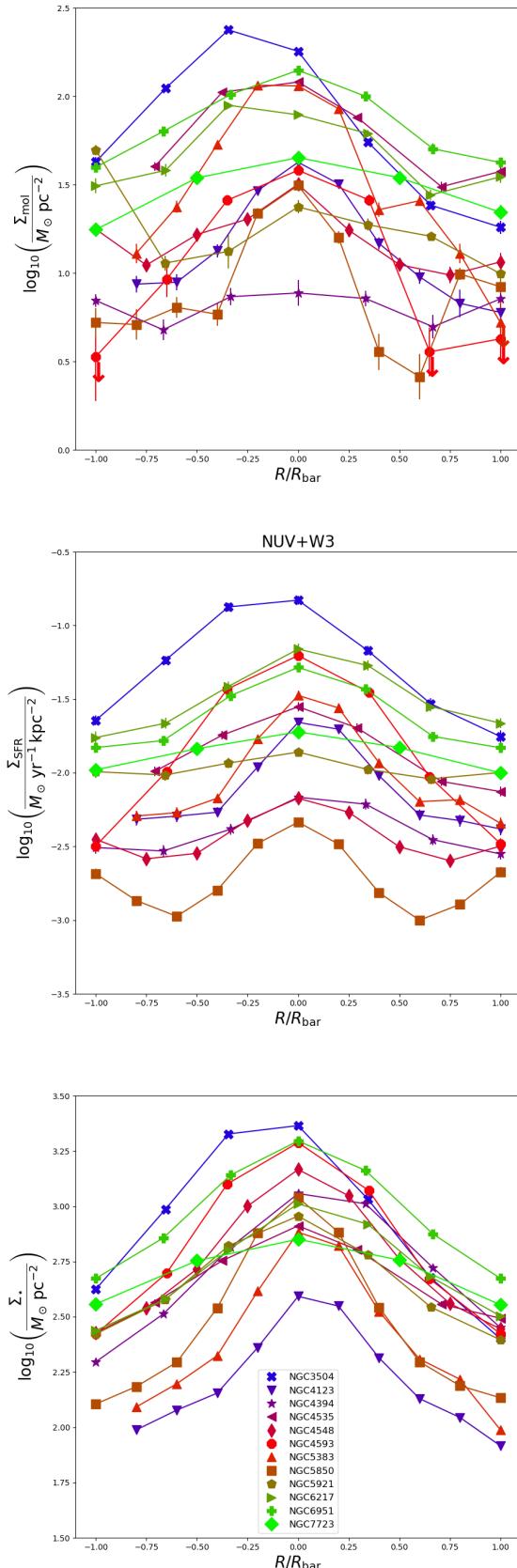


Fig. 4. Profiles of molecular gas mass (top), star formation rate (middle), and stellar surface densities (bottom) along the bar major axis, two-folded so that negative and positive numbers indicate the west and east sides of the bar, respectively. *Upper panel:* the downward-pointing arrows correspond to the pointings with no CO(1–0) detection (Eq. (5)).

In the event of no detection (which is only the case for four spectra of the outermost pointings of NGC 4593 and NGC 5383) we calculated the upper limit as

$$I_{\text{CO}(1-0)} = 3\sigma \sqrt{\Delta V \delta V}. \quad (5)$$

For CO(1–0) we adopted a value of $\Delta V = 200 \text{ km s}^{-1}$, which is the median line width of CO(1–0) in this work. These pointings are shown as arrows in the plots. For CO(2–1) we followed the same steps as in the case of CO(1–0). We adopted $\Delta V = \Delta V$ (CO(1–0)) if CO(1–0) was detected, and $\Delta V = 200 \text{ km s}^{-1}$ otherwise.

We calculate the molecular gas mass (M_{mol}) from the CO(1–0) luminosity (L'_{CO}) following Solomon et al. (1997),

$$L'_{\text{CO}} [\text{K km s}^{-1} \text{ pc}^{-2}] = 3.25 \times 10^7 S_{\text{CO,tot}} \nu_{\text{rest}}^{-2} D^2 (1+z)^{-1}, \quad (6)$$

where $S_{\text{CO,tot}}$ is the CO line flux (in Jy km s^{-1}), D is the distance in Mpc, z the redshift, and ν_{rest} is the rest frequency of the line in GHz, and then

$$M_{\text{mol}} [M_{\odot}] = 1.36 \alpha_{\text{CO}} L'_{\text{CO}}. \quad (7)$$

We adopt the Galactic value $\alpha_{\text{CO}} = \alpha_{\text{CO,gal}} = 3.2 M_{\odot} (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$ (Bolatto et al. 2013). This conversion factor corresponds to $X = N_{\text{H}_2}/I_{\text{CO}} = 2 \times 10^{20} \text{ cm}^{-2} (\text{K km s}^{-1})^{-1}$ (Dickman et al. 1986; Bigiel et al. 2011). The factor 1.36 includes the mass of helium and heavy metals (as in e.g., Bigiel et al. 2011; Bolatto et al. 2013). Molecular gas surface density profiles (Σ_{mol} , in units of $M_{\odot} \text{ pc}^{-2}$) are obtained as

$$\Sigma_{\text{mol}} = 1.36 \alpha_{\text{CO}} I_{\text{CO}} \cos(i), \quad (8)$$

where i corresponds to the disk inclination. Inclination corrections (see also Eqs. (12) and (13)) are performed following Leroy et al. (2008) and Leroy et al. (2019), among others. We do not take into account the 3D geometry of the bar or bulge; this parametrisation is beyond the scope of this paper.

The Σ_{mol} profiles along the bar major axes are shown in the upper panel of Fig. 4. The molecular gas is typically concentrated in the centre of galaxies, and in some cases we detect peaks at one or both edges of the bars (e.g., NGC 4394, NGC 5850, NGC 4548, NGC 5921). In a few cases (e.g., NGC 3504 or NGC 5921) the profiles are rather asymmetric, while in NGC 7723 it is shallower than in the rest of the sample. In Appendix B we list, for each of the pointings, the corresponding J2000 equatorial coordinates, velocity-integrated intensities, molecular gas masses, and surface densities.

We report Σ_{mol} values that are consistent with those in the literature. Bigiel et al. (2008) find average values of Σ_{mol} for disks of local non-barred spiral galaxies of $\sim 10\text{--}15 M_{\odot} \text{ pc}^{-2}$. At the bar region, Σ_{mol} is typically in the range $\approx 20\text{--}90 M_{\odot} \text{ pc}^{-2}$ (Reynaud & Downes 1998; Regan et al. 1999; Kuno et al. 2007; Muraoka et al. 2016; Maeda et al. 2018, 2020; Yajima et al. 2019), under the assumption of a common $\alpha_{\text{CO}} = 3.2 M_{\odot} (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$, as used here. Our use of a bigger sample reveals a somewhat larger range: Σ_{mol} varies by $\approx 1.5\text{--}2$ orders of magnitude between the probed galaxies. All apertures along the bar of NGC 4394 have $\Sigma_{\text{mol}} < 10 M_{\odot} \text{ pc}^{-2}$. The central peaks of NGC 4535, NGC 5383, and NGC 6951 are larger than $50 M_{\odot} \text{ pc}^{-2}$ (see Fig. 4); in the case of NGC 3504 Σ_{mol} goes beyond $100 M_{\odot} \text{ pc}^{-2}$ and is rather asymmetric, which is likely explained by the efficiency of these strong bars in sweeping the disk gas and enhancing its central concentration (e.g., Sellwood & Wilkinson 1993, and references therein).

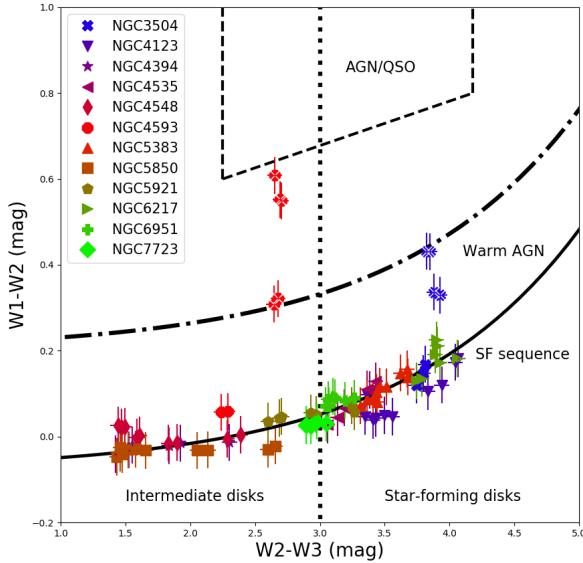


Fig. 5. WISE colour-colour magnitude plot within each IRAM pointing. A solid black line indicates the fit of the star formation sequence from Jarrett et al. (2019, see their Fig. 10 and Eq. (1)), while the vertical dotted line separates the intermediate and star-forming disks. The polygon at the top of the figure, outlined with a dashed line, maps the region of WISE-obscured AGNs, QSOs, LINERs, and ULIRGs. The dash-dotted line indicates the region where low-power Seyferts and LINERs reside (see also Jarrett et al. 2011; Yao et al. 2020), which is where the central pointings of NGC 4593 and NGC 3504 (x symbols) lie.

4. Star formation rates and stellar masses

In order to calculate the SFR and M_{\star} within each IRAM pointing, we perform aperture photometry over the GALEX FUV and NUV and WISE 1-2-3-4 images with the aid of Python’s photutils and astropy packages. We proceed in the following way. We multiply the corresponding background-subtracted GALEX and WISE images, $I(x, y)$, with the IRAM beam pattern (approximated as a normalised Gaussian beam) placed at the position where the CO beam was pointed during the IRAM-30 m CO(1–0) observations (x_0, y_0):

$$I_{\text{beam}}(x, y) = I(x, y) \times \exp\left(-\frac{(x - x_0)^2 + (y - y_0)^2}{2\sigma^2}\right). \quad (9)$$

The Gaussian standard deviation (σ) is related to the FWHM as $\sigma = \text{FWHM}/(2\sqrt{2\ln(2)}) = \text{FWHM}/2.35$. From the resulting maps, $I_{\text{beam}}(x, y)$, we then measure the total fluxes (in Jy). The resulting values are listed in Appendix D, together with the errors, taken as the quadratical sum of the calibration error (which is dominant) and the photometric error. We then compute the luminosities L_{ν} (in $\text{erg s}^{-1} \text{Hz}^{-1}$) and $L = \nu L_{\nu}$ (in erg s^{-1}).

We compute dust attenuation-corrected SFRs combining GALEX NUV emission-line and WISE infrared continuum measurements at 12 μm (W3) and 22 μm (W4). We use the linear ‘hybrid’ SFR calibrations from the GALEX-SDSS-WISE Legacy Catalog (Salim et al. 2016, 2018), as reported by Leroy et al. (2019, see their Table. 7):

$$\text{SFR}_{\text{NUV}, \text{W3}} [M_{\odot} \text{yr}^{-1}] = \frac{L_{\text{NUV}} [\text{erg s}^{-1}]}{10^{43.24}} + \frac{L_{\text{W3}} [\text{erg s}^{-1}]}{10^{42.86}}, \quad (10)$$

$$\text{SFR}_{\text{NUV}, \text{W4}} [M_{\odot} \text{yr}^{-1}] = \frac{L_{\text{NUV}} [\text{erg s}^{-1}]}{10^{43.24}} + \frac{L_{\text{W4}} [\text{erg s}^{-1}]}{10^{42.79}}. \quad (11)$$

W3 and W4 are robust mid-IR SF tracers used to trace dust-enshrouded SF (e.g., Cluver et al. 2017; Hall et al. 2018). The

W3 images are an order of magnitude more sensitive to SF than the W4 (e.g., Leroy et al. 2019). Nevertheless, Leroy et al. (2019) argue that W3 suffers from larger systematic uncertainties and report variations in the W3-to-W4 ratio due to PAHs (Engelbracht et al. 2005). The centre of the W3 band is indeed close to the 11.3 μm PAH emission, but is still dominated by continuum because of its relatively large bandwidth (see discussion in Cluver et al. 2017). PAH abundances are high in regions of active SF and are likely produced in molecular clouds (Sandstrom et al. 2010); PAHs can be destroyed near stars and/or AGNs where the radiation field is very intense.

In order to make sure that our analysis is not biased by the choice of the SF tracer, we tested different SF recipes (see Appendix C), including estimates from NUV, FUV, and W3-4, and different hybrid combinations. We checked and confirmed that the trends presented in this paper are qualitatively the same when FUV is used instead of NUV, but we found that the SFEs and the scatter of the KS law are somewhat sensitive to wavelength (see discussion in Sect. 5). Interestingly enough, the bulk of the SFRs in our galaxies are traced by WISE (Fig. C.2). Therefore, whether we use NUV or FUV as a SF tracer does not make a big difference in the results presented in the next sections as long as they are combined with W3 or W4 (see Fig. C.1). Unless stated otherwise, we estimate SFRs from NUV and W3 combined.

The star formation rate surface density (Σ_{SFR}), corrected for inclination (i), in each aperture is calculated as

$$\Sigma_{\text{SFR}} [M_{\odot} \text{yr}^{-1} \text{pc}^{-2}] = \frac{\text{SFR} \times \cos(i)}{1.13 \times (\text{FWHM}_{\text{pc}})^2}, \quad (12)$$

where the denominator corresponds to the 2D integral of a Gaussian beam, which we take as the FWHM (21.5'') of the IRAM beam (in units of pc) at the distance of the object, FWHM_{pc} . This expression, together with Eq. (9), means that we calculate the weighted mean value of the SFR with the IRAM beam, where the weighting function is the IRAM beam shape.

We calculate the stellar mass from the W1 fluxes within each IRAM pointing ($I_{3.4 \mu\text{m}, \text{beam}}$) adopting a constant mass-to-light ratio $\Upsilon_{\star}^{3.6} = 0.53$ (in solar units at 3.6 μm , Eskew et al. 2012) and neglecting the small difference between 3.6 μm (the Spitzer-IRAC 1 central wavelength) and 3.4 μm (the W1 central wavelength), as done in Leroy et al. (2019), among others. With this we derive the inclination-corrected stellar mass surface density as

$$\begin{aligned} \Sigma_{\star} [M_{\odot} \text{pc}^{-2}] &= \frac{L_{3.4, \odot} \times \Upsilon_{\star}^{3.6} \times \cos(i)}{1.13 \times \text{FWHM}^2} \\ &= 1.88 \times 10^5 \times I_{3.4 \mu\text{m}, \text{beam}} [\text{Jy}] \times \Upsilon_{\star}^{3.6} \times \cos(i), \end{aligned} \quad (13)$$

where $L_{3.4, \odot}$ is the luminosity νL_{ν} at 3.4 μm in units of solar luminosity at 3.6 μm ($1.4 \times 10^{32} \text{ erg s}^{-1}$, Cook et al. 2014). W1-based stellar masses are likewise utilised to normalise SFRs and obtain specific star formation rates (sSFRs) across bars. We opted not to estimate M_{\star} from WISE 1 and 2 colours (Zucker et al. 2016); the resulting Σ_{\star} profiles were not reliable (e.g., minima of Σ_{\star} in the centre of NGC 4593), quite possibly because of non-stellar contaminants such as hot dust and PAHs (see e.g., Zibetti & Groves 2011; Meidt et al. 2012).

In order to assess whether AGNs might have a contribution to the computed WISE fluxes, we construct WISE colour-colour magnitude diagrams (W1–W2 versus W3–W2) for all galaxies and within each aperture (Fig. 5) and include the classifications of Jarrett et al. (2017), presented in Jarrett et al. (2019, see their Fig. 10) into star-forming galaxies and AGNs or quasars (QSOs). The majority of the sampled regions in our bars belong

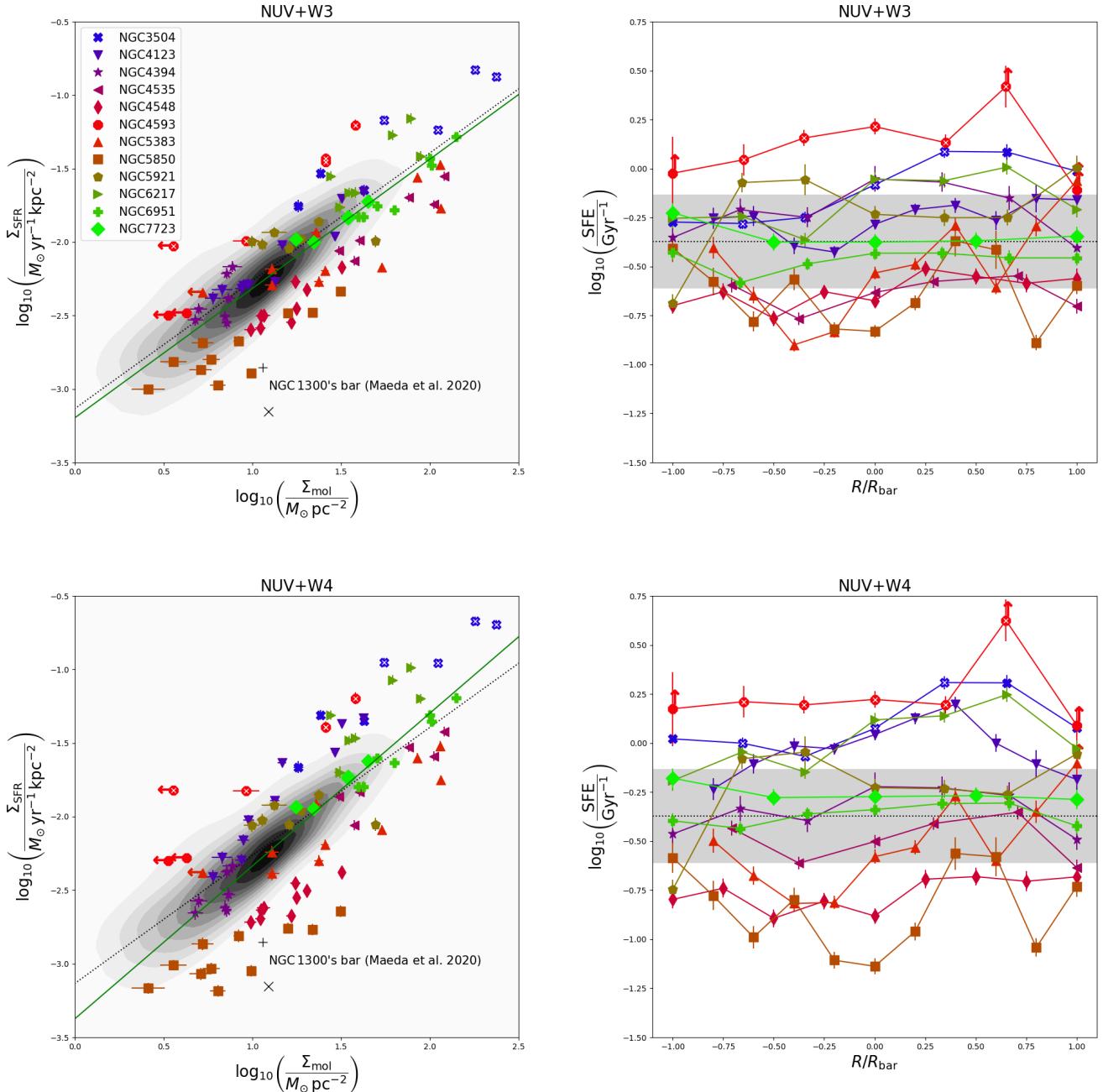


Fig. 6. Kennicutt-Schmidt law and star formation efficiency within bars. *Left panels:* star formation rate surface density traced from NUV combined with WISE 3 (upper panels) and WISE 4 (lower panels) vs molecular gas surface density, traced from the CO(1–0) emission. The colour palette and symbols indicate different galaxies (see legend). The grey contours show the measurements from Bigiel et al. (2011), obtained from FUV and 22 μm images, whose linear fit is also shown with a black dotted line. The white x symbols indicate those points with high W2–W1 (Fig. 5) that may correspond to less reliable SFR estimates due to AGNs. The leftward arrows correspond to the pointings with no CO(1–0) detection (Eq. (5)). The green line corresponds to the linear fit to the cloud of points. The grey cross and x symbol indicate the Σ_{SFR} and Σ_{mol} values at the bar region of NGC 1300 reported by Maeda et al. (2020) from the combination of FUV and W4 ('Bar-A' and 'Bar-B' in their Table B.1). *Right panels:* profiles along the bar major axis of star formation efficiency. The grey rectangle traces the median SFE (black dotted line) plus or minus 1 σ from Bigiel et al. (2011).

to the SF sequence; roughly half belong to the ‘intermediate galaxy’ class ($W2-W3 < 3$, mostly quiescent bars as reported by Díaz-García et al. 2020) and the other half to active star-forming disks ($W2-W3 > 3$). There are, however, a few points that lie above this SF sequence, corresponding to the central regions of NGC 4593 and NGC 3504. NGC 4593 is classified a Seyfert 1 by Véron-Cetty & Véron (2010) which can explain the elevated $W1-W2$ values that bring it close to the AGN–QSO

region. The inferred SFRs and stellar masses within those apertures are thus less reliable. Véron-Cetty & Véron (2010) report nuclear H II regions in NGC 3504, which show high $W2-W3$ and $W1-W2$ values. Interestingly, the WISE colours of NGC 4548, which is classified as a LINER (Véron-Cetty & Véron 2010), are not affected by this nuclear activity and are entirely compatible with the SF sequence. This is also the case of the active galaxy NGC 6951 (Ho et al. 1997; Pérez et al. 2000). We note that in

Table 2. Statistics of the star formation efficiency (SFE) along bars.

$\log_{10}\left(\frac{\text{SFE}}{\text{Gyr}^{-1}}\right)$	$ R/R_{\text{bar}} < 1/4$	$1/4 < R/R_{\text{bar}} < 1/2$	$1/2 < R/R_{\text{bar}} < 3/4$	$3/4 < R/R_{\text{bar}} $
NUV+W3	-0.41 ($\sigma = 0.30$) [N = 18]	-0.29 ($\sigma = 0.28$) [N = 20]	-0.33 ($\sigma = 0.25$) [N = 19]	-0.39 ($\sigma = 0.23$) [N = 24]
NUV+W4	-0.39 ($\sigma = 0.45$) [N = 18]	-0.21 ($\sigma = 0.33$) [N = 20]	-0.25 ($\sigma = 0.34$) [N = 19]	-0.41 ($\sigma = 0.30$) [N = 24]
FUV+W3	-0.44 ($\sigma = 0.29$) [N = 17]	-0.32 ($\sigma = 0.29$) [N = 18]	-0.38 ($\sigma = 0.24$) [N = 17]	-0.44 ($\sigma = 0.22$) [N = 22]
FUV+W4	-0.45 ($\sigma = 0.46$) [N = 17]	-0.26 ($\sigma = 0.36$) [N = 18]	-0.32 ($\sigma = 0.34$) [N = 17]	-0.49 ($\sigma = 0.31$) [N = 22]

Notes. For each radial interval (2nd–5th columns) and hybrid star formation tracer (1st column, see text for details) we indicate the mean SFE. In parenthesis we show standard deviation (σ). In brackets we indicate the number of elements (i.e., IRAM-30 m pointings with CO(1–0) detection) within the radial bin.

some cases the identification of AGN footprints may be hindered by the large size of the beams used. In Fig. 5 and the following figures we use x symbols to indicate pointings whose elevated W2–W1 colours are suggestive of the presence of AGNs.

The profiles along bars of star formation rate surface density and stellar surface density are shown in the central and lower panels of Fig. 4, respectively. Their shapes are quite similar to those of molecular gas surface density, shown in the upper panel. All Σ_{SFR} profiles are centrally peaked; some of these peaks are associated with the presence of nuclear rings (e.g., NGC 5383 and NGC 5850), including the case of NGC 6951 (e.g., Kohno et al. 1999). A few galaxies (e.g., NGC 4548 and NGC 5850) show humps at the bar ends. In general, three galaxies in our sample (NGC 4394, NGC 4548, and NGC 5850) show low levels of SF along the bar major-axis. The visual inspection of the FUV images of NGC 4123, NGC 6951, and NGC 6217 (only in the north-east) also reveal SF gaps along the bar (see e.g., Díaz-García et al. 2020); these gaps are typically smaller than the IRAM-30 m pointings and cannot be resolved. The Σ_{SFR} values directly obtained from FUV in this work (Appendix C) for S0/a-Sc galaxies are consistent with the average profiles along bars presented by Díaz-García et al. (2020) for $T < 5$. They report a standard deviation of 1–1.5 mag of the FUV emission in the inner parts of galaxies with $\log_{10}(M_{\star}/M_{\odot}) = 10\text{--}11$.

None of the Σ_{\star} profiles show enhancements at the bar ends. This is not surprising as these galaxies do not have ansae structures (stellar blobs at the end of the bars; Danby 1965) even though most ansae are detected in early-type spiral galaxies, like some of the ones probed in this work (e.g., Martinez-Valpuesta et al. 2007; Laurikainen et al. 2007). Our galaxies are characterised by large central mass concentrations. At least four galaxies (NGC 4394, NGC 4548, NGC 4593, and NGC 5850) host barlenses (which are lens-like stellar structures embedded in bars postulated to be the face-on counterparts of boxy or peanut bulges; e.g., Laurikainen et al. 2014; Athanassoula et al. 2015), nuclear rings and bars (NGC 5850), or nuclear lenses (NGC 3504 and NGC 4394), according to Buta et al. (2015) and/or the NED, as shown in Table 1.

The resulting values of Σ_{SFR} and Σ_{\star} are tabulated in Appendix D. The relationship between the values of SFRs and M_{\star} , presented here, and molecular gas masses is analysed in the next section.

5. Star formation efficiency along bars

Here we study the relationship between Σ_{SFR} and Σ_{mol} in bars, so that we can determine whether the SF activity is mainly controlled by the H₂ mass (i.e., GMCs and diffuse molecular gas) or by distinct physical conditions in bars. We quantify the link between star formation and molecular gas by measuring the star formation efficiency $\text{SFE} = \Sigma_{\text{SFR}}/\Sigma_{\text{mol}}$. SFE is the inverse of the

gas depletion time (τ_{dep}), which is the time needed to use up the existing gas reservoirs at a given SFR.

The SFEs are calculated within each of the pointings sampling the bars in our sample. Σ_{SFR} and Σ_{mol} are correlated even within bars (see Fig. 6). This means that, on average, the presence of ionised regions in strong bars is to a large extent controlled by the availability of H₂ gas, even though the timescales for SF might well be larger than those of the inflow of GMCs along the bars (see e.g., the models by Pérez & Freeman 2006). The local KS law depends on the spatial resolution of the data; for example, a higher slope of the scaling relation has been found for a coarser spatial resolution (e.g., Verley et al. 2010). Therefore it is instructive to compare these results to those obtained from CO(2–1) which has a two times higher angular resolution. In Appendix E we show that the reported $\Sigma_{\text{SFR}} - \Sigma_{\text{mol}}$ relation is similar when CO(2–1) spectra (beam size of $\sim 1\text{--}2$ kpc) are used to compute Σ_{mol} (see Fig. E.1) instead of CO(1–0). On the other hand, our pointings are sufficiently large to probe physical areas ($\gtrsim 500$ pc) where the molecular gas mass and the star formation rates are tightly correlated, as increasing resolution can eventually wash out the scaling reported on large scales and the KS law breaks down (Schrua et al. 2010; Kruissen & Longmore 2014) (for a similar discussion based on the analysis of the spatially resolved star-forming main sequence, see Hall et al. 2018, and references therein).

The SFEs vary between different galaxies by up to an order of magnitude (Fig. 6), ranging for $0.1 \lesssim \text{SFE/Gyr}^{-1} \lesssim 1.8$ when Σ_{SFR} is estimated from the combination NUV and W3. For a given galaxy there is no systematic trend in SFE along the bar major axis (right panel of Fig. 6). In Table 2 we show the mean and dispersion of SFE as a function of R/R_{bar} for the different hybrid SF recipes used in this work (Sect. 4). The mean SFE in the central region ($|R/R_{\text{bar}}| < 1/4$) and bar-ends ($|R/R_{\text{bar}}| > 3/4$) is $\sim 0.4 \text{ Gyr}^{-1}$, which is roughly 0.1 dex larger than the average in the mid-part of the bar ($1/4 < |R/R_{\text{bar}}| < 3/4$). However, this difference is smaller than the dispersion of SFE across bars (~ 0.27 dex) and SFE gradients are not obvious in individual galaxies; we thus conclude that SFEs in the bar are not remarkably different than those in the central regions. While the SFE depends on the adopted SF tracer and recipe to some extent (depending on the choice the mean integrated SFE within bars varies by a maximum of 0.08 dex; Table 2), the discussed radial trends versus radius are qualitatively the same.

We note that a smaller scatter in the $\Sigma_{\text{SFR}} - \Sigma_{\text{mol}}$ relation is obtained when FUV or NUV are combined with W3, instead of W4, to account for the dust (Fig. 6 and Appendix C). This is in agreement with Cluver et al. (2017), who analysed the relation between the luminosity at W3 and W4 passbands and the SF estimated from the total infrared luminosity: they conclude that the W3 SFR relation has a 1σ scatter of 0.15, while the W4 SFR relation shows more scatter (0.18 dex). In this work we obtain

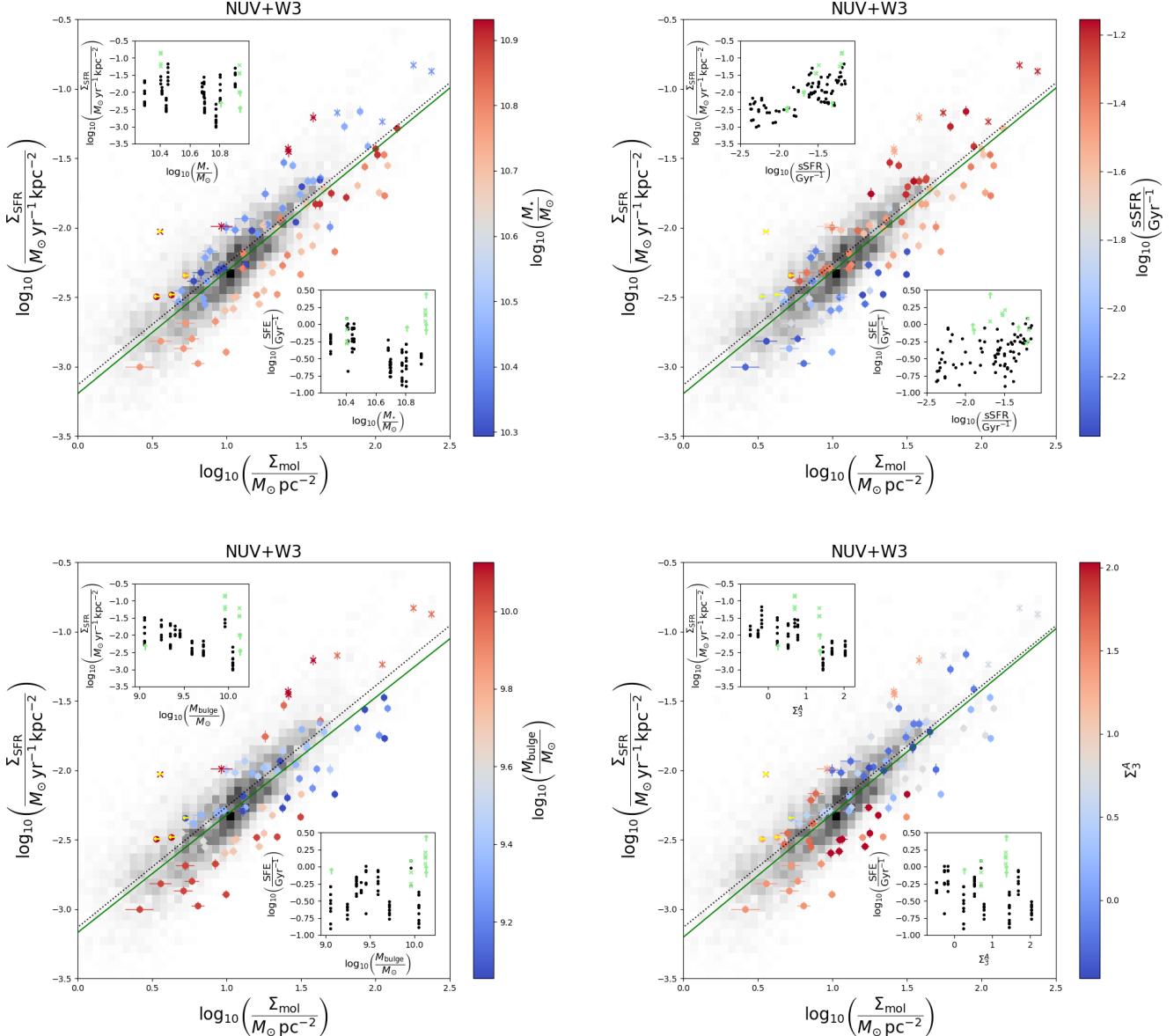


Fig. 7. Same as left panel of Fig. 6, but the colour-coding (see colour bars on the right) indicates the total stellar mass of the host galaxy (*top left*), the specific star formation of each aperture along the bars (*top right*), the bulge stellar mass (*bottom left*), and the projected surface density to the third-nearest neighbour galaxy (*bottom right*). The background black and grey density plots trace the measurements from Bigiel et al. (2011). The x symbols indicate pointings whose elevated W2–W1 colours suggest the presence of AGNs. The insets show Σ_{SFR} (upper left corners) and SFE (lower right corners) vs the parameters used for the colour-coding of the main plot; black solid points correspond to the pointings with CO(1–0) detection whose WISE colours lie in the SF sequence, while non-detections and AGN candidates are respectively highlighted with green x symbols and arrows (i.e., less reliable SFEs and SFRs).

a 1σ scatter of 0.27 dex in the KS law when using $\text{SFR}_{\text{NUV},\text{W}3}$; the scatter increases to 0.37 dex for $\text{SFR}_{\text{NUV},\text{W}4}$. In addition, the 24 μm diffuse emission within the IRAM-30 m pointing may not be related to SFR, but rather to the circumstellar dusty envelopes of AGB stars (e.g., Verley et al. 2009); this can be another physical reason for the above-mentioned scatter when using W4.

On average, our SFE within bars are consistent with median values SFEs obtained in the literature for spiral galaxies. In particular, we compare our measurements with those reported by Bigiel et al. (2011) (median SFE 0.43 Gyr^{-1} , $\sigma = 0.24$ dex, under the assumption of the same α_{CO} applied in this work), shown in black and grey in the left panels of Fig. 6 (see also Figs. C.2 and C.1). There is good agreement between the linear

fits of the KS law in this work and that of Bigiel et al. (2011) for $\text{SFR}_{\text{NUV},\text{W}3}$, shown with solid green and black dotted lines in Fig. 6, respectively. We again report differences between W3 and W4: the slope of the KS law is $\sim 15\%$ larger for $\text{SFR}_{\text{NUV},\text{W}4}$ than for $\text{SFR}_{\text{NUV},\text{W}3}$ (see Table C.1).

The SFEs are almost a factor of 2 lower than the average in a few of our galaxies, such as NGC 5850 ($<\text{SFE}>=0.25\pm0.03 \text{ Gyr}^{-1}$) and NGC 4548 ($<\text{SFE}>=0.24\pm0.01 \text{ Gyr}^{-1}$); the resulting SFEs are even lower (by a factor of ~ 1.5) when using $\text{SFR}_{\text{NUV},\text{W}4}$. This is also the case of NGC 4535 ($<\text{SFE}>=0.23\pm0.02 \text{ Gyr}^{-1}$), which hosts H II regions in the mid-parts of the bars. NGC 3504 and NGC 4593 are the galaxies in the sample with the highest SFE values (even larger than 1 Gyr^{-1}); we note, however, that the measurement in the central pointings may

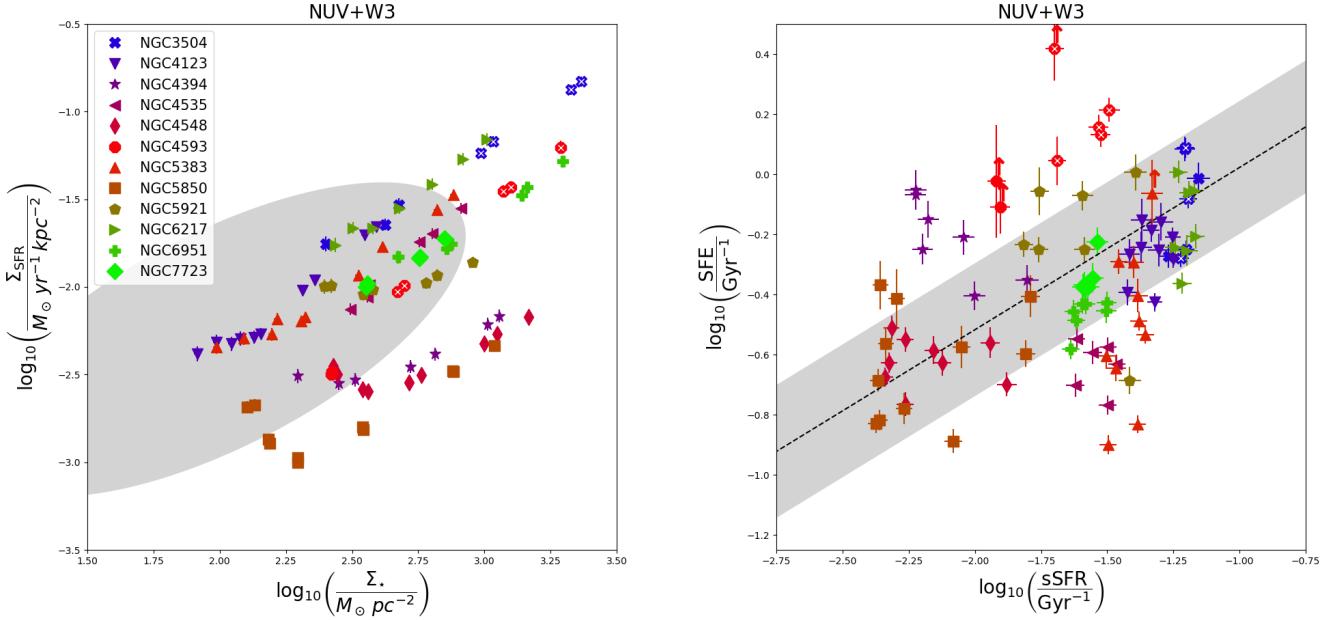


Fig. 8. Scaling relations involving surface densities of stellar mass, molecular gas mass, and star formation rate. *Left:* star-forming main sequence measured within the CO pointings covering the bars, for each galaxy in our sample (see legend). Star formation rates are calculated from the combination of NUV and WISE 3 images (Eq. (10)), while stellar masses are estimated from WISE 1 images (Eq. (13)). The grey ellipse corresponds to the spatially resolved main sequence reported by Ellison et al. (2018), traced from their Fig. 4 (see Rosado-Belza et al. 2019), based on Mapping Nearby Galaxies at APO (MaNGA; Bundy et al. 2015). *Right:* star formation efficiency vs specific star formation. The grey area corresponds to the best fit (black dotted line) reported by Saintonge et al. (2011, COLD GASS, see their Table 1; we only use their secure detections) plus or minus 1σ . For the sake of consistency, we add the mass of helium and heavy metals ($\times 1.36$) to the molecular gas masses used to calculate SFEs in COLD GASS. The colours and symbols are the same as in Fig. 6.

be affected to some extent by AGN contamination (x symbols in Figs. 6 and 7), as discussed in Sect. 4.

It is not trivial to explain the differences in SFE between different galaxies as they are all grand design hosting high-amplitude bars. Interestingly enough, we find a segregation as a function of the total stellar mass of the host galaxy (Fig. 7): the massive galaxies in our sample ($10^{10.7} \lesssim M_*/M_\odot \lesssim 10^{11}$) have lower SFEs than their fainter counterparts ($10^{10.2} \lesssim M_*/M_\odot \lesssim 10^7$). In Fig. 7 we also show that the bars with the lowest Σ_{SFR} and sSFR values cohabit with massive bulges, according to 2D photometric decompositions by Salo et al. (2015), and higher Σ_3^A (Table 1, Sect. 2). Naturally, these statistical trends may imply that environmental and internal effects control the global gas availability and KS law, in the bars and elsewhere in the disk. These trends are also clear when we use FUV or WISE 4 for the SFR estimates (Fig. C.3). This is further discussed in Sect. 6.2.

One possible source of uncertainty in our analysis of SFEs in bars is the assumption of a constant CO-to-H₂ conversion factor, using the same value as found in our Galaxy (Bolatto et al. 2013). It is well accepted that α_{CO} can vary in different types of galaxies, and within galaxies (e.g., Sandstrom et al. 2013; Blanc et al. 2013), mainly due to its dependence on metallicity and star formation activity (see Bolatto et al. 2013, and references therein). In particular, α_{CO} can be considerably higher for low-metallicity galaxies (below $12 + \log(\text{O/H}) \sim 8.4$, e.g., Leroy et al. 2011; Bolatto et al. 2013; Hunt et al. 2015) and lower by a factor of 3–10 in extreme starbursts, as in ultraluminous infrared galaxies (ULIRGs) (Downes & Solomon 1998, 2003). Our sample does not contain any low-metallicity galaxies nor does it contain extreme starburst galaxies, as we can see from Fig. 8 (left panel) showing that all our pointings follow the spatially resolved star-forming main sequence (Ellison et al. 2018, see their Fig. 4) or lie below it (interestingly in the case of

NGC 4394, NGC 4548, and NGC 5850, which lack H II knots and clumps in the mid-part of the bar; Díaz-García et al. 2020). This means that we do not expect large variations in α_{CO} due to a low metallicity or a high SFR within our sample.

A lower CO-to-H₂ conversion factor in the bar region than in the rest of the disk has been suggested in previous observational work (e.g., the case of NGC 3627, studied by Morokuma-Matsui et al. 2015), but remains a matter of debate (see also Watanabe et al. 2011). Sorai et al. (2012) argues that SFEs in the bar of Maffei 2 can be underestimated by a factor of 0.5–0.8 by assuming a constant α_{CO} , which is linked to the presence of diffuse, non-bound, and non-optically thick molecular gas (see also the analysis of NGC 1300 by Maeda et al. 2020). We cannot entirely exclude this possibility based on our data, and we thus acknowledge that part of the dispersion of SFE along bars might be a consequence of the α_{CO} uncertainty. However, even if α_{CO} were lower, we can still conclude that SFE is not systematically low along the bars probed in this work. We also note that the central pointings of AGN hosts can have more uncertain α_{CO} (Papadopoulos & Sequist 1999; Bolatto et al. 2013). Improving the α_{CO} calibration is beyond the scope of this paper and is not within reach with the data presented here.

Finally, in order to illustrate the consistency of our results with COLD GASS (Saintonge et al. 2011), in the right panel of Fig. 8 we show the star formation efficiency versus specific star formation rate within bars. The grey area corresponds to the best fit ($\pm 1\sigma$) reported in COLD GASS based on global integrated measurements of the CO(1–0) line from IRAM-30 m for 222 galaxies (with and without bars); the authors show the consistency in their reported H₂ depletion times with integrated and resolved measurements in HERACLES (Leroy et al. 2009). Altogether, the slope of the relationship in this work coincides with the best fit reported by Saintonge et al. (2011).

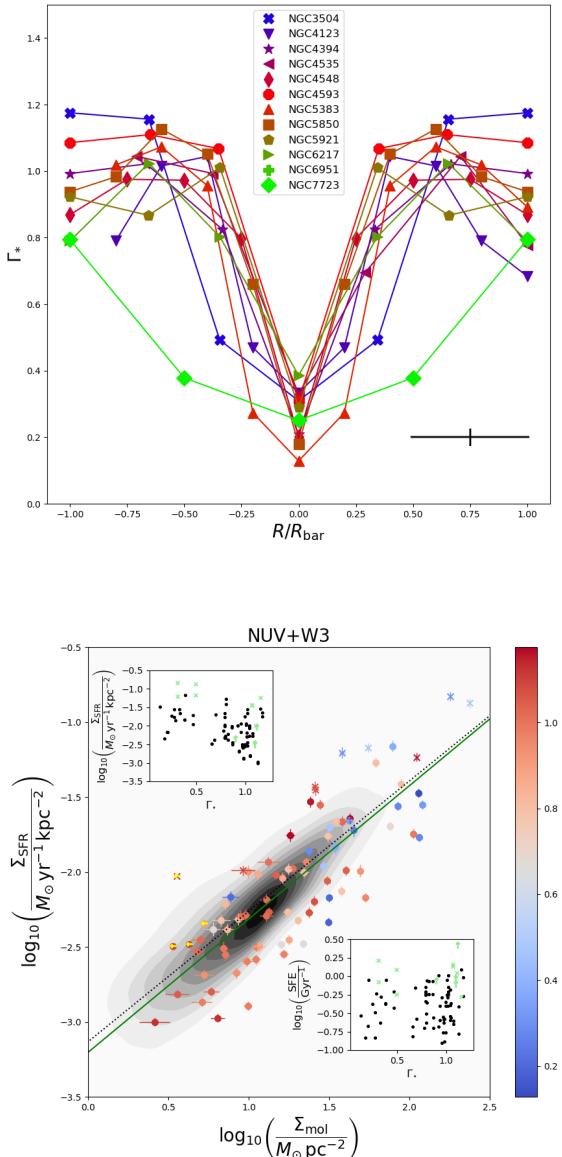


Fig. 9. Shear and its connection to star formation. *Upper panel:* shear evaluated from the stellar component of the circular velocity within each IRAM-30 m aperture as a function of the radius along the bar major axis. The black cross indicates the mean bar range covered by the IRAM-30 m pointings (x -axis, i.e., $\text{FWHM}/R_{\text{bar}}$) and error on Γ_* . The latter corresponds to a 13% typical uncertainty on the slope of the disk+bulge rotation curve (Díaz-García et al. 2016b), which is associated with the uncertainty on the disk thickness determination. *Lower panel:* as in Fig. 7, but colour-coded by the stellar contribution to the shear. The grey contours show the measurements from Bigiel et al. (2011).

Link between gravitational torques, shear, and star formation

The shear (Γ) in a galaxy is expected to be correlated with the star formation rate (Seigar 2005). The value of Γ can be estimated from the slope of the rotation curves (V),

$$\Gamma = -d\ln \Omega / d\ln r, \quad (14)$$

where $\Omega(r) = V/r$ is the angular velocity (e.g., Fujii et al. 2018) at a given radius r , and $\Gamma = 1$ in the flat regime. However, Γ cannot be directly estimated for the galaxies in our sample as we lack ancillary high-resolution rotation curves or 2D

integral field unit (IFU) or Fabry-Pérot (FP) dynamics. However we have at our disposal the stellar component (bulge+disk) of the circular velocity, from which we here calculate the shear (Γ_*). Specifically, Díaz-García et al. (2016b) estimated the radial force field (F_R) by applying the NIR-QB code (Salo et al. 1999; Laurikainen & Salo 2002) to $3.6 \mu\text{m}$ S⁴G images, and calculated the stellar contribution to the circular velocity as

$$V_{\text{disk+bulge}}(r) = \sqrt{\Upsilon_{3.6 \mu\text{m}} \langle F_R(r) \rangle r}, \quad (15)$$

where r is the galactocentric radius, F_R corresponds to the radial force obtained for $M/L=1$, and $\Upsilon_{3.6 \mu\text{m}} = 0.53$ is the mass-to-light ratio at $3.6 \mu\text{m}$ obtained by Eskew et al. (2012), which is assumed to be constant throughout the disk. Here we evaluate Γ_* in the same radial ranges where the IRAM-30 m dish was pointed. The value of Γ_* is a lower bound of Γ as it does not include the contribution of dark matter and gas components to the potential well, nor the contribution of non-circular motions, and thus the conclusions in this section on the Γ -SFR connection are not definitive.

As expected, the shear rate parameter is lowest in the central regions of the galaxies (Fig. 9, upper panel) where the inner slope of the rotation curve is highest. We find that the SFR surface density is highest in the pointings with lowest shear rates (blue points in the lower panel of Fig. 9). The implications of the connection between Γ and SFR are addressed in Sect. 6.2. Having characterised the SFEs, SFRs, and molecular gas masses within bars, we next discuss the results in light of galaxy evolution.

6. Discussion: Physical mechanisms that drive star formation in galactic bars

Bars in galaxies have been studied intensively during the last decades (e.g., Athanassoula 2013; Kormendy 2013; Buta 2013, and references therein), mainly with the motivation to investigate their role in the transport of material towards the central regions of galaxies (e.g., Kormendy & Kennicutt 2004; Bureau & Athanassoula 2005; Pérez et al. 2009; Seidel et al. 2015; Zurita et al. 2021). The fraction of dense gas measured by HCN emission, for example, can be enhanced in bars (such as in NGC 2903; Leon et al. 2008). In some cases the funnelled gas is known to enable episodes of SF and nuclear activity (e.g., Knapen et al. 2000; Ellison et al. 2011; Oh et al. 2012; Cisternas et al. 2013; Emsellem et al. 2015; Florido et al. 2015; Vera et al. 2016; Catalán-Torrecilla et al. 2017; Lin et al. 2020) and can increase the central stellar mass (e.g., Díaz-García et al. 2016a).

The relation between H₂ and SF within bars has so far received much less attention, and is not well understood yet. This is of particular interest given the extreme physical conditions of bars as places with strong shear, shocks, and non-circular velocities (e.g., Pence & Blackman 1984; Zurita et al. 2004) or significant magnetic field strengths (e.g., Beck et al. 2002). These make them ideal testing rooms for exploring the parameters that trigger or inhibit SF since there are no obvious differences between the H II regions in bars and those in the outer disks (e.g., Martin & Friedli 1999). We traced the molecular gas (Sect. 3) and SF (Sect. 4) along the bars of 12 spiral galaxies (Sect. 2). These galaxies host bars with different degrees of SF, with no prior selection on gas content. We took a different approach in comparison to previous studies that either observed individual objects or a small set of galaxies (see e.g., Muraoka et al. 2016; Yajima et al. 2019), often reporting a low

SFE in bars (see [Maeda et al. 2020](#), and references therein). Instead, we used a fairly large sample of 12 galaxies and show that, on average, the SFE is not remarkably diminished in bars relative to the typical values found in spiral galaxies (for further details, see Sect. 6.2).

6.1. Spatial distribution of SF in bars

Different spatial distributions for the recent SF in bars, as traced by the H α emission, have been reported (e.g., [Sheth et al. 2002](#); [Zurita & Pérez 2008](#); [Neumann et al. 2019](#); [Fraser-Mckelvie et al. 2020](#); [Díaz-García et al. 2020](#)). When any H α is detected in the bar it seems to come (1) from H II regions distributed along the bar, (2) from the nuclear or circumnuclear region with little or no emission from the bar, or (3) from the bar and the nuclear region (i.e., an intermediate case between 1 and 2) ([Martin & Friedli 1997](#); [Verley et al. 2007a](#)). All of the strongly barred galaxies studied in this work have prominent circumnuclear SF relative to the underlying disk, and show different degrees of SF at the mid-part of the bar (Sect. 4).

The different H α bar morphologies are interpreted as stages of an evolutionary sequence of the bar by [Verley et al. \(2007a\)](#) based on the AMIGA project ([Verdes-Montenegro et al. 2005](#)) (see also earlier work by [Martin & Friedli 1997](#)). The sequence starts with SF distributed along the bar. The gas is then progressively depopulated from the bar through gas inflow towards the centre of the galaxy until H α emission is only seen in the nuclear or circumnuclear region. The later stage would occupy most of the bar life times, given the observed frequency. In this scenario the absence of SF in many barred galaxies is explained by the lack of gas. Furthermore, [Díaz-García et al. \(2020\)](#) show that strongly barred early-type spiral galaxies are characterised by a ~ 0.5 mag brighter central UV emission (i.e., $\gtrsim 50\%$ larger Σ_{SFR}), compared to their weakly barred counterparts. This scenario can indeed be explained by the efficiency of a strong bar potential at inducing central gas concentration and starbursts. Many of the barred galaxies in our sample have central Σ_{SFR} values that are substantially larger than those in [Bigiel et al. \(2011\)](#) (Fig. 6) or [Ellison et al. \(2018\)](#), based on MaNGA DR13 datacubes (Fig. 8), and those correspond to places with high specific SFRs (Fig. 7).

In addition, [Díaz-García et al. \(2020\)](#) showed that, on average, inner-ringed galaxies are characterised by a UV and H α deficit in the central parts (see also [Neumann et al. 2019](#), and references therein) in both barred and non-barred systems. This picture can be partially explained by gas being trapped at the 1/4 ultraharmonic resonance ([Schwarz 1984](#); [Buta & Combes 1996](#)), where inner rings tend to form (e.g., [Sellwood & Wilkinson 1993](#); [Rautiainen & Salo 2000](#); [Díaz-García et al. 2019a](#)), slowing down its migration to the nuclear regions (e.g., [Neumann et al. 2019](#)). In this work the almost ubiquitous presence of inner pseudo-rings in our sample ($> 80\%$, see Sect. 2 and Table 1) does not systematically imply a lack of molecular gas and SF along the bars. This suggests that either gas is not irreversibly caught at the 1/4 ultraharmonic resonance or that only closed, well-defined rings have a strong effect in trapping the gas particles, unlike pseudo-rings. The only galaxy in our sample with a closed ring, NGC 5850, is indeed characterised by a low-amplitude centrally peaked Σ_{mol} profile that drops along the bar showing very little SF (Fig. 4), and gets enhanced again at one of the bar ends, at the ring radius.

Distinct distributions of massive SF within bars in galaxies of different morphological types have been reported by [Díaz-García et al. \(2020\)](#) using both stacking techniques (NUV and FUV) and visual classifications (in both H α and FUV).

They also find differences in the statistical distributions of star-forming and quiescent bars as a function of physical properties, such as cold gas fraction and tangential-to-radial force ratios; for a segregation in the SFR- M_\star plane in the Mapping Nearby Galaxies at APO (MaNGA; [Bundy et al. 2015](#)), see [Fraser-Mckelvie et al. \(2020\)](#). S0 galaxies tend to host SF exclusively in the circumnuclear regions, but are not studied in this work. Star-forming bars are most common among late-type galaxies ($T > 5$); [Díaz-García et al. \(2020\)](#) argue that this is a consequence of low shear given their lower central masses, which in turn may favour the gravitational collapse of GMCs and the formation of stars (e.g., [Seigar 2005](#)) (see next Sect. 6.2).

Star formation at the bar ends, but scant along the bar, is typical for early- and intermediate-type spirals ([Díaz-García et al. 2020](#)). This most likely results from the interplay of gas flow, shocks, and enhanced shear in centrally concentrated galaxies with large bar amplitudes. Even so, the authors identify quite a few galaxies ($\sim 1/3$) with $0 \leq T \leq 5$ and H II regions along bars, and a key parameter missing from their analysis is the fuel for SF. In the present paper, that purposely targets star-forming as well as quiescent bars hosted by S0/a-Sc galaxies, we show that the degree of SF in bars correlates with the mass surface density of molecular gas, in agreement with the KS law reported for disks as a whole. We detect H₂ along the bars of all probed galaxies, regardless of their T -type, very clearly in the circumnuclear regions (Fig. 4) and at the bar ends in some objects. The value of Σ_{mol} varies by almost 2 orders of magnitude between the probed galaxies.

Bar stacks by [Díaz-García et al. \(2020\)](#) show that the average FUV emission in the mid- and outer part of strong bars hosted by more than 100 S0/a-Sc galaxies is ~ 26.3 mag arcsec $^{-2}$, which corresponds to $\Sigma_{\text{SFR}} \approx 10^{-3.2} M_\odot \text{ yr}^{-1} \text{ kpc}^{-2}$ (see their Fig. 3). This value is consistent with the mean Σ_{SFR} of the 12 galaxies probed in this work, as shown in the upper panel of Fig. C.2; in the case of NGC 4548 and NGC 5850 Σ_{SFR} , as traced from FUV, is 0.5–1 orders of magnitude lower. The quiescent bars in our sample (NGC 4394, NGC 4548, and NGC 5850), all having $\Sigma_{\text{SFR}} \lesssim 10^{-2.25} M_\odot \text{ yr}^{-1} \text{ kpc}^{-2}$ when NUV is combined with W4 or W3, show low Σ_{mol} . Interestingly, these three bars clearly cluster below the main sequence (Fig. 8) and are also characterised by low specific SFRs (see lower panel of Fig. 7). Their atomic gas content is likely to be low as well (e.g., this is clear from the inspection of the VIVA H I map for NGC 4548 in [Chung et al. 2009](#)), but the analysis of H I is beyond the scope of this paper. Differences in SFEs between galaxies are discussed next.

6.2. The SFE is not always diminished in bars

In general, the SFEs in the bars of our galaxy sample (Fig. 6) are similar to those reported by [Bigiel et al. \(2011\)](#) (0.43 Gyr^{-1}) from the analysis of 30 nearby disks at a resolution of 1 kpc (see also [Bigiel et al. 2008](#)). Of the three quiescent bars that we resolve in this work, with $\Sigma_{\text{SFR}} \lesssim 10^{-2.25} M_\odot \text{ yr}^{-1} \text{ kpc}^{-2}$, NGC 4548 and NGC 5850 have a factor of ~ 2 lower SFEs, while NGC 4394 ($< \text{SFE} > = 0.64 \pm 0.07 \text{ Gyr}^{-1}$) does not. On the other hand, NGC 4535, which hosts SF along the bar, has a fairly low SFE ($< \text{SFE} > = 0.23 \pm 0.02 \text{ Gyr}^{-1}$).

Combining data from the Nobeyama 45m and ALMA 12m (with no sensitivity to diffuse gas), [Maeda et al. \(2020\)](#) study the star formation activity in NGC 1300 and report a lower SFE at the bar region than in the bar-end or spiral arms (see their Fig. A.2). They find values of SFE at the bar region that are much lower than those in [Bigiel et al. \(2011\)](#). They argue that molecular gas exists in strong bars with no clear H II regions

(Maeda et al. 2018). Likewise, based on single-dish CO observations of NGC 2903 and NGC 4303, Muraoka et al. (2016) and Yajima et al. (2019) conclude that the H₂ gas density in the bar regions is lower than that in the arms. Here, SFEs in the central regions are not systematically different than those in outer parts of the bar and at the beginning of the arms or rings. Maeda et al. (2020) argue that the presence of a large amount of diffuse molecular gas across bars makes the SFE low in appearance. However, they conclude that the bar SFE remains low even when diffuse gas is excluded, and thus other mechanisms might explain this trend. In addition, Maeda et al. (2020) find a tight correlation between the SFE and $R_{21} = I_{\text{CO}(2-1)}/I_{\text{CO}(1-0)}$. R_{21} is a proxy of the gas excitation (e.g., Koda et al. 2012), which Maeda et al. (2020) report to be low in the bar region of NGC 1300. We do not confirm these trends in the bars in our sample (Fig. F.1; see Appendix F for further discussion).

An explanation for the presence or absence of SF are the local conditions of the gas in bars. Cloud-cloud collisions, inner-cloud turbulence, or the bar tidal field, can tear apart GMCs (e.g., Hüttemeister et al. 2000; Yajima et al. 2019) and enhance the fraction of diffuse gas. In addition, as discussed by Regan & Elmegreen (1997) and Sheth et al. (2002), GMCs can be disrupted, or their formation inhibited, by diverging streamlines on the trailing side of the bar dust lanes. Sheth et al. (2000) argue, however, that this effect can be neutralised by self-gravity. Recent models by Khoperskov et al. (2018) show that SF is less efficient in the turbulent interstellar medium, and point to the effect of bars quenching SF without needing to deplete the gas. One possible explanation for the low SFE in their galaxy (or in galaxies like NGC 4548 and NGC 5850 here) is indeed that GMCs in bars are gravitationally unbound (e.g., Meidt et al. 2013; Nimori et al. 2013), but we argue that this is not the general picture for all strong bars.

Evidence exists for the inhibition of SF based on CO observations (e.g., Reynaud & Downes 1998). Fluid dynamic simulations of bars (e.g., Athanassoula 1992, 2000) predict that the highest gas density loci are also the loci of strong shocks and high shear within strong bars. Observations of H α velocity gradients suggest that shear makes SF drop, whereas shocks enhance it in general (Zurita et al. 2004). In Sect. 5 we showed that the pointings having the highest Σ_{SFR} are indeed those with the lowest shear rate parameters, calculated from the stellar contribution to the circular velocity (from Díaz-García et al. 2016b). In this paper we use a lower limit of the true shear since we neglect the contribution of dark matter, gas, and non-circular motions.

The occurrence of SF at the bar ends is witnessed in simulations as well (e.g., Renaud et al. 2015). It may result from a combination of kiloparsec-scale dynamics (gas flows) and parsec-scale turbulence and clouds collisions, under a low shear, as discussed in Fraser-Mckelvie et al. (2020). The same physics explain why H II regions are preferentially located on the leading side of the bars (e.g., Athanassoula 1992; Rozas et al. 2000; Sheth et al. 2002; Popping et al. 2010; Neumann et al. 2019; Díaz-García et al. 2020; Fraser-Mckelvie et al. 2020): it is likely a consequence of shear and turbulence forces inhibiting SF everywhere but on the bar leading side (Emsellem et al. 2015; Renaud et al. 2015; Fraser-Mckelvie et al. 2020) and at the bar ends.

A possible connection between bar strength and the occurrence of SF has been discussed in the past (e.g., Jogee et al. 2002), given the expectation that weaker bars have weaker shocks and shear (Athanassoula 1992). However, to our knowledge this is not supported by any recent observational work with large unbiased samples. In some galaxies H II regions

are indeed found in bars where strong shocks and shear are also detected (e.g., Martin & Friedli 1997; Sheth et al. 2002; Zurita et al. 2004; Zurita & Pérez 2008). We searched for the existence of correlations of the SFE with various structural parameters, such as bar-to-total and bulge-to-total flux ratios (from Kim et al. 2014; Salo et al. 2015), bar strength (e.g., normalised $m = 2$ Fourier amplitude or ellipticity, from Díaz-García et al. 2016b), or even tangential-to-radial forces evaluated locally (e.g., Díaz-García et al. 2019b) within the area covered by the IRAM-30 m pointings; we found no evidence of any correlations (plots not shown here). This implies that either the range of bar parameters covered is too narrow (our sample only contains strong bars, Fig. 2) or that other, more fundamental, global dynamical and/or neutral gas properties can explain the differences in SFE.

We find an interesting segregation as a function of the total stellar mass of the host galaxy (Fig. 7): galaxies with highest M_* tend to host bars with low SFE. Even though we have used the largest sample to date for the stated scientific goals, the robustness of the analysis is still limited by sample size. Thus, we encourage this result to be verified elsewhere with larger surveys, such as EDGE-CALIFA (Bolatto et al. 2017; Sánchez et al. 2021) or Physics at High Angular Resolution in Nearby GalaxieS (PHANGS; e.g., Rosolowsky et al. 2021; Anand et al. 2021; Lee et al. 2021). The analysis of SF scaling relations at ~ 100 pc from PHANGS was released by Pessa et al. (2021) during the review process of this paper and will be analysed elsewhere.

Verley et al. (2010) conclude that for M 33 the stellar disk, which is gravitationally dominant with respect to the gaseous disk, plays a major role in driving the SFR, and the usual phase shift between gas and stellar components may also play a role to explain the aforementioned tendency. Whether the segregation of the KS law as a function of M_* hints at the SF suppression by large-amplitude bars (typically hosted by massive galaxies; Díaz-García et al. 2016a) or points to more fundamental dynamical differences remains to be confirmed. We note that Leroy et al. (2013, see their Sect. 4.1.2) already discussed a weak correlation relating gas depletion times to galaxy masses and average surface densities (see their Fig. 7; see also Saintonge et al. 2011, 2012).

To sum up, H α and far- and near-UV maps of disk galaxies show a massive SF along some bars, whereas there is a lack of SF in others. Among early- and intermediate-type galaxies, $\sim 2/3$ lack H II regions along the bar, but $\sim 1/3$ do not (Díaz-García et al. 2020). We argue, for S0/a-Sc galaxies, that the degree of SF in bars is not primarily controlled by the SFE. Similarly, George et al. (2020) recently presented a multi- λ analysis of bars in NGC 3351, NGC 4579, and NGC 4725, and concludes that bars that are devoid of SF (as traced from UV) are also devoid of molecular and neutral hydrogen. Our observations are not easy to reconcile with the theoretical expectation of a lower SFE in the central parts of gas-rich barred galaxies, compared to their unbarred counterparts, or a flattening of the slope of the KS relation in the bar region (Khoperskov et al. 2018); we note, however, that our sample is not limited to gas-rich galaxies (Sect. 2). We conclude that gas depletion times are not always high along strong bars, either in the central, mid-, or end parts. Strong bars with lowest SFEs are typically hosted by massive galaxies ($M_* \gtrsim 10^{10.7} M_\odot$).

6.3. Environmental versus internal quenching

We have shown that NGC 4394, NGC 4548, and NGC 5850 host bars with remarkably low values of Σ_{SFR} and sSFR. What

physical mechanisms, other than those discussed so far, could have caused the quenching of SF in their bars?

Given their high Σ_{mol}^A values (lower right panel of Fig. 7), environmental effects are likely responsible for the low level of star formation in these systems. NGC 5850 is perturbed, as shown by Higdon et al. (1998), due to the encounter with the nearby massive elliptical NGC 5846. Using interferometric IRAM observations, Leon et al. (2000) report CO emission from the nuclear ring, whose distribution is asymmetric due to the interaction. Likewise, NGC 4394 and NGC 4548 belong to the Virgo cluster. NGC 4548 shows signatures of a past interaction, such as an inner polar ring (Sil'chenko 2002), and is perturbed and warped (Vollmer et al. 1999).

On the other hand, other internal mechanisms, besides those discussed in Sect. 6.2, may account for the low SFE in NGC 4548 or NGC 5850. According to numerical models by Martig et al. (2009), given that SF occurs in gravitationally unstable gas disks, massive bulges can morphologically quench SF in galaxies. The stellar mass of the bulges hosted by NGC 4548 and NGC 5850 are $10^{9.7} M_{\odot}$ and $10^{10.1} M_{\odot}$, respectively. This means that they are indeed two of the galaxies with the largest bulges in our sample (see the lower left panel of Figs. 7 and C.3); however, we have checked that no correlation exists between the residuals of the KS law and bulge-to-total flux ratios. We conclude that both morphological and environmental quenching are likely responsible for the low Σ_{SFR} and SFE in NGC 4548 and NGC 5850.

7. Summary and conclusions

The principal aim of this work was to study the processes that enhance or inhibit star formation within strong stellar bars of spiral galaxies. Bars are excellent laboratories for conducting exploratory research on SF as they are characterised by extreme physical conditions, such as strong shear, shocks, non-circular motions, and magnetic fields. We selected a sample of 12 nearby, practically face-on, strongly barred galaxies with total stellar masses $\log_{10}(M_{\star}/M_{\odot}) \in [10.2, 11]$, Hubble types $T \in [0, 5]$, and different degrees of star formation in bars.

In order to trace the distribution of molecular gas within bars we obtained CO(1–0) spectra (21.5'' FWHM) and CO(2–1) spectra (10.75'' FWHM) with the IRAM-30 m radiotelescope, performing several pointings along their major axes. At the distance to our galaxies ($\lesssim 40$ Mpc), the resolution elements associated with the IRAM pointings were in the range 1.7–3.9 kpc for CO(1–0) and 0.85–1.95 kpc for CO(2–1). The CO spectra are made available in electronic form at the CDS linked to this publication. CO emission was detected in the full set of galaxies. We computed the molecular gas surface density (Σ_{mol}) from the velocity-integrated line intensity of the CO(1–0) spectra. A remarkable central enhancement of Σ_{mol} was observed in the whole sample, while four galaxies show secondary peaks at the bar ends.

To trace the star formation (SF), we used GALEX near- and far-UV and WISE 12 and 22 μm photometry. We performed aperture photometry centred on the CO pointings and computed the SFRs. We find similar distributions of the surface density of SF rate (Σ_{SFR}) in bars; the 12 galaxies host central starbursts, which is not surprising as strong bar torques are known to actively funnel the gas towards the central regions (e.g., Sellwood & Wilkinson 1993). Three quiescent bars in our sample (NGC 4394, NGC 4548, and NGC 5850) have $\Sigma_{\text{SFR}} \lesssim 10^{-2.25} M_{\odot} \text{yr}^{-1} \text{kpc}^{-2}$ as quantified from the combination of NUV and WISE 3 and WISE 4 emission along the bar major

axis; interestingly, these are the galaxies in the densest environments (highest surface density to the third-nearest neighbour galaxy).

Σ_{SFR} and Σ_{mol} are tightly correlated within bars. This means that the SF activity in bars of strongly barred galaxies is primarily controlled by the content of molecular gas. We calculated the star formation efficiencies ($\text{SFE} = \Sigma_{\text{SFR}}/\Sigma_{\text{mol}}$). The SFEs within bars vary between different galaxies by up to an order of magnitude, spanning $0.1 \lesssim \text{SFE/Gyr}^{-1} \lesssim 1.8$ when measured from NUV and WISE 3 combined, but are roughly constant along the bar major axes.

The derived SFRs and SFEs follow the typical values found in disks of spirals galaxies (e.g., Bigiel et al. 2011). We do not find remarkable differences in the KS law in the central regions of barred and non-barred galaxies, as predicted from numerical models by Khoperskov et al. (2018) for gas-rich galaxies. However, in some galaxies the SFE within bars can be a factor of ~ 2 lower than the average in spirals. Interestingly, the two with lowest SFE and Σ_{SFR} (NGC 4548 and NGC 5850) are also those hosting massive stellar bulges, while having signatures of recent interactions; both environmental and internal quenching are plausible explanations for their low level of star formation.

In general, whether or not a strong bar hosts star-forming regions is not only determined by the degree of SFE, as has been previously claimed, and thus other physical factors must also come into play. Interestingly, the higher the total stellar mass of the galaxies, the lower the SFE within their bars. The latter needs to be verified with larger samples that include faint systems as well as weak bars as it might link the SF suppression to large-amplitude bars (typically hosted by massive galaxies, e.g., Díaz-García et al. 2016a).

The novelty of this work is that it presents a blind study of the molecular gas in bars for galaxies with different star formation distributions and properties, with the aid of multi- λ data. In most previous works the emphasis was placed on objects for which molecular gas was present, or objects harbouring a lot of star formation in their bars. We have shown that the star formation efficiency is not uniformly inhibited in strong bars using a fairly large sample of 12 galaxies.

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Appendix A: IRAM-30 m pointings and CO spectra

We present the CO(1-0) and CO(2-1) spectra used in this work for the 12 galaxies in our sample (Figs. A.1–A.12). The upper and central panels correspond to the WISE 1 and GALEX NUV images. The bottom panel shows the spectra in every beam along bars, with every subplot located according to the RA and DEC offsets of the pointings. Velocity and intensity ranges are indicated in the lower right corner of the plots. In the images we show in blue the regions where the CO spectra were taken, which were chosen based on the visual estimate of the bar length and position angle. These are in good agreement with the values reported by Díaz-García et al. (2016b).

We also show, with a red ellipse, the bar isophote measured by Herrera-Endoqui et al. (2015) via ellipse fitting (see Sect. 2.1), which is known to slightly underestimate the true size of bars (Wozniak & Pierce 1991; Laurikainen & Salo 2002; Erwin & Sparke 2003; Michel-Dansac & Wozniak 2006; Díaz-García et al. 2016b). Visual analysis and ellipse fitting yield similar measurements, the only exception being NGC 5383 (Fig. A.7) where the presence of H II regions in the spiral arms at the bar end make the ellipse fitting output less reliable. Uncertainties on bar sizes are not a big issue for the current analysis, given fairly large FWHM ($21.5''$) of the apertures where spectra were taken.

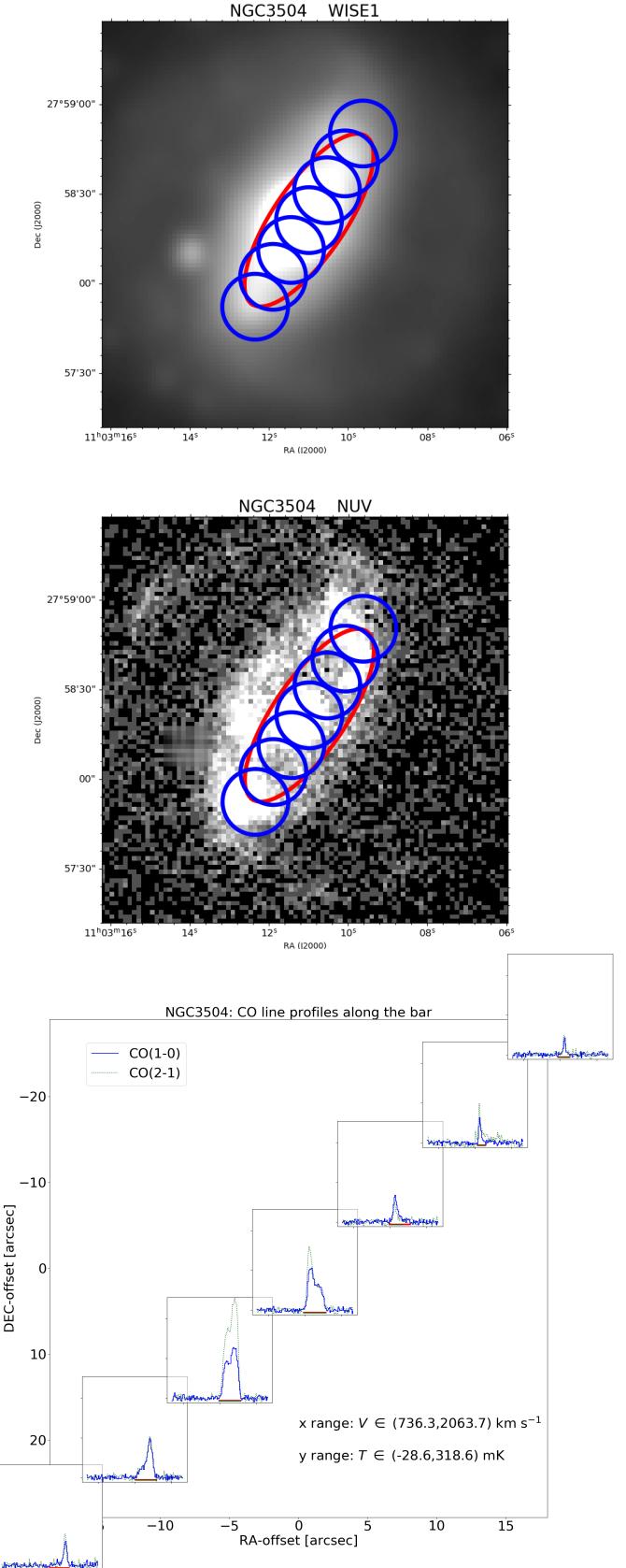
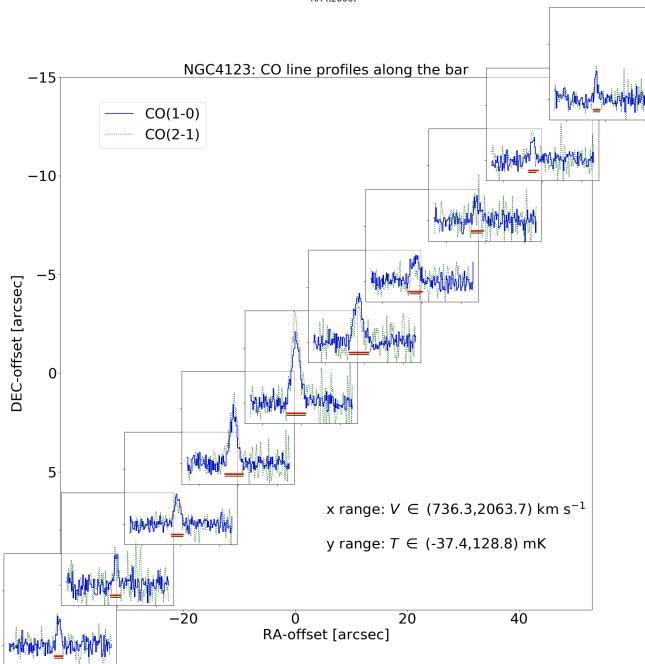
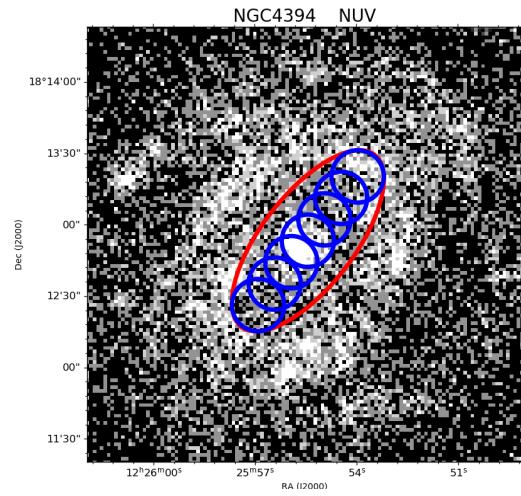
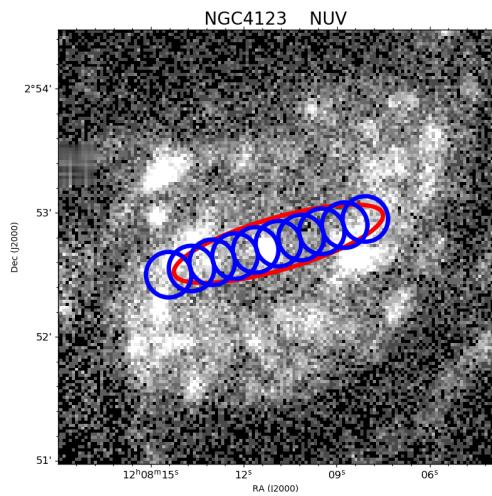
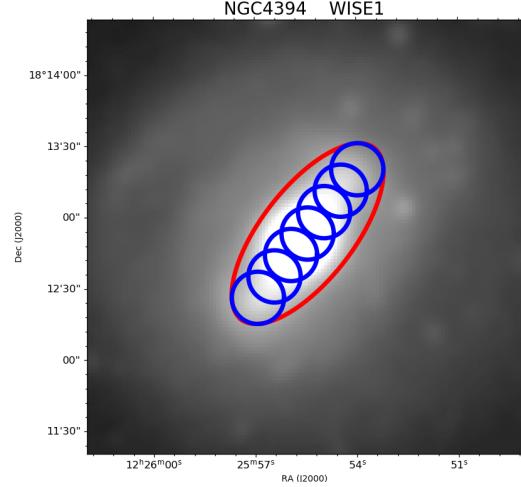
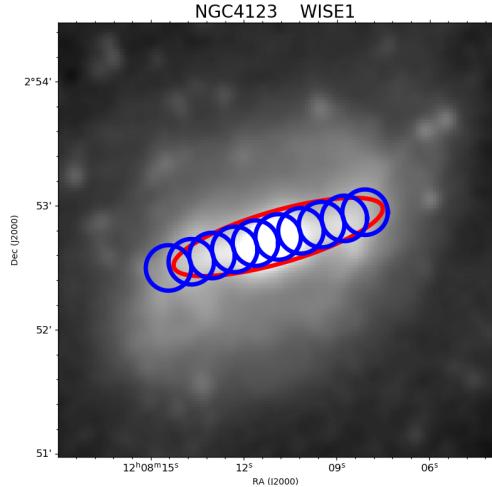
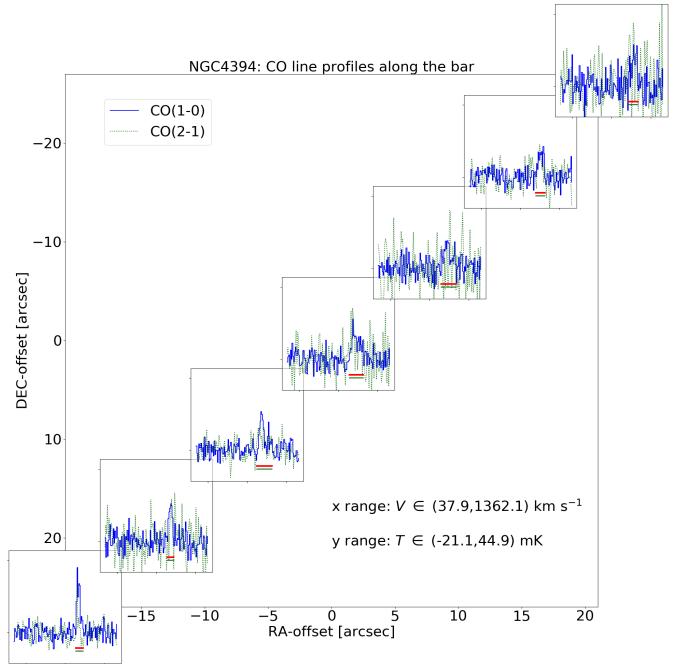


Fig. A.1. As in Fig. 3, but for NGC 3504.

**Fig. A.2.** As in Fig. 3, but for NGC 4123.**Fig. A.3.** As in Fig. 3, but for NGC 4394.

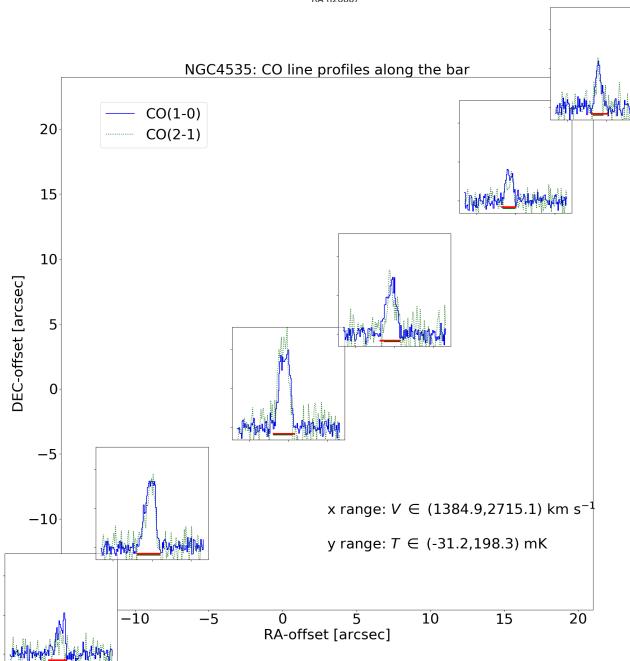
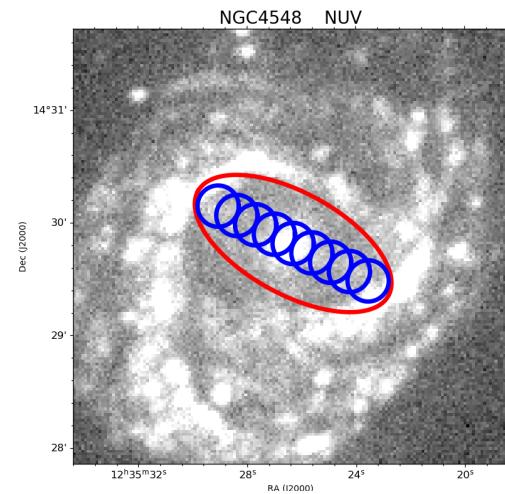
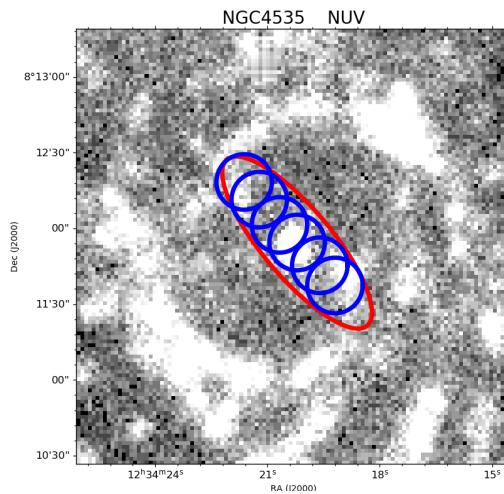
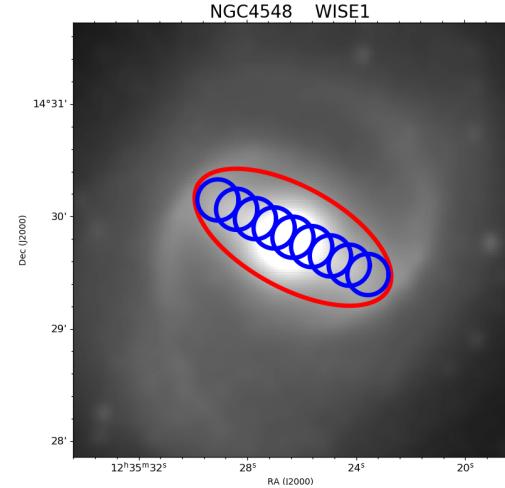
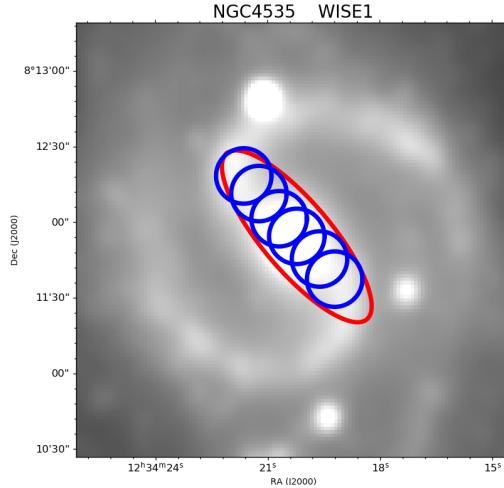


Fig. A.4. As in Fig. 3, but for NGC 4535.

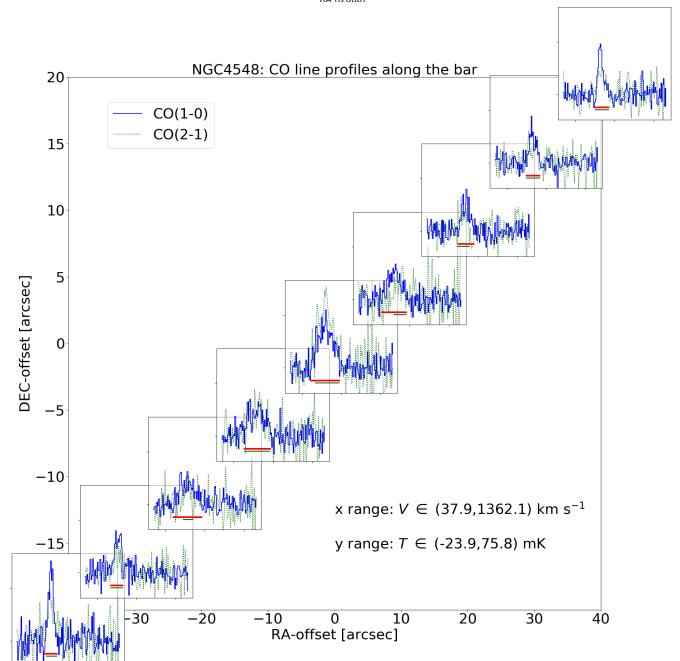


Fig. A.5. As in Fig. 3, but for NGC 4548.

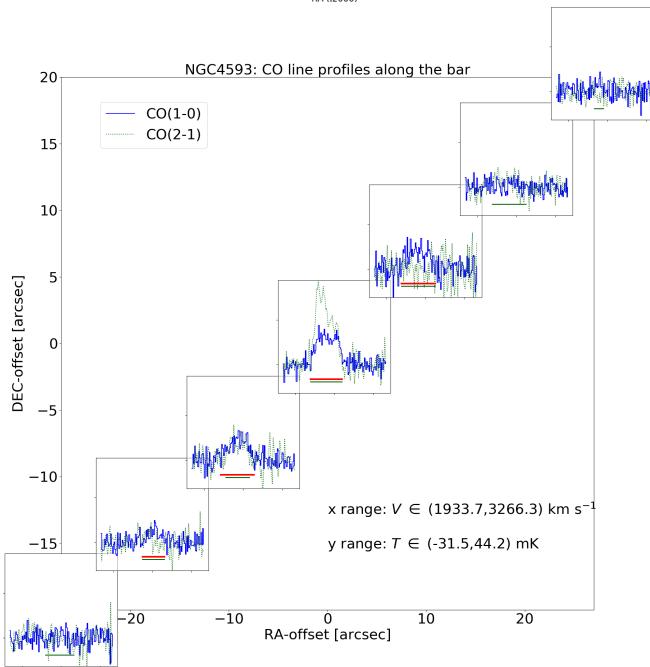
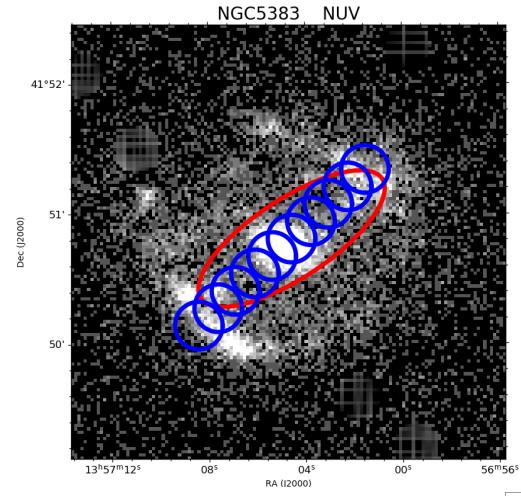
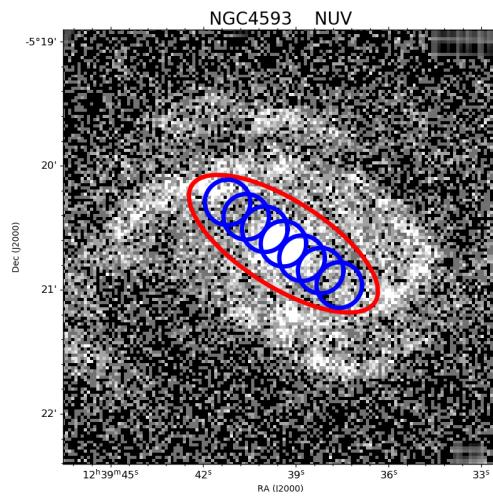
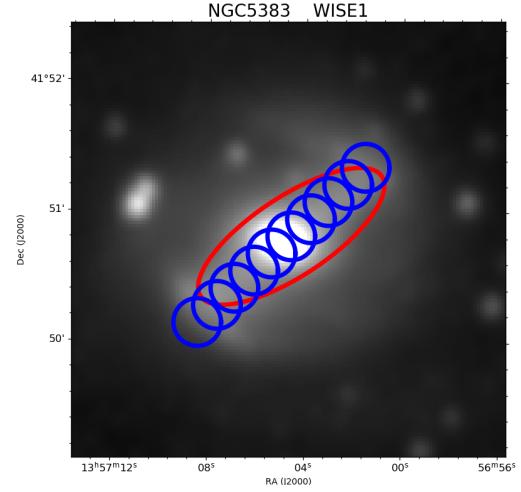
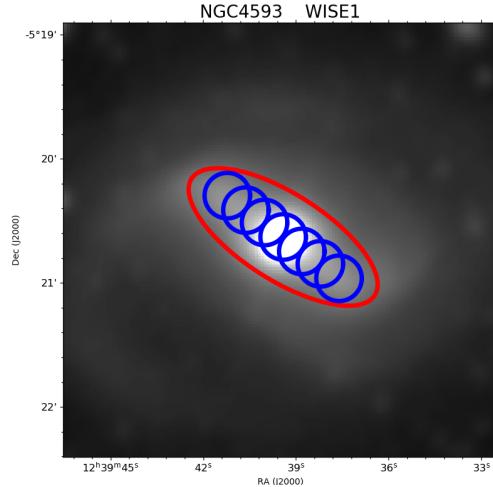


Fig. A.6. As in Fig. 3, but for NGC 4593.

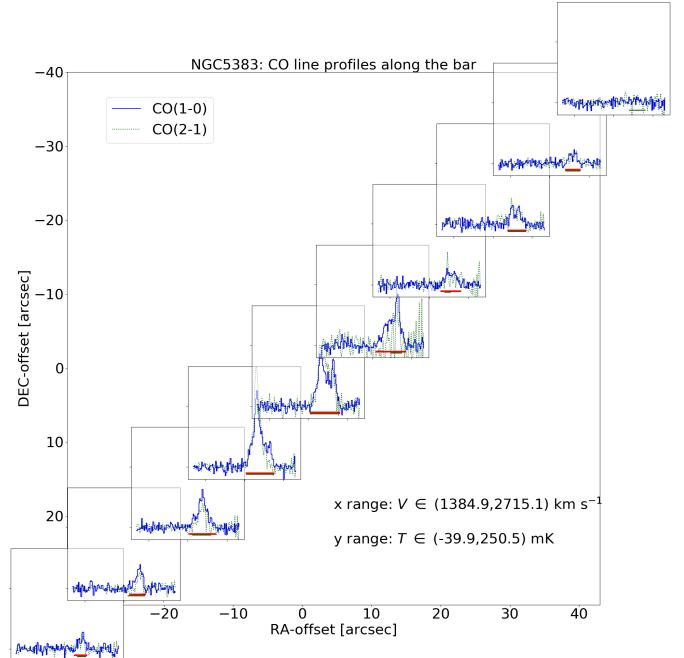


Fig. A.7. As in Fig. 3, but for NGC 5383.

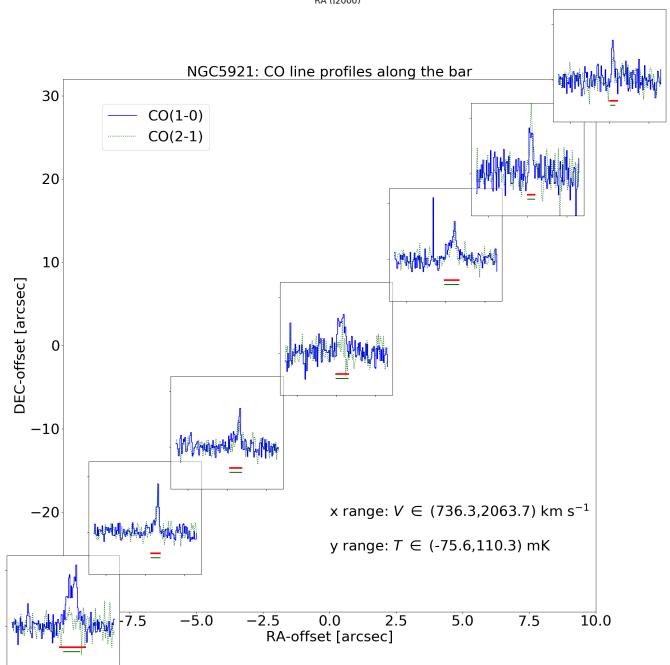
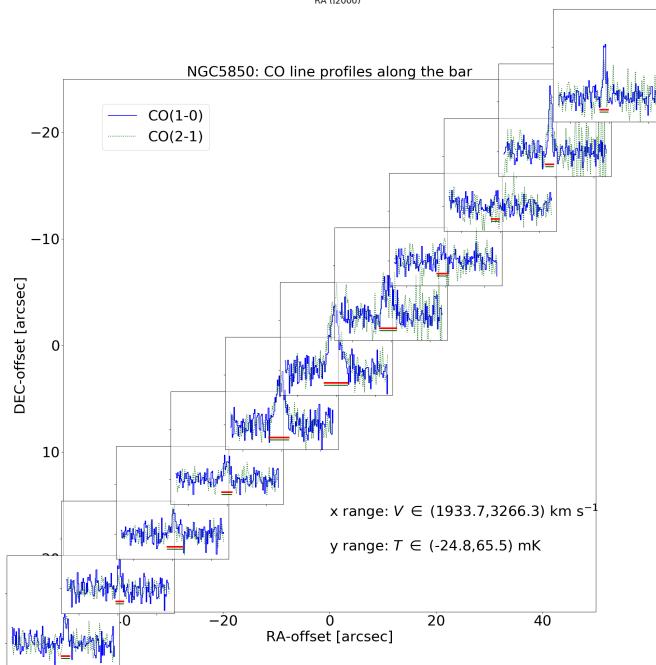
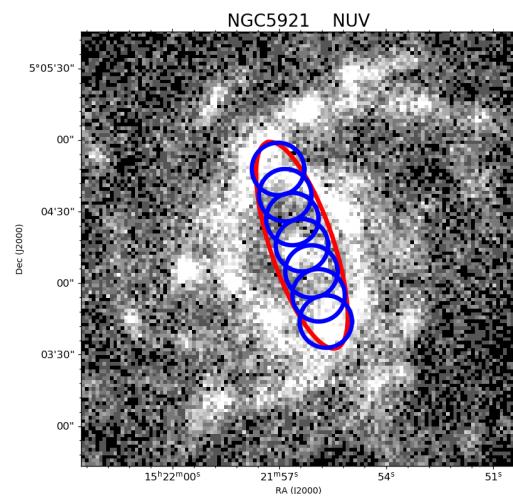
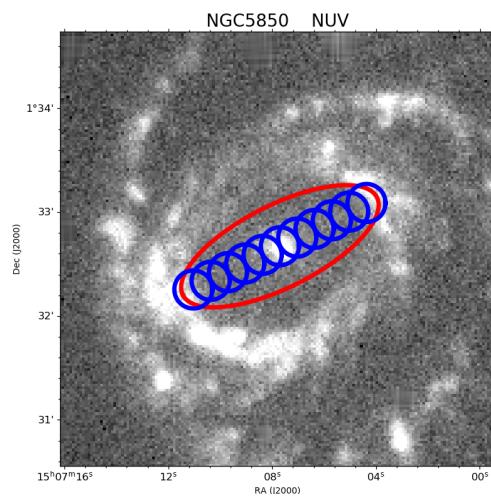
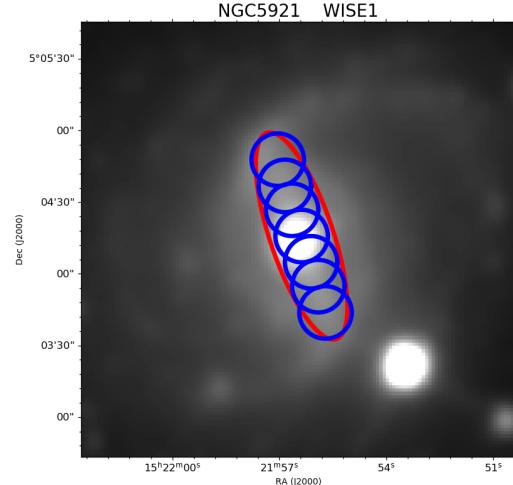
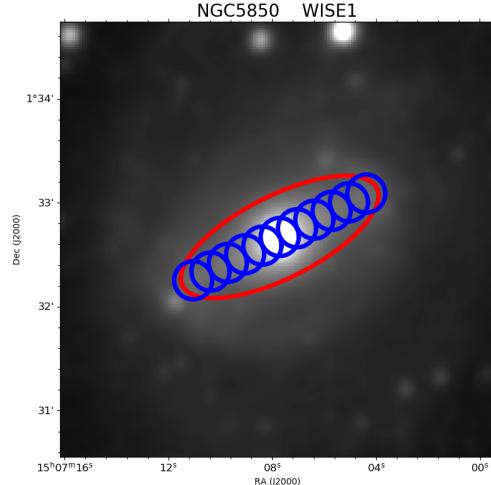
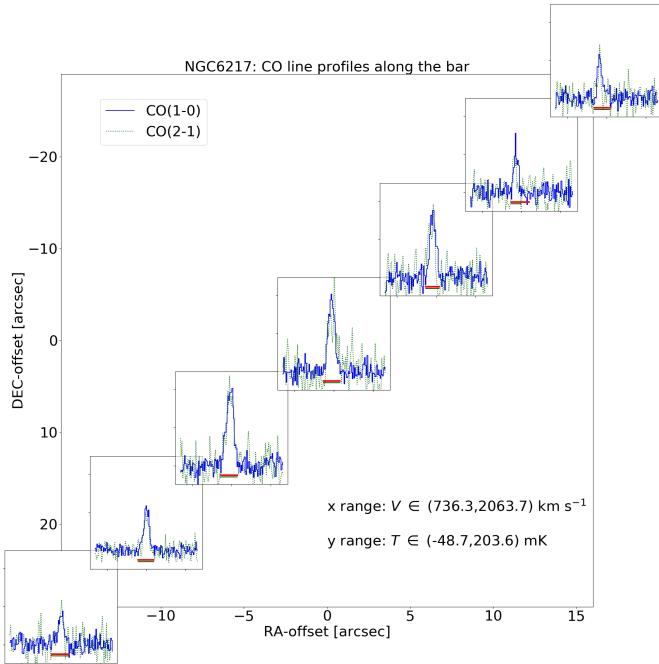
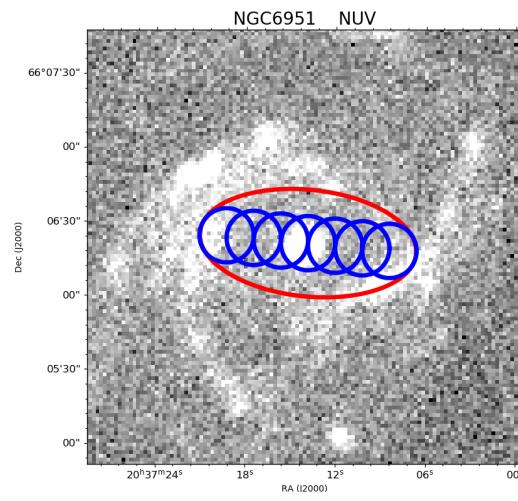
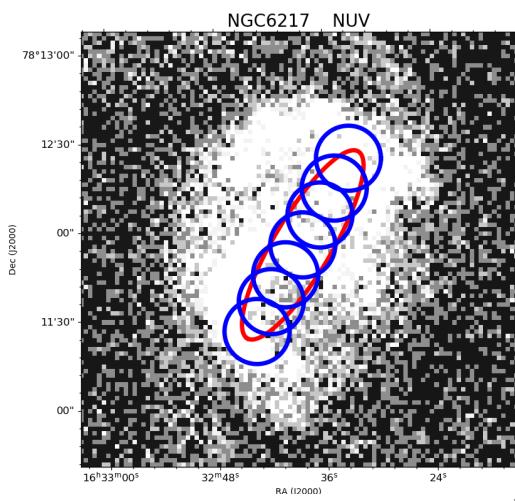
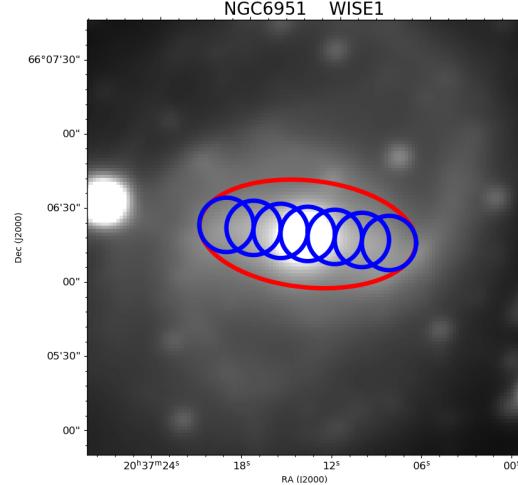
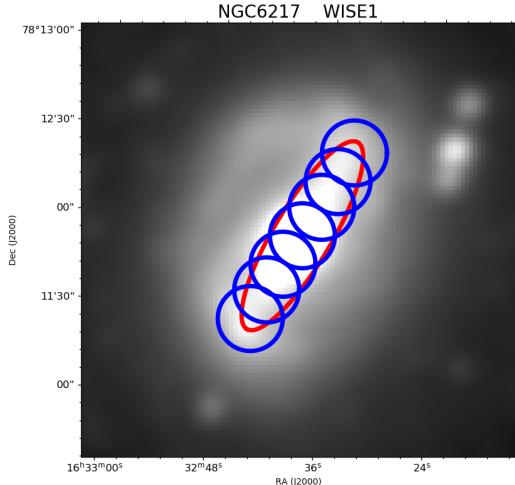
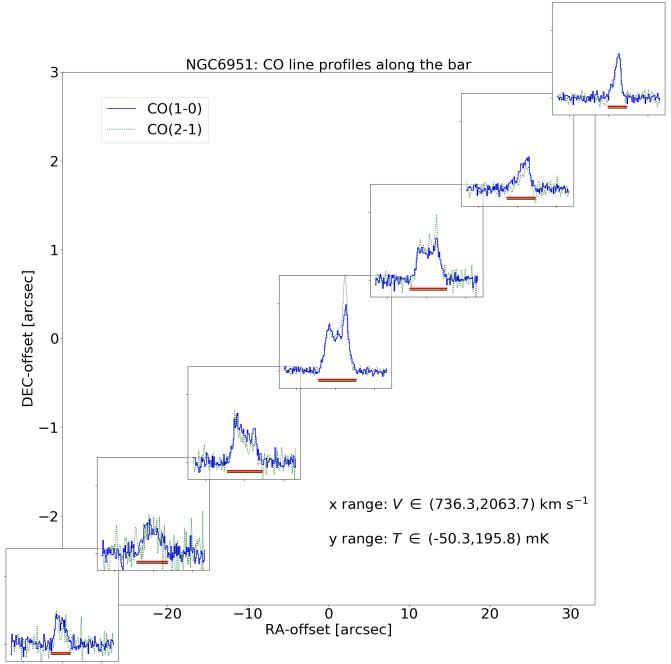


Fig. A.8. As in Fig. 3, but for NGC 5850.

Fig. A.9. As in Fig. 3, but for NGC 5921.

**Fig. A.10.** As in Fig. 3, but for NGC 6217.**Fig. A.11.** As in Fig. 3, but for NGC 6951.

Appendix B: Velocity-integrated intensities

The following data are listed in Table B.1:

- Column 1: Galaxy name;
- Column 2: Right ascension of centre of the IRAM-30 m pointing (in degrees);
- Column 3: Declination of centre of the IRAM-30 m pointing (in degrees);
- Column 4: Velocity-integrated intensity $I_{\text{CO}(1-0)}$ (in K km s^{-1}) and uncertainty on spectrum integration;
- Column 5: Root mean square (rms) noise (in mK) associated with each CO(1-0) spectrum at a velocity resolution of $\delta V = 10.5 \text{ km s}^{-1}$;
- Column 6: Zero-level velocity width of the CO(1-0) line (in km s^{-1}) of each spectrum (ΔV);
- Column 7: Velocity-integrated intensity $I_{\text{CO}(2-1)}$ (in K km s^{-1}) and its error;
- Column 8: Root mean square (rms) noise (in mK) associated with each CO(2-1) spectrum at a velocity resolution of $\delta V = 10.5 \text{ km s}^{-1}$;
- Column 9: Mass of molecular gas (in M_{\odot}) and its error (Eq. 7);
- Column 10: Surface density of molecular gas mass (in $M_{\odot} \text{ pc}^{-2}$) and its error (Eq. 8).

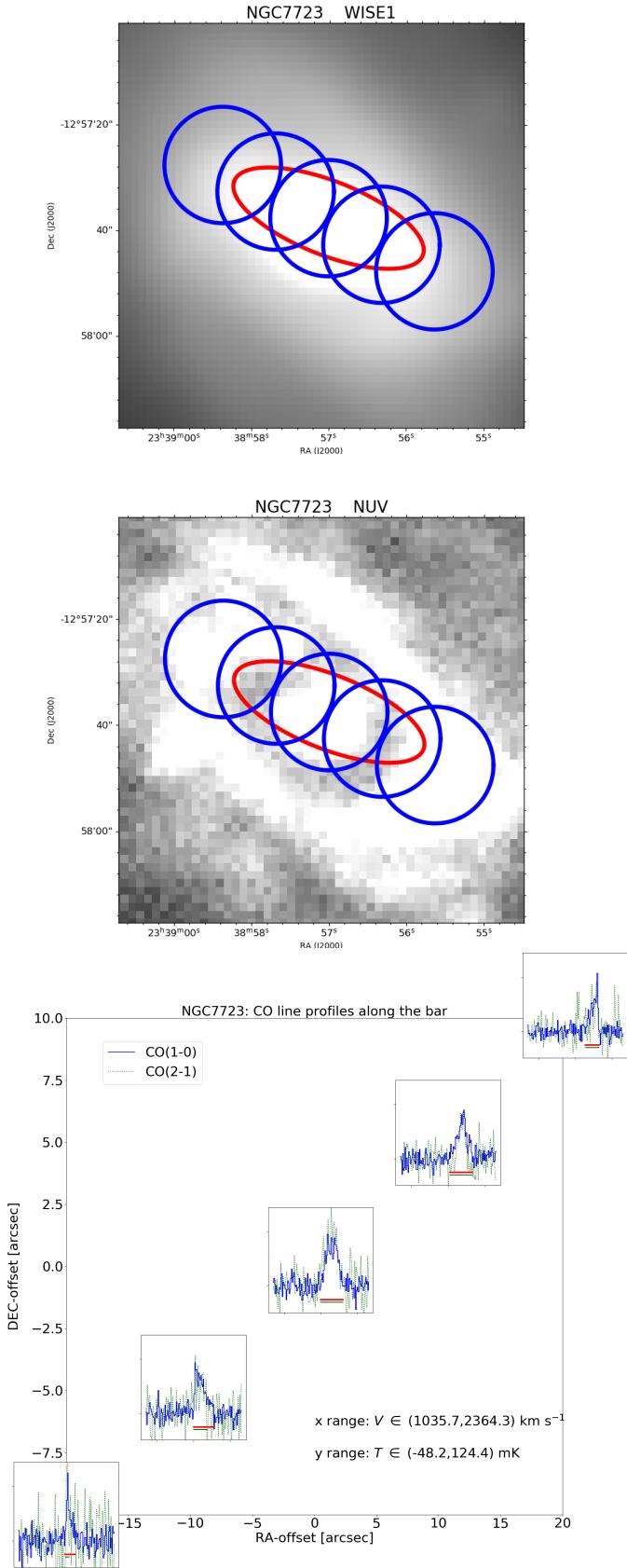


Fig. A.12. As in Fig. 3, but for NGC 7723.

Table B.1. CO integrated spectra and molecular gas.

Galaxy	RA ($^{\circ}$) (1)	DEC ($^{\circ}$) (2)	$I_{\text{CO}(1-0)}$ (K km s $^{-1}$) (4)	$\text{rms}_{\text{CO}(1-0)}$ (mK) (5)	ΔV (km s $^{-1}$) (6)	$I_{\text{CO}(2-1)}$ (K km s $^{-1}$) (7)	$\text{rms}_{\text{CO}(2-1)}$ (mK) (8)	M_{mol} (M_{\odot}) (9)	Σ_{mol} (M_{\odot} pc $^{-2}$) (10)
NGC3504	165.79059	27.98028	10.03 \pm 0.59	11.34	260	8.77 \pm 0.83	15.92	8.32 \pm 1.25	42.58 \pm 1.84
	165.79248	27.97750	26.11 \pm 0.49	9.06	283	26.62 \pm 0.77	14.10	8.74 \pm 1.73	110.81 \pm 1.54
	165.79436	27.97500	55.88 \pm 0.58	10.43	293	107.10 \pm 0.68	12.32	9.07 \pm 1.99	237.18 \pm 1.80
	165.79625	27.97222	42.30 \pm 0.53	9.41	300	52.53 \pm 0.68	12.22	8.95 \pm 1.91	179.51 \pm 1.64
	165.79814	27.96944	12.97 \pm 0.52	9.54	281	6.57 \pm 0.76	14.02	8.43 \pm 1.42	55.06 \pm 1.61
	165.80002	27.96694	5.70 \pm 0.32	9.54	105	10.17 \pm 0.85	25.66	8.08 \pm 1.28	24.20 \pm 0.99
	165.80191	27.96417	4.28 \pm 0.37	9.23	155	4.82 \pm 0.52	12.92	7.95 \pm 1.10	18.15 \pm 1.16
NGC4123	182.03407	2.88211	2.91 \pm 0.32	9.86	95	2.96 \pm 0.49	15.46	7.71 \pm 1.00	8.65 \pm 0.70
	182.03685	2.88128	2.98 \pm 0.30	11.85	114	< 3.23	23.52	7.72 \pm 1.04	8.87 \pm 0.66
	182.03990	2.88044	4.49 \pm 0.38	7.95	144	3.57 \pm 0.48	12.34	7.90 \pm 1.10	13.36 \pm 0.84
	182.04269	2.87961	9.83 \pm 0.50	8.83	226	13.08 \pm 0.92	18.91	8.24 \pm 1.32	29.22 \pm 1.09
	182.04575	2.87878	14.23 \pm 0.43	10.21	230	18.70 \pm 1.35	27.63	8.40 \pm 1.54	42.31 \pm 0.93
	182.04881	2.87794	10.74 \pm 0.48	8.77	241	7.34 \pm 1.12	22.23	8.27 \pm 1.37	31.94 \pm 1.04
	182.05159	2.87711	4.94 \pm 0.38	9.87	175	2.86 \pm 0.70	16.41	7.94 \pm 1.15	14.70 \pm 0.83
	182.05466	2.87628	3.20 \pm 0.31	10.27	146	< 2.89	21.05	7.75 \pm 1.05	9.53 \pm 0.68
	182.05743	2.87544	2.26 \pm 0.41	8.71	115	2.93 \pm 0.60	17.42	7.60 \pm 0.81	6.73 \pm 0.90
	182.06049	2.87461	2.01 \pm 0.29	9.51	80	< 2.24	16.31	7.55 \pm 0.91	5.99 \pm 0.62
NGC4394	186.47528	18.22185	1.86 \pm 0.16	4.54	92	< 0.75	5.45	7.38 \pm 1.11	6.99 \pm 0.44
	186.47733	18.21935	1.27 \pm 0.17	5.90	79	< 1.68	12.27	7.22 \pm 0.93	4.77 \pm 0.47
	186.47937	18.21685	1.96 \pm 0.22	4.33	191	< 0.87	6.33	7.40 \pm 1.00	7.35 \pm 0.60
	186.48141	18.21435	2.05 \pm 0.34	5.09	174	< 1.57	11.48	7.43 \pm 0.84	7.71 \pm 0.95
	186.48347	18.21185	1.91 \pm 0.19	5.96	189	< 1.97	14.35	7.39 \pm 1.04	7.17 \pm 0.53
	186.48552	18.20935	1.32 \pm 0.20	4.88	114	< 1.02	7.47	7.23 \pm 0.87	4.97 \pm 0.56
	186.48756	18.20685	1.90 \pm 0.16	8.09	120	< 2.17	15.82	7.39 \pm 1.11	7.14 \pm 0.44
NGC4535	188.58041	8.19330	10.21 \pm 0.58	9.68	224	< 2.85	20.81	8.15 \pm 1.27	40.08 \pm 1.68
	188.58211	8.19553	26.85 \pm 0.60	10.99	281	23.22 \pm 1.38	25.44	8.57 \pm 1.66	105.35 \pm 1.72
	188.58463	8.19803	30.66 \pm 0.51	12.04	262	34.57 \pm 1.83	34.91	8.63 \pm 1.79	120.34 \pm 1.46
	188.58659	8.19997	19.38 \pm 0.60	11.73	249	14.95 \pm 1.77	34.75	8.43 \pm 1.52	76.04 \pm 1.73
	188.58884	8.20275	7.84 \pm 0.48	11.36	169	6.16 \pm 0.76	18.11	8.03 \pm 1.24	30.78 \pm 1.38
	188.59052	8.20469	9.57 \pm 0.56	12.22	204	7.41 \pm 1.18	25.59	8.12 \pm 1.25	37.56 \pm 1.63
NGC4548	188.84856	14.49091	5.24 \pm 0.31	9.89	154	3.81 \pm 0.83	20.63	7.75 \pm 1.26	17.70 \pm 0.76
	188.85143	14.49230	3.28 \pm 0.30	7.53	146	1.85 \pm 0.41	10.41	7.55 \pm 1.08	11.08 \pm 0.74
	188.85429	14.49369	4.89 \pm 0.41	6.43	356	< 1.83	13.37	7.72 \pm 1.11	16.52 \pm 1.03
	188.85716	14.49508	5.97 \pm 0.35	8.14	323	5.56 \pm 0.87	14.98	7.81 \pm 1.25	20.19 \pm 0.88
	188.86003	14.49647	9.40 \pm 0.42	7.64	360	13.01 \pm 1.14	18.66	8.01 \pm 1.37	31.77 \pm 1.05
	188.86290	14.49786	5.17 \pm 0.46	7.59	308	< 2.52	18.35	7.75 \pm 1.09	17.49 \pm 1.15
	188.86577	14.49925	3.30 \pm 0.29	6.78	190	2.12 \pm 0.46	10.30	7.55 \pm 1.10	11.15 \pm 0.71
	188.86864	14.50064	2.88 \pm 0.31	6.06	161	2.12 \pm 0.46	11.17	7.50 \pm 1.01	9.75 \pm 0.77
	188.87151	14.50202	3.42 \pm 0.41	6.87	166	2.14 \pm 0.43	10.40	7.57 \pm 0.97	11.56 \pm 1.02
NGC4593	189.90707	-5.34977	< 0.93	8.52	600	< 1.68	12.23	< 7.68 \pm 0.44	< 3.36 \pm 1.43
	189.90958	-5.34783	2.53 \pm 0.59	7.19	283	< 1.33	9.70	8.11 \pm 0.72	9.20 \pm 1.58
	189.91209	-5.34616	7.11 \pm 0.47	7.87	424	5.45 \pm 0.76	11.46	8.56 \pm 1.21	25.81 \pm 1.26
	189.91460	-5.34421	10.49 \pm 0.53	6.75	405	20.15 \pm 0.67	10.23	8.73 \pm 1.32	38.10 \pm 1.42
	189.91711	-5.34227	7.14 \pm 0.53	10.88	433	< 2.55	18.60	8.56 \pm 1.16	25.92 \pm 1.41
	189.91962	-5.34060	< 0.99	7.08	183	< 1.22	8.87	< 7.70 \pm 0.62	< 3.58 \pm 0.84
	189.92213	-5.33866	< 1.17	8.17	118	< 1.22	8.91	< 7.78 \pm 0.69	< 4.25 \pm 0.80
NGC5383	209.25691	41.85528	3.61 \pm 0.46	10.69	147	2.34 \pm 0.39	10.01	8.35 \pm 0.94	12.90 \pm 1.22
	209.25989	41.85305	6.62 \pm 0.51	10.40	192	2.68 \pm 0.43	9.64	8.61 \pm 1.14	23.70 \pm 1.35
	209.26324	41.85083	14.92 \pm 0.60	9.51	322	7.83 \pm 0.80	13.77	8.97 \pm 1.41	53.39 \pm 1.57
	209.26622	41.84861	32.24 \pm 0.70	11.48	347	34.17 \pm 0.73	12.19	9.30 \pm 1.67	115.36 \pm 1.83
	209.26958	41.84639	31.96 \pm 0.65	11.86	369	33.93 \pm 1.04	16.79	9.30 \pm 1.70	114.34 \pm 1.72
	209.27293	41.84417	23.75 \pm 0.67	11.46	373	7.94 \pm 2.76	44.22	9.17 \pm 1.56	84.99 \pm 1.78
	209.27592	41.84195	6.38 \pm 0.58	10.31	247	< 2.49	18.18	8.60 \pm 1.08	22.81 \pm 1.53
	209.27928	41.83972	7.17 \pm 0.46	11.58	223	7.49 \pm 0.68	14.08	8.65 \pm 1.22	25.64 \pm 1.21
	209.28226	41.83750	3.60 \pm 0.45	10.52	180	2.63 \pm 0.39	8.94	8.35 \pm 0.95	12.87 \pm 1.19
	209.28561	41.83528	< 1.47	10.81	188	< 2.03	14.78	< 7.96 \pm 0.61	< 5.25 \pm 1.25

Table B.1. continued.

Galaxy	RA ($^{\circ}$)	DEC ($^{\circ}$)	$I_{\text{CO}(1-0)}$ (K km s $^{-1}$)	rms $\text{CO}(1-0)$ (mK)	ΔV (km s $^{-1}$)	$I_{\text{CO}(2-1)}$ (K km s $^{-1}$)	rms $\text{CO}(2-1)$ (mK)	M_{mol} (M_{\odot})	Σ_{mol} (M_{\odot} pc $^{-2}$)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
NGC5850	226.76862	1.55108	1.49 \pm 0.29	7.45	102	1.52 \pm 0.25	7.52	7.54 \pm 0.79	5.25 \pm 0.74
	226.77139	1.54969	1.45 \pm 0.29	7.88	93	< 0.76	5.55	7.53 \pm 0.78	5.11 \pm 0.74
	226.77417	1.54830	1.81 \pm 0.25	6.95	194	< 0.84	6.15	7.62 \pm 0.92	6.39 \pm 0.64
	226.77696	1.54691	1.66 \pm 0.25	5.51	125	0.97 \pm 0.24	6.63	7.58 \pm 0.88	5.84 \pm 0.65
	226.77974	1.54552	6.19 \pm 0.40	9.18	247	6.13 \pm 0.44	8.69	8.16 \pm 1.22	21.81 \pm 1.04
	226.78252	1.54413	8.89 \pm 0.41	8.78	294	11.66 \pm 0.52	9.44	8.31 \pm 1.35	31.30 \pm 1.07
	226.78529	1.54274	4.51 \pm 0.37	7.96	212	3.76 \pm 0.73	15.43	8.02 \pm 1.12	15.89 \pm 0.97
	226.78807	1.54136	1.02 \pm 0.24	6.44	133	0.91 \pm 0.37	9.82	7.37 \pm 0.72	3.59 \pm 0.62
	226.79085	1.53997	0.74 \pm 0.22	6.73	98	< 1.40	10.18	7.23 \pm 0.65	2.59 \pm 0.56
	226.79362	1.53858	2.80 \pm 0.22	6.87	98	< 2.77	20.18	7.81 \pm 1.14	9.88 \pm 0.57
	226.79640	1.53719	2.37 \pm 0.20	6.39	98	< 1.59	11.57	7.74 \pm 1.10	8.36 \pm 0.53
NGC5921	230.48247	5.06178	13.62 \pm 0.90	12.00	321	< 2.28	16.63	8.38 \pm 1.21	49.46 \pm 2.41
	230.48331	5.06483	3.14 \pm 0.31	7.90	111	2.77 \pm 0.33	9.55	7.74 \pm 1.05	11.38 \pm 0.82
	230.48415	5.06761	3.65 \pm 0.82	9.51	147	2.15 \pm 0.44	11.23	7.81 \pm 0.74	13.26 \pm 2.19
	230.48526	5.07067	6.51 \pm 0.56	15.59	155	< 2.10	15.30	8.06 \pm 1.10	23.65 \pm 1.50
	230.48637	5.07372	5.17 \pm 0.41	9.02	175	3.06 \pm 0.46	10.78	7.96 \pm 1.14	18.77 \pm 1.09
	230.48721	5.07650	4.44 \pm 0.24	20.96	89	4.94 \pm 0.64	21.03	7.89 \pm 1.29	16.13 \pm 0.64
	230.48805	5.07956	2.73 \pm 0.37	13.97	93	< 2.01	14.65	7.68 \pm 0.92	9.91 \pm 1.00
NGC6217	248.14055	78.20620	7.79 \pm 0.76	15.09	222	6.68 \pm 1.47	30.50	8.34 \pm 1.05	31.10 \pm 2.23
	248.14734	78.20342	9.54 \pm 0.79	9.14	199	6.58 \pm 0.75	16.44	8.43 \pm 1.12	38.07 \pm 2.31
	248.15413	78.20092	22.27 \pm 0.76	16.06	201	22.26 \pm 1.60	34.89	8.79 \pm 1.48	88.90 \pm 2.24
	248.16228	78.19814	19.65 \pm 0.65	15.74	208	17.71 \pm 1.77	37.89	8.74 \pm 1.49	78.45 \pm 1.91
	248.17043	78.19537	15.40 \pm 0.68	17.27	170	13.60 \pm 1.43	34.00	8.63 \pm 1.38	61.49 \pm 1.99
	248.17722	78.19286	6.92 \pm 0.45	16.65	227	5.12 \pm 1.48	30.33	8.29 \pm 1.22	27.64 \pm 1.31
	248.18401	78.19009	8.77 \pm 0.68	13.93	197	6.26 \pm 1.23	27.05	8.39 \pm 1.14	35.02 \pm 2.01
NGC6951	309.28613	66.10472	11.79 \pm 0.31	12.95	242	8.02 \pm 1.16	23.13	8.43 \pm 1.59	39.65 \pm 0.77
	309.29367	66.10500	18.80 \pm 0.63	17.12	391	14.16 \pm 3.23	50.51	8.63 \pm 1.49	63.21 \pm 1.56
	309.30121	66.10528	30.44 \pm 0.55	12.34	439	22.99 \pm 1.53	22.52	8.84 \pm 1.75	102.38 \pm 1.36
	309.30875	66.10555	41.73 \pm 0.54	6.17	474	48.63 \pm 0.48	6.82	8.98 \pm 1.89	140.33 \pm 1.34
	309.31628	66.10583	29.69 \pm 0.86	9.88	464	34.44 \pm 1.23	17.64	8.83 \pm 1.55	99.86 \pm 2.12
	309.32382	66.10611	15.00 \pm 1.04	8.09	355	8.35 \pm 0.87	14.28	8.53 \pm 1.19	50.44 \pm 2.58
	309.33139	66.10639	12.54 \pm 0.63	7.70	230	11.94 \pm 0.55	11.23	8.45 \pm 1.32	42.18 \pm 1.57
NGC7723	354.73221	-12.96361	6.02 \pm 0.49	16.66	140	3.50 \pm 1.64	42.96	8.21 \pm 1.12	17.60 \pm 1.06
	354.73508	-12.96222	11.80 \pm 0.62	12.15	268	8.48 \pm 1.54	29.13	8.50 \pm 1.30	34.52 \pm 1.33
	354.73792	-12.96083	15.37 \pm 0.54	12.87	305	19.77 \pm 1.58	27.96	8.62 \pm 1.47	44.97 \pm 1.17
	354.74075	-12.95944	11.84 \pm 0.68	11.70	299	5.85 \pm 1.28	22.82	8.50 \pm 1.27	34.64 \pm 1.46
	354.74362	-12.95806	7.54 \pm 0.73	9.63	184	5.15 \pm 1.05	24.04	8.31 \pm 1.05	22.07 \pm 1.57

Appendix C: SFEs with different direct and hybrid recipes of SFRs

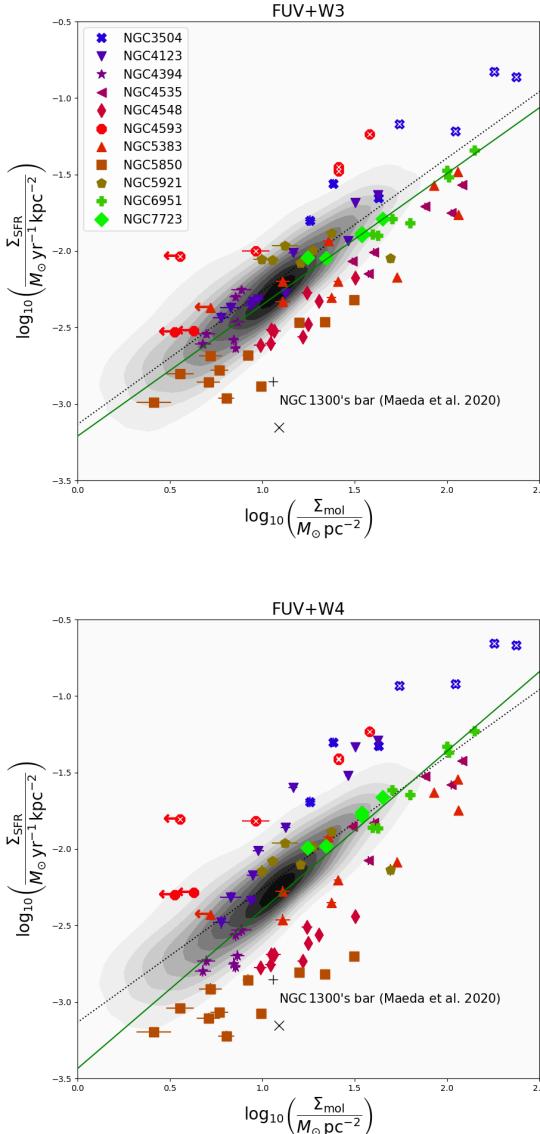


Fig. C.1. As in the left panels of Fig. 6, but using hybrid SFR recipes corresponding to Eq. C.5 (FUV+WISE 3) and Eq. C.7 (FUV+WISE 4), as indicated on top of the panels.

Here we show the relationship between star formation rate and molecular gas surface densities (Sects. 4 and 5) using a variety of hybrid (Fig. C.1) and direct (Fig. C.2) recipes for the calculation of SF calibrated by Leroy et al. (2019):

$$\text{SFR}_{\text{FUV}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{FUV}} [\text{erg s}^{-1}]}{10^{43.42}}, \quad (\text{C.1})$$

Table C.1. Slope, y -intercept, and $1-\sigma$ forward scatter of residuals in linear fit $\log_{10}\left(\frac{\Sigma_{\text{SFR}}}{M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}}\right) = m \times \log_{10}\left(\frac{\Sigma_{\text{mol}}}{M_{\odot} \text{ pc}^{-2}}\right) + b$ for the different SF tracers and recipes (1st column; see text).

KS	slope (m)	y -intercept (b)	scatter ($1-\sigma$)
NUV	0.81 ± 0.08	-3.62 ± 0.11	0.35
FUV	0.69 ± 0.10	-3.87 ± 0.13	0.41
W3	0.92 ± 0.06	-3.26 ± 0.08	0.25
W4	1.15 ± 0.09	-3.53 ± 0.13	0.40
NUV+W3	0.88 ± 0.06	-3.20 ± 0.09	0.27
NUV+W4	1.04 ± 0.09	-3.37 ± 0.12	0.37
FUV+W3	0.86 ± 0.06	-3.21 ± 0.09	0.26
FUV+W4	1.04 ± 0.09	-3.44 ± 0.13	0.39

$$\text{SFR}_{\text{NUV}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{NUV}} [\text{erg s}^{-1}]}{10^{43.24}}, \quad (\text{C.2})$$

$$\text{SFR}_{\text{W3}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{W3}} [\text{erg s}^{-1}]}{10^{42.70}}, \quad (\text{C.3})$$

$$\text{SFR}_{\text{W4}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{W4}} [\text{erg s}^{-1}]}{10^{42.63}}. \quad (\text{C.4})$$

We take advantage of the combination of UV and infrared (IR) data, which allows us to obtain dust-corrected SFRs:

$$\text{SFR}_{\text{FUV,W3}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{FUV}} [\text{erg s}^{-1}]}{10^{43.42}} + \frac{L_{\text{W3}} [\text{erg s}^{-1}]}{10^{42.79}}, \quad (\text{C.5})$$

$$\text{SFR}_{\text{NUV,W3}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{NUV}} [\text{erg s}^{-1}]}{10^{43.24}} + \frac{L_{\text{W3}} [\text{erg s}^{-1}]}{10^{42.86}}, \quad (\text{C.6})$$

$$\text{SFR}_{\text{FUV,W4}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{FUV}} [\text{erg s}^{-1}]}{10^{43.42}} + \frac{L_{\text{W4}} [\text{erg s}^{-1}]}{10^{42.73}}, \quad (\text{C.7})$$

$$\text{SFR}_{\text{NUV,W4}} [M_{\odot} \text{ yr}^{-1}] = \frac{L_{\text{NUV}} [\text{erg s}^{-1}]}{10^{43.24}} + \frac{L_{\text{W4}} [\text{erg s}^{-1}]}{10^{42.79}}. \quad (\text{C.8})$$

In Table C.1 we indicate the slope, y -intercept, and 1σ forward scatter of the residuals associated with the linear fit (log space) of the KS law. For a discussion of the differences of the KS law for the different recipes, see Sects. 4 and 5 in this paper.

We also estimate Σ_{SFR} from NUV and WISE 4 combined (Eq. C.8) to study the KS relation (Fig. C.3) as a function of total stellar mass, specific star formation rate, bulge stellar mass, and projected surface density to the third-nearest neighbour galaxy. The trends are qualitatively the same as reported in Sect. 5, where Σ_{SFR} was estimated from NUV and WISE 3 (Fig. 7).

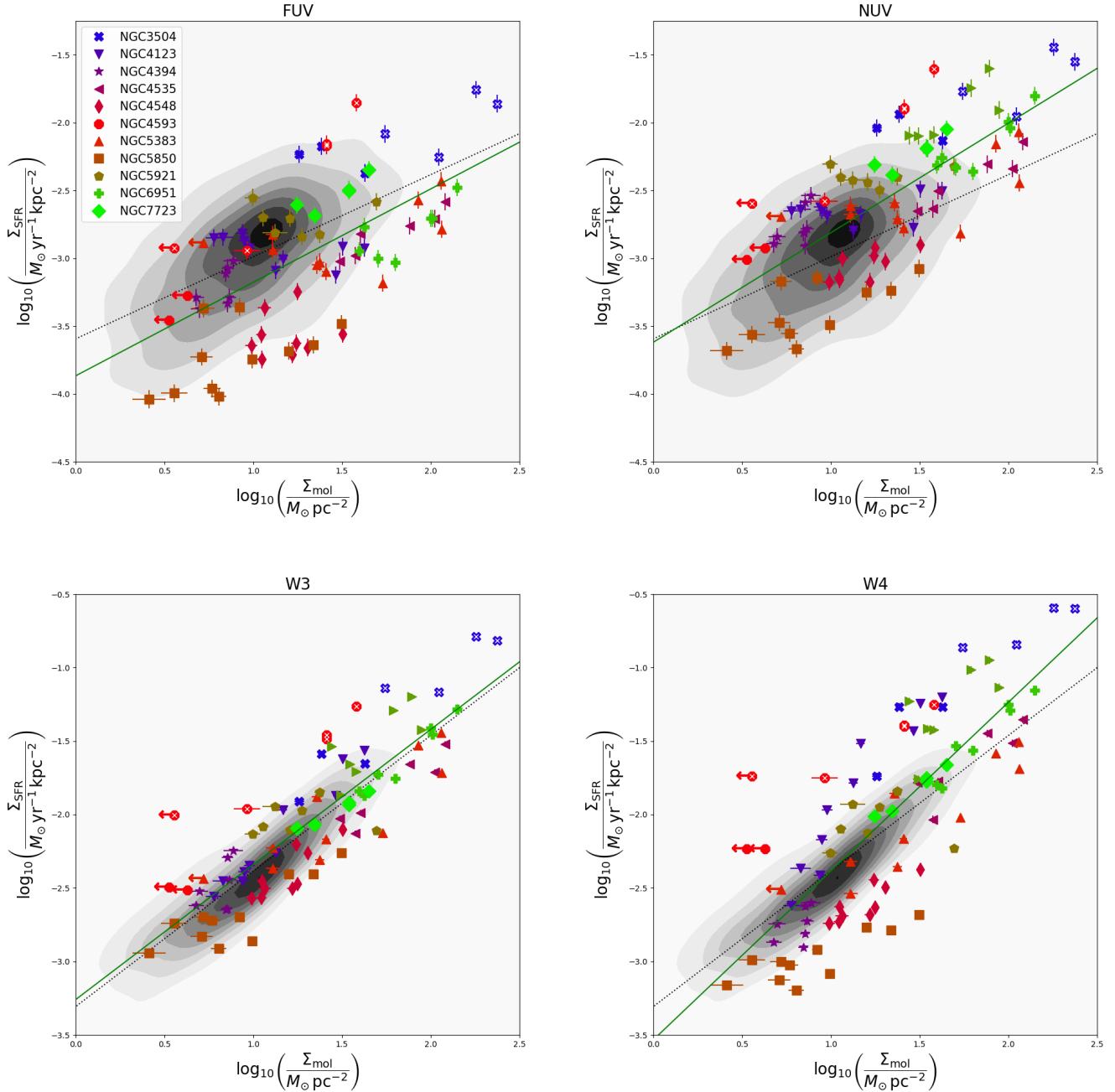


Fig. C.2. As in the left panels of Fig. 6, but using direct SFR recipes corresponding to Eq. C.1 (FUV), Eq. C.2 (NUV), Eq. C.3 (WISE 3), and Eq. C.4 (WISE 4), as indicated on top of the panels. The grey contours trace the measurements from Bigiel et al. (2011) using FUV (upper panels) and Spitzer 22 μm (Eq. 1 in Bigiel et al. 2008).

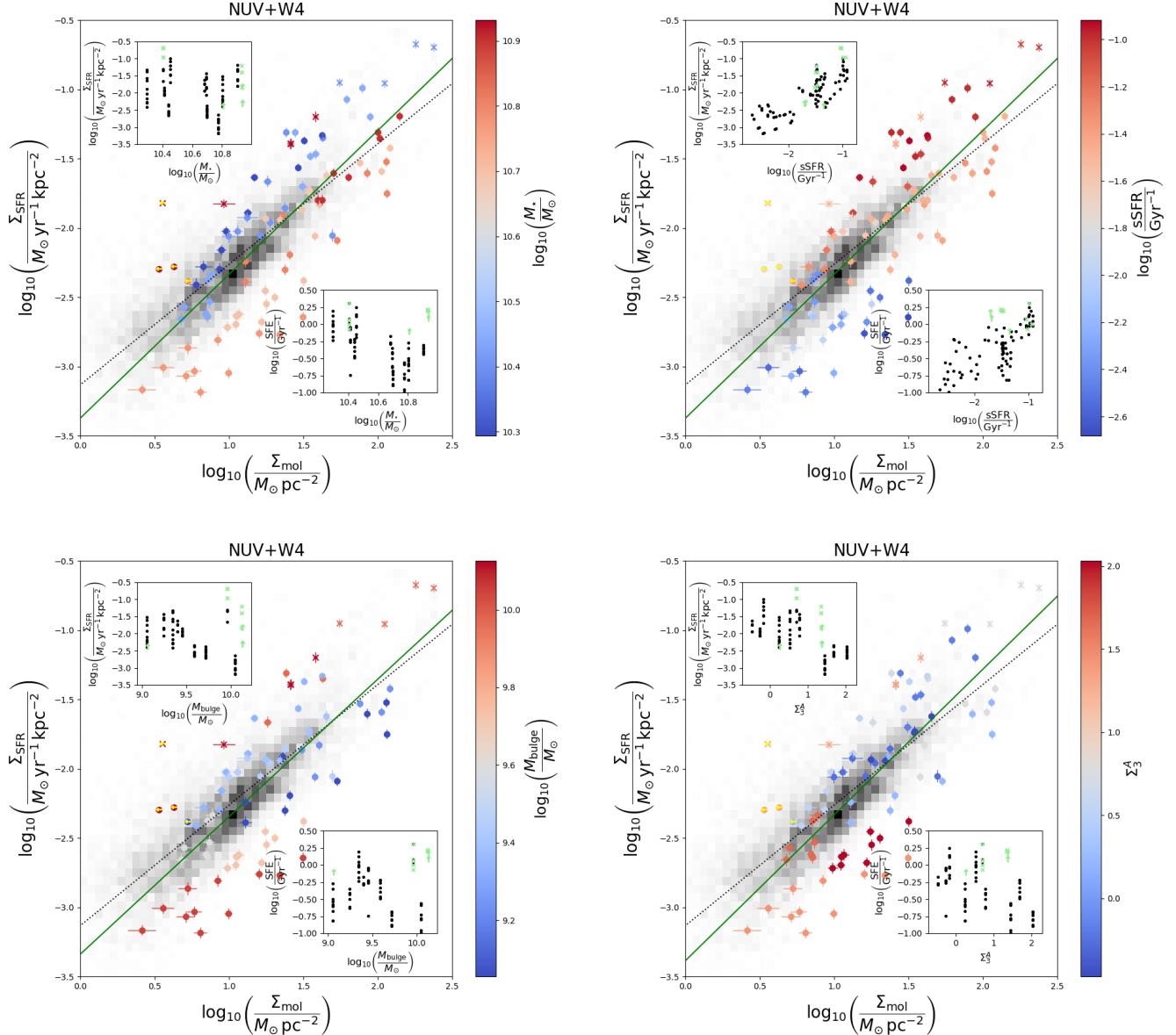


Fig. C.3. As in Fig. 7, but calculating SFRs from the combination of NUV and WISE 4.

Appendix D: Aperture photometry output

The following data are listed in Table D.1:

- Column 1: Galaxy name;
- Column 2: Right ascension of centre of the IRAM-30 m pointing (in degrees);
- Column 3: Declination of centre of the IRAM-30 m pointing (in degrees);
- Column 4: Decimal logarithm of the flux within beam in the NUV images (in janskys) and its uncertainty;
- Column 5: Decimal logarithm of the flux within beam in the FUV images (in janskys) and its uncertainty;
- Column 6: Decimal logarithm of the flux within beam in the WISE 1 images (in janskys) and its uncertainty;

- Column 7: Decimal logarithm of the flux within beam in the WISE 2 images (in janskys) and its uncertainty;
- Column 8: Decimal logarithm of the flux within beam in the WISE 3 images (in janskys) and its uncertainty;
- Column 9: Decimal logarithm of the flux within beam in the WISE 4 images (in janskys) and its uncertainty;
- Column 10: Σ_{SFR} - Star formation surface density (in $M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$) obtained from NUV and WISE 3 images using Eqs. 10 and 12, and its error;
- Column 11: Σ_{\star} - Stellar mass surface density (in $M_{\odot} \text{ pc}^{-2}$) obtained from WISE 1 images using Eq. 13, and its error.

Table D.1. Aperture photometry, star formation rate and stellar surface densities.

Galaxy	RA ($^{\circ}$)	DEC ($^{\circ}$)	$\log_{10} F_{\text{NUV}}$ (Jy)	$\log_{10} F_{\text{FUV}}$ (Jy)	$\log_{10} F_{\text{W1}}$ (Jy)	$\log_{10} F_{\text{W2}}$ (Jy)	$\log_{10} F_{\text{W3}}$ (Jy)	$\log_{10} F_{\text{W4}}$ (Jy)	$\log_{10} \Sigma_{\text{SFR}}$ ($M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$)	$\log_{10} \Sigma_{\star}$ ($M_{\odot} \text{ pc}^{-2}$)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
NGC3504	165.79059	27.98028	-2.99 \pm 0.06	-3.23 \pm 0.06	-1.88 \pm 0.01	-2.08 \pm 0.01	-1.33 \pm 0.02	-0.75 \pm 0.02	-1.64 \pm 0.03	2.62 \pm 0.01
	165.79248	27.97750	-2.81 \pm 0.06	-3.11 \pm 0.06	-1.52 \pm 0.01	-1.64 \pm 0.01	-0.84 \pm 0.02	-0.32 \pm 0.02	-1.24 \pm 0.03	2.99 \pm 0.01
	165.79436	27.97500	-2.41 \pm 0.06	-2.71 \pm 0.06	-1.18 \pm 0.01	-1.26 \pm 0.01	-0.49 \pm 0.02	-0.08 \pm 0.02	-0.87 \pm 0.03	3.33 \pm 0.01
	165.79625	27.97222	-2.30 \pm 0.06	-2.61 \pm 0.06	-1.14 \pm 0.01	-1.22 \pm 0.01	-0.46 \pm 0.02	-0.07 \pm 0.02	-0.83 \pm 0.03	3.36 \pm 0.01
	165.79814	27.96944	-2.63 \pm 0.06	-2.94 \pm 0.06	-1.47 \pm 0.01	-1.59 \pm 0.01	-0.81 \pm 0.02	-0.34 \pm 0.02	-1.17 \pm 0.03	3.03 \pm 0.01
	165.80002	27.96694	-2.80 \pm 0.06	-3.03 \pm 0.06	-1.83 \pm 0.01	-2.02 \pm 0.01	-1.26 \pm 0.02	-0.74 \pm 0.02	-1.53 \pm 0.04	2.68 \pm 0.01
	165.80191	27.96417	-2.90 \pm 0.06	-3.09 \pm 0.06	-2.11 \pm 0.01	-2.31 \pm 0.01	-1.58 \pm 0.02	-1.22 \pm 0.02	-1.75 \pm 0.04	2.40 \pm 0.01
NGC4123	182.03407	2.88211	-3.33 \pm 0.06	-3.51 \pm 0.06	-2.36 \pm 0.01	-2.60 \pm 0.01	-1.97 \pm 0.02	-1.74 \pm 0.02	-2.32 \pm 0.04	1.99 \pm 0.01
	182.03685	2.88128	-3.35 \pm 0.06	-3.57 \pm 0.06	-2.27 \pm 0.01	-2.51 \pm 0.01	-1.91 \pm 0.02	-1.49 \pm 0.02	-2.29 \pm 0.04	2.08 \pm 0.01
	182.03990	2.88044	-3.50 \pm 0.06	-3.79 \pm 0.06	-2.20 \pm 0.01	-2.43 \pm 0.01	-1.78 \pm 0.02	-1.11 \pm 0.02	-2.27 \pm 0.03	2.15 \pm 0.01
	182.04269	2.87961	-3.48 \pm 0.06	-3.82 \pm 0.06	-1.99 \pm 0.01	-2.20 \pm 0.01	-1.39 \pm 0.02	-0.75 \pm 0.02	-1.96 \pm 0.03	2.36 \pm 0.01
	182.04575	2.87878	-3.21 \pm 0.06	-3.62 \pm 0.06	-1.76 \pm 0.01	-1.94 \pm 0.01	-1.08 \pm 0.02	-0.52 \pm 0.02	-1.66 \pm 0.03	2.59 \pm 0.01
	182.04881	2.87794	-3.20 \pm 0.06	-3.61 \pm 0.06	-1.80 \pm 0.01	-1.99 \pm 0.01	-1.14 \pm 0.02	-0.57 \pm 0.02	-1.70 \pm 0.03	2.55 \pm 0.01
	182.05159	2.87711	-3.38 \pm 0.06	-3.71 \pm 0.06	-2.04 \pm 0.01	-2.25 \pm 0.01	-1.49 \pm 0.02	-0.84 \pm 0.02	-2.02 \pm 0.03	2.31 \pm 0.01
	182.05466	2.87628	-3.39 \pm 0.06	-3.62 \pm 0.06	-2.22 \pm 0.01	-2.46 \pm 0.01	-1.87 \pm 0.02	-1.29 \pm 0.02	-2.29 \pm 0.04	2.13 \pm 0.01
	182.05743	2.87544	-3.34 \pm 0.06	-3.55 \pm 0.06	-2.31 \pm 0.01	-2.54 \pm 0.01	-1.97 \pm 0.02	-1.69 \pm 0.02	-2.32 \pm 0.04	2.04 \pm 0.01
	182.06049	2.87461	-3.35 \pm 0.06	-3.55 \pm 0.06	-2.44 \pm 0.01	-2.65 \pm 0.01	-2.07 \pm 0.02	-1.94 \pm 0.02	-2.38 \pm 0.04	1.92 \pm 0.01
NGC4394	186.47528	18.22185	-3.62 \pm 0.06	-3.91 \pm 0.06	-2.16 \pm 0.01	-2.42 \pm 0.01	-2.26 \pm 0.02	-2.33 \pm 0.02	-2.51 \pm 0.04	2.30 \pm 0.01
	186.47733	18.21935	-3.70 \pm 0.06	-4.09 \pm 0.06	-1.94 \pm 0.01	-2.20 \pm 0.01	-2.24 \pm 0.02	-2.29 \pm 0.02	-2.53 \pm 0.04	2.51 \pm 0.01
	186.47937	18.21685	-3.59 \pm 0.06	-4.09 \pm 0.06	-1.64 \pm 0.01	-1.90 \pm 0.01	-2.06 \pm 0.02	-2.15 \pm 0.02	-2.38 \pm 0.04	2.81 \pm 0.01
	186.48141	18.21435	-3.35 \pm 0.06	-3.82 \pm 0.06	-1.40 \pm 0.01	-1.67 \pm 0.01	-1.86 \pm 0.02	-2.03 \pm 0.02	-2.17 \pm 0.04	3.06 \pm 0.01
	186.48347	18.21185	-3.39 \pm 0.06	-3.87 \pm 0.06	-1.44 \pm 0.01	-1.71 \pm 0.01	-1.91 \pm 0.02	-2.05 \pm 0.02	-2.21 \pm 0.04	3.01 \pm 0.01
	186.48552	18.20935	-3.65 \pm 0.06	-4.18 \pm 0.06	-1.73 \pm 0.01	-1.99 \pm 0.01	-2.14 \pm 0.02	-2.17 \pm 0.02	-2.45 \pm 0.04	2.72 \pm 0.01
	186.48756	18.20685	-3.71 \pm 0.06	-4.13 \pm 0.06	-2.00 \pm 0.01	-2.26 \pm 0.01	-2.26 \pm 0.02	-2.24 \pm 0.02	-2.55 \pm 0.04	2.45 \pm 0.01
NGC4535	188.58041	8.19330	-3.33 \pm 0.06	-3.64 \pm 0.06	-1.91 \pm 0.01	-2.13 \pm 0.01	-1.63 \pm 0.02	-1.22 \pm 0.02	-1.99 \pm 0.03	2.56 \pm 0.01
	188.58211	8.19553	-3.16 \pm 0.06	-3.53 \pm 0.06	-1.72 \pm 0.01	-1.93 \pm 0.01	-1.35 \pm 0.02	-0.96 \pm 0.02	-1.74 \pm 0.03	2.75 \pm 0.01
	188.58463	8.19803	-2.97 \pm 0.06	-3.40 \pm 0.06	-1.56 \pm 0.01	-1.76 \pm 0.01	-1.16 \pm 0.02	-0.80 \pm 0.02	-1.55 \pm 0.03	2.91 \pm 0.01
	188.58659	8.19997	-3.13 \pm 0.06	-3.58 \pm 0.06	-1.67 \pm 0.01	-1.88 \pm 0.01	-1.30 \pm 0.02	-0.89 \pm 0.02	-1.69 \pm 0.03	2.80 \pm 0.01
	188.58884	8.20275	-3.48 \pm 0.06	-3.84 \pm 0.06	-1.91 \pm 0.01	-2.15 \pm 0.01	-1.66 \pm 0.02	-1.23 \pm 0.02	-2.06 \pm 0.03	2.56 \pm 0.01
	188.59052	8.20469	-3.46 \pm 0.06	-3.80 \pm 0.06	-1.98 \pm 0.01	-2.22 \pm 0.01	-1.77 \pm 0.02	-1.48 \pm 0.02	-2.13 \pm 0.03	2.49 \pm 0.01
NGC4548	188.84856	14.49091	-3.68 \pm 0.06	-4.00 \pm 0.06	-1.98 \pm 0.01	-2.23 \pm 0.01	-2.05 \pm 0.02	-2.01 \pm 0.02	-2.45 \pm 0.03	2.43 \pm 0.01
	188.85143	14.49230	-3.90 \pm 0.06	-4.32 \pm 0.06	-1.87 \pm 0.01	-2.13 \pm 0.01	-2.14 \pm 0.02	-2.11 \pm 0.02	-2.58 \pm 0.03	2.54 \pm 0.01
	188.85429	14.49369	-3.93 \pm 0.06	-4.46 \pm 0.06	-1.69 \pm 0.01	-1.95 \pm 0.01	-2.08 \pm 0.02	-2.06 \pm 0.02	-2.55 \pm 0.03	2.72 \pm 0.01
	188.85716	14.49508	-3.78 \pm 0.06	-4.41 \pm 0.06	-1.41 \pm 0.01	-1.65 \pm 0.01	-1.83 \pm 0.02	-1.88 \pm 0.02	-2.32 \pm 0.03	3.00 \pm 0.01
	188.86003	14.49647	-3.66 \pm 0.06	-4.31 \pm 0.06	-1.24 \pm 0.01	-1.48 \pm 0.01	-1.68 \pm 0.02	-1.76 \pm 0.02	-2.17 \pm 0.03	3.17 \pm 0.01
	188.86290	14.49786	-3.74 \pm 0.06	-4.38 \pm 0.06	-1.36 \pm 0.01	-1.60 \pm 0.01	-1.78 \pm 0.02	-1.82 \pm 0.02	-2.27 \pm 0.03	3.05 \pm 0.01
	188.86577	14.49925	-3.92 \pm 0.06	-4.50 \pm 0.06	-1.65 \pm 0.01	-1.90 \pm 0.01	-2.02 \pm 0.02	-2.01 \pm 0.02	-2.50 \pm 0.03	2.76 \pm 0.01
	188.86864	14.50064	-3.94 \pm 0.06	-4.40 \pm 0.06	-1.85 \pm 0.01	-2.11 \pm 0.01	-2.14 \pm 0.02	-2.12 \pm 0.02	-2.60 \pm 0.03	2.56 \pm 0.01
	188.87151	14.50202	-3.76 \pm 0.06	-4.12 \pm 0.06	-1.96 \pm 0.01	-2.22 \pm 0.01	-2.07 \pm 0.02	-2.07 \pm 0.02	-2.50 \pm 0.03	2.44 \pm 0.01

Table D.1. continued.

Galaxy	RA ($^{\circ}$)	DEC ($^{\circ}$)	$\log_{10} F_{\text{NUV}}$ (Jy)	$\log_{10} F_{\text{FUV}}$ (Jy)	$\log_{10} F_{\text{W1}}$ (Jy)	$\log_{10} F_{\text{W2}}$ (Jy)	$\log_{10} F_{\text{W3}}$ (Jy)	$\log_{10} F_{\text{W4}}$ (Jy)	$\log_{10} \Sigma_{\text{SFR}}$ ($M_{\ast} \text{ yr}^{-1} \text{kpc}^{-2}$)	$\log_{10} \Sigma_{\star}$ ($M_{\ast} \text{ pc}^{-2}$)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
NGC4593	189.90707	-5.34977	-3.80 \pm 0.06	-4.24 \pm 0.06	-2.02 \pm 0.01	-2.25 \pm 0.01	-2.10 \pm 0.02	-1.64 \pm 0.02	-2.50 \pm 0.03	2.42 \pm 0.01
	189.90958	-5.34783	-3.37 \pm 0.06	-3.73 \pm 0.06	-1.74 \pm 0.01	-1.87 \pm 0.01	-1.57 \pm 0.02	-1.16 \pm 0.02	-1.99 \pm 0.03	2.70 \pm 0.01
	189.91209	-5.34616	-2.68 \pm 0.06	-2.94 \pm 0.06	-1.34 \pm 0.01	-1.37 \pm 0.01	-1.06 \pm 0.02	-0.81 \pm 0.02	-1.43 \pm 0.03	3.10 \pm 0.01
	189.91460	-5.34421	-2.40 \pm 0.06	-2.64 \pm 0.06	-1.15 \pm 0.01	-1.16 \pm 0.01	-0.87 \pm 0.02	-0.66 \pm 0.02	-1.20 \pm 0.04	3.29 \pm 0.01
	189.91711	-5.34227	-2.69 \pm 0.06	-2.96 \pm 0.06	-1.37 \pm 0.01	-1.40 \pm 0.01	-1.09 \pm 0.02	-0.80 \pm 0.02	-1.45 \pm 0.03	3.07 \pm 0.01
	189.91962	-5.34060	-3.39 \pm 0.06	-3.71 \pm 0.06	-1.77 \pm 0.01	-1.90 \pm 0.01	-1.61 \pm 0.02	-1.15 \pm 0.02	-2.03 \pm 0.03	2.67 \pm 0.01
	189.92213	-5.33866	-3.72 \pm 0.06	-4.06 \pm 0.06	-2.02 \pm 0.01	-2.25 \pm 0.01	-2.12 \pm 0.02	-1.64 \pm 0.02	-2.48 \pm 0.03	2.42 \pm 0.01
NGC5383	209.25691	41.85528	-3.46 \pm 0.06	-3.72 \pm 0.06	-2.34 \pm 0.01	-2.56 \pm 0.01	-1.96 \pm 0.02	-1.94 \pm 0.02	-2.29 \pm 0.04	2.09 \pm 0.01
	209.25989	41.85305	-3.50 \pm 0.06	-3.81 \pm 0.06	-2.23 \pm 0.01	-2.46 \pm 0.01	-1.90 \pm 0.02	-1.76 \pm 0.02	-2.27 \pm 0.03	2.20 \pm 0.01
	209.26324	41.85083	-3.60 \pm 0.06	-3.96 \pm 0.06	-2.11 \pm 0.01	-2.33 \pm 0.01	-1.72 \pm 0.02	-1.42 \pm 0.02	-2.17 \pm 0.03	2.32 \pm 0.01
	209.26622	41.84861	-3.23 \pm 0.06	-3.57 \pm 0.06	-1.82 \pm 0.01	-2.01 \pm 0.01	-1.31 \pm 0.02	-1.09 \pm 0.02	-1.77 \pm 0.03	2.62 \pm 0.01
	209.26958	41.84639	-2.85 \pm 0.06	-3.21 \pm 0.06	-1.55 \pm 0.01	-1.74 \pm 0.01	-1.04 \pm 0.02	-0.91 \pm 0.02	-1.47 \pm 0.03	2.88 \pm 0.01
	209.27293	41.84417	-2.94 \pm 0.06	-3.35 \pm 0.06	-1.61 \pm 0.01	-1.81 \pm 0.01	-1.13 \pm 0.02	-0.99 \pm 0.02	-1.56 \pm 0.03	2.82 \pm 0.01
	209.27592	41.84195	-3.38 \pm 0.06	-3.83 \pm 0.06	-1.91 \pm 0.01	-2.12 \pm 0.01	-1.48 \pm 0.02	-1.26 \pm 0.02	-1.93 \pm 0.03	2.52 \pm 0.01
	209.27928	41.83972	-3.56 \pm 0.06	-3.88 \pm 0.06	-2.12 \pm 0.01	-2.34 \pm 0.01	-1.77 \pm 0.02	-1.57 \pm 0.02	-2.20 \pm 0.03	2.31 \pm 0.01
	209.28226	41.83750	-3.40 \pm 0.06	-3.61 \pm 0.06	-2.21 \pm 0.01	-2.43 \pm 0.01	-1.82 \pm 0.02	-1.73 \pm 0.02	-2.18 \pm 0.03	2.22 \pm 0.01
	209.28561	41.83528	-3.48 \pm 0.06	-3.67 \pm 0.06	-2.44 \pm 0.01	-2.65 \pm 0.01	-2.04 \pm 0.02	-1.92 \pm 0.02	-2.34 \pm 0.04	1.99 \pm 0.01
NGC5850	226.76862	1.55108	-3.95 \pm 0.06	-4.14 \pm 0.06	-2.32 \pm 0.01	-2.58 \pm 0.01	-2.29 \pm 0.02	-2.40 \pm 0.02	-2.69 \pm 0.03	2.11 \pm 0.01
	226.77139	1.54969	-4.25 \pm 0.06	-4.50 \pm 0.06	-2.24 \pm 0.01	-2.51 \pm 0.01	-2.42 \pm 0.02	-2.52 \pm 0.02	-2.87 \pm 0.03	2.18 \pm 0.01
	226.77417	1.54830	-4.45 \pm 0.06	-4.79 \pm 0.06	-2.13 \pm 0.01	-2.40 \pm 0.01	-2.50 \pm 0.02	-2.59 \pm 0.02	-2.97 \pm 0.03	2.29 \pm 0.01
	226.77696	1.54691	-4.33 \pm 0.06	-4.73 \pm 0.06	-1.89 \pm 0.01	-2.15 \pm 0.01	-2.31 \pm 0.02	-2.42 \pm 0.02	-2.80 \pm 0.03	2.54 \pm 0.01
	226.77974	1.54552	-4.02 \pm 0.06	-4.41 \pm 0.06	-1.54 \pm 0.01	-1.82 \pm 0.01	-1.99 \pm 0.02	-2.19 \pm 0.02	-2.48 \pm 0.03	2.88 \pm 0.01
	226.78252	1.54413	-3.85 \pm 0.06	-4.26 \pm 0.06	-1.39 \pm 0.01	-1.66 \pm 0.01	-1.85 \pm 0.02	-2.08 \pm 0.02	-2.33 \pm 0.03	3.04 \pm 0.01
	226.78529	1.54274	-4.03 \pm 0.06	-4.46 \pm 0.06	-1.54 \pm 0.01	-1.81 \pm 0.01	-2.00 \pm 0.02	-2.17 \pm 0.02	-2.48 \pm 0.03	2.88 \pm 0.01
	226.78807	1.54136	-4.34 \pm 0.06	-4.77 \pm 0.06	-1.88 \pm 0.01	-2.15 \pm 0.01	-2.33 \pm 0.02	-2.39 \pm 0.02	-2.81 \pm 0.03	2.54 \pm 0.01
	226.79085	1.53997	-4.46 \pm 0.06	-4.81 \pm 0.06	-2.13 \pm 0.01	-2.39 \pm 0.01	-2.53 \pm 0.02	-2.56 \pm 0.02	-3.00 \pm 0.03	2.30 \pm 0.01
	226.79362	1.53858	-4.27 \pm 0.06	-4.52 \pm 0.06	-2.24 \pm 0.01	-2.50 \pm 0.01	-2.45 \pm 0.02	-2.48 \pm 0.02	-2.89 \pm 0.03	2.19 \pm 0.01
	226.79640	1.53719	-3.92 \pm 0.06	-4.13 \pm 0.06	-2.29 \pm 0.01	-2.56 \pm 0.01	-2.29 \pm 0.02	-2.32 \pm 0.02	-2.68 \pm 0.03	2.13 \pm 0.01
NGC5921	230.48247	5.06178	-3.11 \pm 0.06	-3.37 \pm 0.06	-2.02 \pm 0.01	-2.24 \pm 0.01	-1.71 \pm 0.02	-1.64 \pm 0.02	-1.99 \pm 0.04	2.42 \pm 0.01
	230.48331	5.06483	-3.19 \pm 0.06	-3.49 \pm 0.06	-1.86 \pm 0.01	-2.09 \pm 0.01	-1.69 \pm 0.02	-1.51 \pm 0.02	-2.01 \pm 0.04	2.58 \pm 0.01
	230.48415	5.06761	-3.22 \pm 0.06	-3.60 \pm 0.06	-1.62 \pm 0.01	-1.85 \pm 0.01	-1.55 \pm 0.02	-1.34 \pm 0.02	-1.93 \pm 0.03	2.82 \pm 0.01
	230.48526	5.07067	-3.19 \pm 0.06	-3.61 \pm 0.06	-1.48 \pm 0.01	-1.72 \pm 0.01	-1.45 \pm 0.02	-1.25 \pm 0.02	-1.86 \pm 0.03	2.96 \pm 0.01
	230.48637	5.07372	-3.29 \pm 0.06	-3.63 \pm 0.06	-1.66 \pm 0.01	-1.89 \pm 0.01	-1.58 \pm 0.02	-1.36 \pm 0.02	-1.98 \pm 0.03	2.78 \pm 0.01
	230.48721	5.07650	-3.23 \pm 0.06	-3.49 \pm 0.06	-1.89 \pm 0.01	-2.13 \pm 0.01	-1.71 \pm 0.02	-1.53 \pm 0.02	-2.04 \pm 0.04	2.54 \pm 0.01
	230.48805	5.07956	-3.10 \pm 0.06	-3.34 \pm 0.06	-2.04 \pm 0.01	-2.27 \pm 0.01	-1.74 \pm 0.02	-1.67 \pm 0.02	-2.00 \pm 0.04	2.40 \pm 0.01
NGC6217	248.14055	78.20620	-2.93 \pm 0.06	-	-2.04 \pm 0.01	-2.24 \pm 0.01	-1.51 \pm 0.02	-1.21 \pm 0.02	-1.76 \pm 0.04	2.44 \pm 0.01
	248.14734	78.20342	-2.92 \pm 0.06	-	-1.90 \pm 0.01	-2.10 \pm 0.01	-1.35 \pm 0.02	-0.87 \pm 0.02	-1.66 \pm 0.03	2.58 \pm 0.01
	248.15413	78.20092	-2.74 \pm 0.06	-	-1.68 \pm 0.01	-1.86 \pm 0.01	-1.07 \pm 0.02	-0.59 \pm 0.02	-1.41 \pm 0.03	2.80 \pm 0.01
	248.16228	78.19814	-2.43 \pm 0.06	-	-1.47 \pm 0.01	-1.63 \pm 0.01	-0.84 \pm 0.02	-0.40 \pm 0.02	-1.16 \pm 0.03	3.01 \pm 0.01
	248.17043	78.19537	-2.58 \pm 0.06	-	-1.56 \pm 0.01	-1.73 \pm 0.01	-0.94 \pm 0.02	-0.47 \pm 0.02	-1.27 \pm 0.03	2.92 \pm 0.01
	248.17722	78.19286	-2.93 \pm 0.06	-	-1.80 \pm 0.01	-1.98 \pm 0.01	-1.18 \pm 0.02	-0.68 \pm 0.02	-1.55 \pm 0.03	2.68 \pm 0.01
	248.18401	78.19009	-3.02 \pm 0.06	-	-1.98 \pm 0.01	-2.16 \pm 0.01	-1.30 \pm 0.02	-0.87 \pm 0.02	-1.66 \pm 0.03	2.50 \pm 0.01
NGC6951	309.28613	66.10472	-3.07 \pm 0.06	-3.70 \pm 0.06	-1.73 \pm 0.01	-1.95 \pm 0.01	-1.41 \pm 0.02	-1.17 \pm 0.02	-1.83 \pm 0.03	2.67 \pm 0.01
	309.29367	66.10500	-3.12 \pm 0.06	-3.79 \pm 0.06	-1.55 \pm 0.01	-1.77 \pm 0.01	-1.33 \pm 0.02	-0.94 \pm 0.02	-1.78 \pm 0.03	2.86 \pm 0.01
	309.30121	66.10528	-2.80 \pm 0.06	-3.46 \pm 0.06	-1.26 \pm 0.01	-1.48 \pm 0.01	-1.03 \pm 0.02	-0.67 \pm 0.02	-1.48 \pm 0.03	3.14 \pm 0.01
	309.30875	66.10555	-2.56 \pm 0.06	-3.23 \pm 0.06	-1.11 \pm 0.01	-1.33 \pm 0.01	-0.85 \pm 0.02	-0.53 \pm 0.02	-1.28 \pm 0.03	3.30 \pm 0.01
	309.31628	66.10583	-2.75 \pm 0.06	-3.46 \pm 0.06	-1.24 \pm 0.01	-1.46 \pm 0.01	-0.98 \pm 0.02	-0.63 \pm 0.02	-1.43 \pm 0.03	3.16 \pm 0.01
	309.32382	66.10611	-3.09 \pm 0.06	-3.76 \pm 0.06	-1.53 \pm 0.01	-1.76 \pm 0.01	-1.30 \pm 0.02	-0.91 \pm 0.02	-1.75 \pm 0.03	2.87 \pm 0.01
	309.33139	66.10639	-3.02 \pm 0.06	-3.52 \pm 0.06	-1.73 \pm 0.01	-1.95 \pm 0.01	-1.44 \pm 0.02	-1.20 \pm 0.02	-1.83 \pm 0.03	2.67 \pm 0.01
NGC7723	354.73221	-12.96361	-3.01 \pm 0.06	-3.30 \pm 0.06	-1.79 \pm 0.01	-2.03 \pm 0.01	-1.60 \pm 0.02	-1.33 \pm 0.02	-1.98 \pm 0.04	2.56 \pm 0.01
	354.73508	-12.96222	-2.88 \pm 0.06	-3.19 \pm 0.06	-1.59 \pm 0.01	-1.83 \pm 0.01	-1.45 \pm 0.02	-1.09 \pm 0.02	-1.84 \pm 0.04	2.75 \pm 0.01

Appendix E: Use of CO(2-1) to estimate molecular gas masses and star formation efficiencies

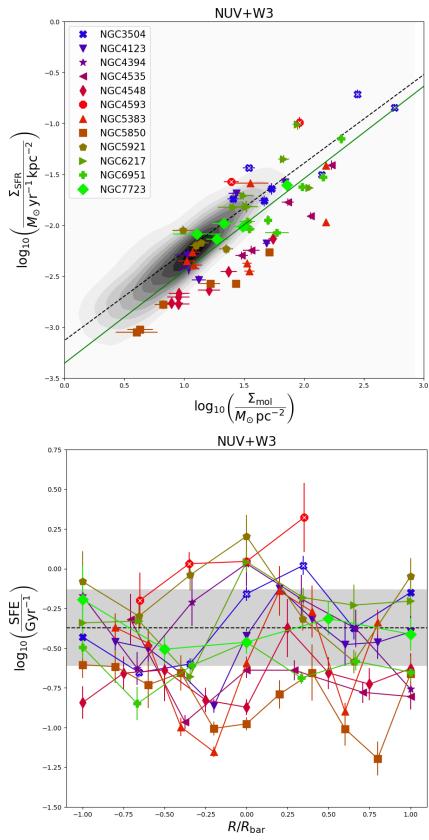


Fig. E.1. As in Fig. 6, but using CO(2-1) spectra and estimating SFEs within $10.75''$ apertures (based on the profiles shown in Fig. E.2). Only the pointings where CO(2-1) was detected are shown.

In order to test the sensitivity to resolution of the KS law within bars, we use the CO(2-1) spectra instead of CO(1-0). The latter, exploited in the previous sections, have better quality and two times lower calibration errors, but also two times worse resolution. We calculated M_{mol} and Σ_{mol} from $I_{\text{CO}(2-1)}$, assuming an intrinsic ratio $R_{21} = I_{\text{CO}(2-1)}/I_{\text{CO}(1-0)} = 0.8$ (Leroy et al. 2009), and then applied Eqs. 6, 7, and 8. We recalculate Σ_{SFR} and Σ_* following the steps described in Sect. 4 for the smaller CO(2-1) beam (i.e. FWHM=10.75'', corresponding to $\sim 1-2$ kpc) if CO(2-1) was detected. The resulting profiles are shown in Fig. E.2. These are more centrally peaked, as expected from the higher resolution, and the SF dips in quiescent bars are more clearly resolved. The $\Sigma_{\text{SFR}}-\Sigma_{\text{mol}}$ relation and SFE profiles (Fig. E.1) are similar to the ones reported in Sect. 5. The mean SFEs are 0.12 dex lower when using CO(2-1) to estimate Σ_{mol} as compared to CO(1-0); however, the differences in the CO(2-1)-based SFE amplitudes are dependent on the adopted R_{21} , which is not necessarily constant in and across galaxies (see next Sect. F). Despite the poorer data quality, the analysis of the KS law with CO(2-1) spectra reinforces the conclusions of this work: low SFE in the massive galaxies NGC 4548 and NGC 5850, which are known to lie in denser environments and host massive bulges (Sect. 5).

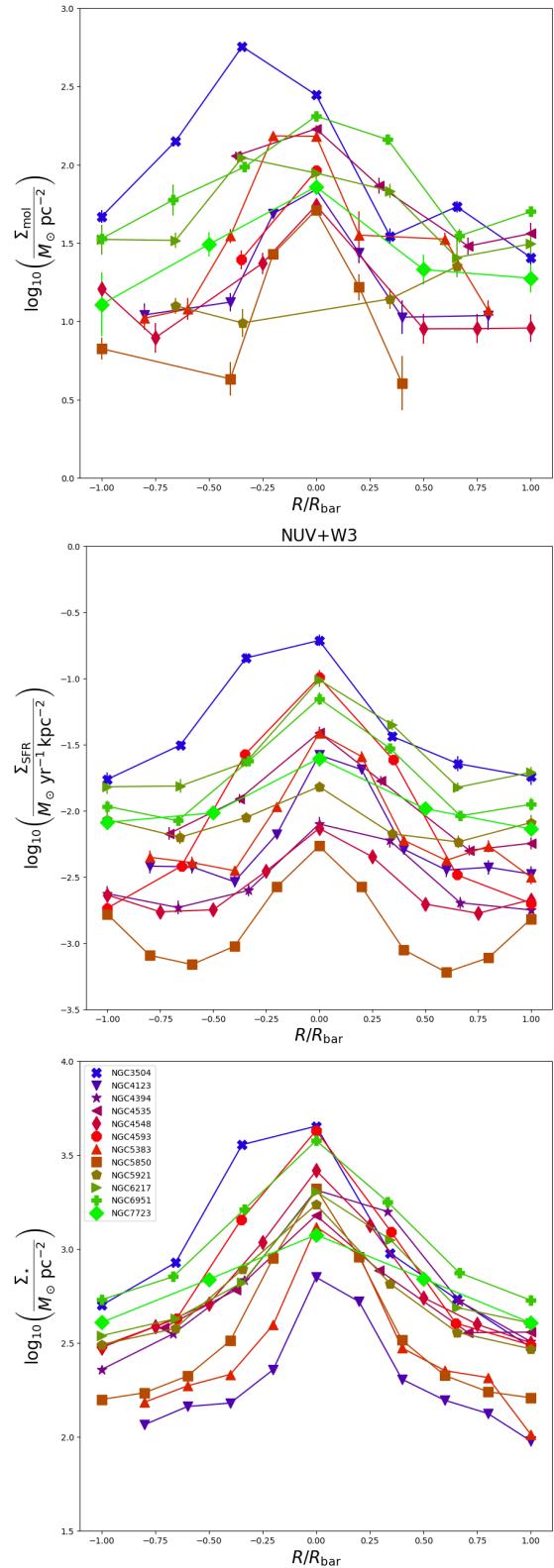


Fig. E.2. As in Fig. 4, but computing molecular gas surface densities from CO(2-1) spectra (*upper panel*) and calculating star formation rate (*central panel*, from the combination of NUV and W3) and stellar mass (*lower panel*) surfaces densities within the CO(2-1) pointings (FWHM = $10.75''$). In the upper panel only the pointings where CO(2-1) was detected are shown.

Appendix F: Gas excitation traced from CO(2-1) and CO(1-0) spectra

In this appendix we analyse the gas excitation across the bars in our sample. We calculate the ratio of the velocity-integrated line intensity of the IRAM-30 m CO(1-0) and CO(2-1) spectra: $R_{21} = I_{\text{CO}(2-1)} / I_{\text{CO}(1-0)}$. The mean value is $R_{21} = 0.89 \pm 0.03$ ($\sigma = 0.33$). In the upper panel of Fig. F.1 we show R_{21} as a function of radius along the individual bars. There is no clear trend of R_{21} along the bars, and no noticeable differences between the different galaxies. In the lower panel of Fig. F.1 we show the KS relation, colour-coded by R_{21} . There is no correlation between R_{21} and Σ_{SFR} or SFE (see inset plots of Fig. F.1).

In order to interpret R_{21} we have to take into account several effects. R_{21} depends on the excitation temperature and the opacity of the molecular gas. In addition, the observed R_{21} depends on the distribution of the molecular gas relative to the beam size since in our measurements the beam size of $I_{\text{CO}(2-1)}$ is half the beam size of $I_{\text{CO}(1-0)}$. The intrinsic value, $R_{21,\text{intrinsic}}$, that would be measured with matched aperture observations is $R_{21,\text{intrinsic}} > 1$ for optically thin gas, and (the more realistic case) $R_{21,\text{intrinsic}} \sim 0.6-1$ for optically thick gas in thermal equilibrium, where $R_{21,\text{intrinsic}}$ depends on the temperature of the gas. Values below ~ 0.5 are an indication of subthermal excitation due to a gas density well below the critical density so that the excitation temperature is below the kinetic temperature.

If the source is more extended than the beams, we observe $R_{21} = R_{21,\text{intrinsic}}$. If the source is less extended than the beams, the observed ratio is higher than the intrinsic ratio with the limiting case of a point-like source, for which we would observe $R_{21} = R_{21,\text{intrinsic}} \times \frac{\text{FWHM}(\text{CO}(1-0))^2}{\text{FWHM}(\text{CO}(2-1))^2} = 4 \times R_{21,\text{intrinsic}}$ (see Solomon et al. 1997; Lisenfeld et al. 2019).

Matched aperture observations of CO(1-0) and CO(2-1) give line ratios of $R_{21} = 0.89 \pm 0.6$ (Braine et al. 1993, for a small sample of nearby spiral galaxies), $R_{21} \sim 0.8$ (Leroy et al. 2009, for the SINGS sample), and $R_{21} \sim 0.6-0.9$ (Casasola et al. 2015, for four low-luminosity AGNs from the NUCleI of GALaxies, NUGA, survey). $R_{21} \sim 0.8$ is consistent with optically thick gas with an excitation temperature of ~ 10 K (Leroy et al. 2009). It is therefore very likely that a similar situation holds in our sample, meaning that $R_{21} > 1$ can be interpreted as optically thick thermalised gas with a spatial extension smaller than (at least) the CO(1-0) beam, and the values between $R_{21} \sim 0.5-1$ as optically thick extended gas in thermal and kinetic equilibrium. There are a few pointings with a lower R_{21} which indicate subthermal excitation due to a low gas density.

Maeda et al. (2020) found a tight correlation between the SFE and R_{21} (observed with matched beams, so that $R_{21} = R_{21,\text{intrinsic}}$), in the sense that areas with a low SFE show a lower R_{21} . We cannot confirm this trend. Even though we took into account the different beam sizes of CO(1-0) and CO(2-1), a possibly small source size can only mean that our measured R_{21} overestimates $R_{21,\text{intrinsic}}$. Thus, the low values of R_{21} indicate that $R_{21,\text{intrinsic}}$ is even lower. From Fig. F.1 (lower panel) it is

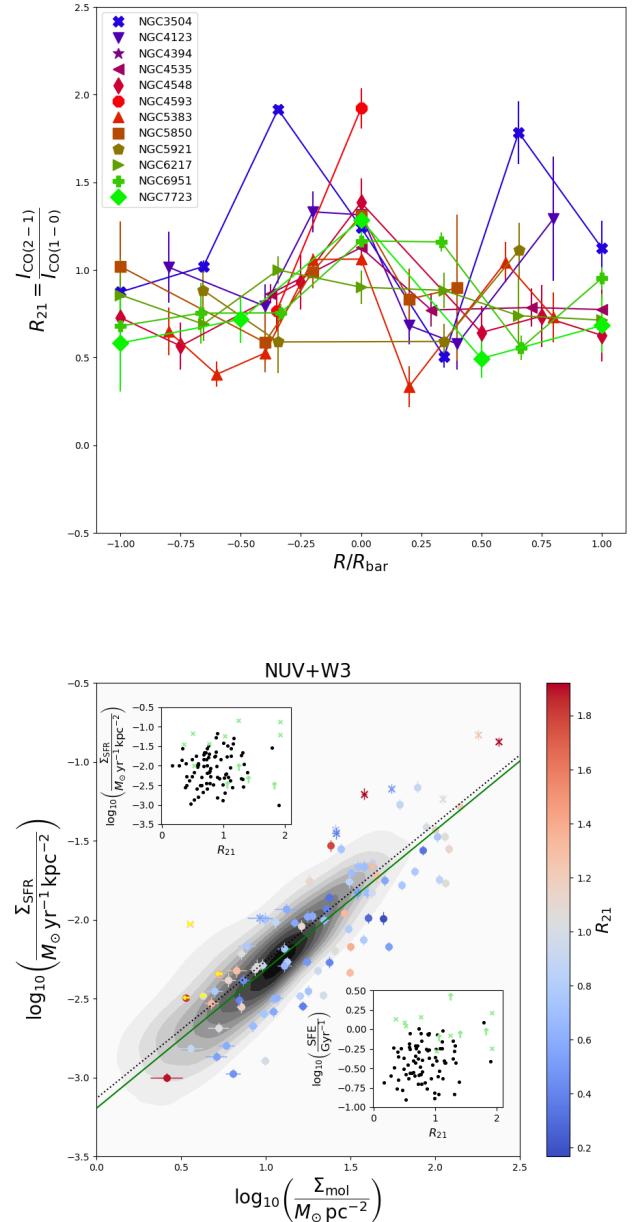


Fig. F.1. Gas excitation across bars. *Upper panel:* Ratio of the velocity-integrated line intensity of CO(2-1) to CO(1-0) spectra along bars of each of the 12 galaxies in our sample (see legend). *Lower panel:* As in Fig. 7, but colour-coded by R_{21} . The insets show Σ_{SFR} (upper left corner) and SFE (lower right corner) vs R_{21} ; the green points indicate non-reliable measurements. The grey contours trace the measurements from Bigiel et al. (2011).

hard to see how we could produce a trend between $R_{21,\text{intrinsic}}$ and the SFE, since the low values ($R_{21} < 0.5$) are rather evenly distributed in pixels of high and low SFEs.